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TOI-1080 b: a temperate, rocky planet orbiting a quiet M4V host

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ABSTRACT

We present the detection and validation of a small, temperate transiting exoplanet orbiting TOI-1080 every $3.9652482^{+0.0000014}_{-0.0000015}$ d. The host is a quiet M4V star at 25.6 pc. The planet signal was first detected by the Transiting Exoplanet Survey Satellite (*TESS)* and validated using *TESS* and ground-based observations. By fitting the available light curves, the planet radius is measured to be 1.200 ± 0.058 $R_\oplus$ and its equilibrium temperature of $368^{+12}_{-10.}$ K. With Near Infra Red Planet Searcher (NIRPS) radial velocities, we are able to place a 3σ upper limit on the mass of TOI-1080 b of 10.7 $M_\oplus$. Our injection-recovery tests enable us to discard additional transiting planets in the TOI-1080 system with radii down to 0.9 $R_\oplus$ and periods between 0.5 and 7.7 d, and planets with radii larger than 1.4 $R_\oplus$ for periods up to 19 d. We demonstrate that it is highly amenable to characterization of its mass and putative atmosphere. In particular, we find that TOI-1080 b is an exceptional target for the ongoing *JWST* + *HST* Rocky Worlds DDT programme, having a priority score that is higher than four out of nine targets currently being investigated by the programme. TOI-1080 b can be added to the sample of nearby benchmark planets accessible for detailed study with *JWST*.

Key words: techniques: spectroscopic – techniques: photometric – exoplanets – stars: individual: TOI-1080 – stars: low-mass – planetary systems.

1 INTRODUCTION

Of the currently known over 6225 exoplanets (J. Schneider et al. 2011, updated on 2026-02-05), ∼75 per cent have been detected via the transit method, but of those, only ∼9 per cent are transiting stars cooler than 4000 K. But because M dwarfs are the most numerous stars in the solar neighbourhood (T. J. Henry & W.-C. Jao 2024; J. D. Kirkpatrick et al. 2024), they could potentially harbour most of the terrestrial planets in the Galaxy (e.g. F. C. Adams, P. Bodenheimer & G. Laughlin 2005; C. D. Dressing & D. Charbonneau 2015; R. Burn et al. 2021; L. Mignon et al. 2025). At first order, given that the transit depth is defined by the relative size between the stellar and planet radius (e.g. S. Seager & G. Mallén-Ornelas 2003), terrestrial planets ($R_{\rm pl} \lesssim 1.6$ $R_\oplus$; e.g. L. A. Rogers 2015; R. Cloutier & K. Menou 2020; H. Parviainen, R. Luque & E. Palle 2024) are more easily detected and studied around these smaller stars. As such, the rocky planets most readily available for characterization now are those orbiting M-dwarf hosts. One of the stated science goals of the *TESS* mission is to identify planets with radius <2.5 $R_\oplus$ around stars in the solar neighbourhood in orbits with periods of up to 10 d (G. R. Ricker et al. 2014). Currently, there are just over 183 of these nearby, small planets (J. L. Christiansen et al. 2025, updated on 2026-02-05).

The growing sample of M-dwarf planets has shown that small planets orbiting M dwarfs can exhibit a wide-range of planet structures (e.g. C. Dorn et al. 2017). Although some rocky planets are expected to have atmospheres (e.g. K. J. Zahnle & D. C. Catling 2017; P. C. August et al. 2025), a recent review by L. Kreidberg &

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Table 1. Stellar parameters adopted for this work.

	TOI-1080, TIC 161032923, 2MASS J18252886-5212514, APASS 25090057, *Gaia* DR2 6653968157663516672, *Gaia* DR3 6653968157663516672, UCAC4 189–186987, *WISE* J182528.54-521252.0, SIPS J1825-5212	
Designations		
Parameter	Value	Source
T mag	12.707 ± 0.007	K. G. Stassun et al. (2019)
B mag	17.138 ± 0.008	N. Zacharias et al. (2013)
V mag	15.349 ± 0.049	N. Zacharias et al. (2013)
G mag	14.0343 ± 0.0006	Gaia Collaboration (2023)
J mag	11.044 ± 0.022	R. M. Cutri et al. (2003)
H mag	10.467 ± 0.023	R. M. Cutri et al. (2003)
K mag	10.17 ± 0.02	R. M. Cutri et al. (2003)
$W1$ mag	10.015 ± 0.023	R. M. Cutri et al. (2021)
$W2$ mag	9.83 ± 0.02	R. M. Cutri et al. (2021)
$W3$ mag	9.584 ± 0.042	R. M. Cutri et al. (2021)
$W4$ mag	> 8.42	R. M. Cutri et al. (2021)
Distance	25.559 ± 0.054 pc	C. A. L. Bailer-Jones et al. (2021)
α	18:25:28.4	Gaia Collaboration (2023)
δ	−52:12:52.46	Gaia Collaboration (2023)
μ_α	-270.293 mas yr^{-1}	Gaia Collaboration (2023)
μ_δ	-69.5518 mas yr^{-1}	Gaia Collaboration (2023)
Parallax	39.132 ± 0.022 mas	Gaia Collaboration (2023)
SpT (Opt + NIR)	M4.0 ± 1.0	This work (Sections 2.1–2.2)
$R_\star$	$0.2019 \pm 0.0075 R_\odot$	This work (Section 2.3)
$M_\star$	$0.1667 \pm 0.0041 M_\odot$	This work (Section 2.3)
$L_\star$	$0.00324 \pm 0.00046 L_\odot$	This work
$T_{\rm eff}$	3065 ± 50 K	This work
$\log g_\star$	5.049 ± 0.043	This work
[Fe/H]	-0.25 ± 0.25 dex	This work (SED)
[Fe/H]	-0.09 ± 0.11 dex	This work (near-infrared)
[Fe/H]	-0.10 ± 0.20 dex	This work (optical)
RV	-78.3 ± 2.1 kms^{-1}	Gaia Collaboration (2022)
Age	$\gtrsim 5-7$ Gyr	This work

K. B. Stevenson (2025) highlights that there has been no unambiguous detection of an atmosphere of a rocky planet with the state-of-the-art *James Webb Space Telescope* (*JWST*) observations. The detections of a volatile-atmosphere around the hot Super-Earth 55 Cnc e (R. Hu et al. 2024), a tenuous rock vapour atmosphere around K2-141 b (S. Zieba et al. 2022), and a thick volatile-atmosphere around TOI-561 b (J. K. Teske et al. 2025), all orbit more massive stars. For example, the Near Infrared Imager and Slitless Spectrograph (NIRISS) transits of Super-Earth GJ 357 b show no atmospheric features (J. Taylor et al. 2025) and the Near Infrared Spectrograph (NIRSpec) transits of GJ 1156 b also show a flat transmission spectrum (K. A. Bennett et al. 2025). For TRAnsiting Planets and PlanetesImals Small Telescope (TRAPPIST-1), thermal emission photometry with Mid-Infrared Instrument (MIRI) discarded dense compact atmospheres around the two inner planets (M. Gillon et al. 2025), while transit spectroscopy with NIRISS and/or NIRSpec has failed to bring strong constraints on the presence/absence atmospheres of the five outer planets so far, because of limited signal-to-noise ratios (SNRs) and the negative impact of stellar noises (e.g. A. Glidden et al. 2025). The Rocky Worlds Discretionary Director's Time programme has been allocated 500 h with *JWST* and 250 orbits with *Hubble Space Telescope* (*HST*) to expand our understanding of atmospheres of terrestrial planets orbiting M dwarfs (S. Redfield et al. 2024). The first results from this programme have found that for two secondary eclipses of GJ 3929 b observed with MIRI at 15μm, the most likely interpretation is a bare rock (Q. Xue et al. 2025).

One of the significant hurdles to a robust atmospheric detection for these rocky M-dwarf planets is the stellar contamination due to their active planet hosts. When the stellar photosphere has inhomogeneities that are not occulted during transit, it leads to biases in the transmission spectra known as the transit light source effect (e.g. B. V. Rackham, D. Apai & M. S. Giampapa 2018). Although strategies that rely on stellar models are currently applied, more development of stellar models are needed for robust mitigation (B. V. Rackham & J. de Wit 2024). The atmospheric characterization of rocky planets can inform our understanding of their interior structure (e.g. T. Lichtenberg et al. 2025), internal sources of energy (e.g. Z. Lin & T. Daylan 2026) and give insights into the planet formation history (e.g. K. I. Öberg, S. Facchini & D. E. Anderson 2023). As such, it has become important to detect and characterize small planets around small, inactive stars.

In this paper, we present the detection and validation of the planetary nature of TOI-1080 b, a temperate, rocky planet orbiting an inactive M4V star. We present first the characterization of the star TOI-1080 in Section 2 based on medium resolution spectroscopy and spectral energy distribution fitting. We assess possible blends with archival images and high-resolution imaging, and discuss stellar activity. In Section 3, we present the transit photometry obtained with *TESS* and the follow-up, time-series photometry obtained with ground-based facilities. In Section 4, we statistically evaluate the planet scenario against other astrophysical scenarios and validate the planet nature of TOI-1080 b. In Section 5, we present the modelling of the transit light curves to characterize the physical properties of the planet TOI-1080 b. The upper limit on the planet mass based on optical and near-infrared precision radial velocities is given in Section 6. We discuss the significance of this detection and the potential for further characterization of the planet in Section 7.

2 STELLAR CHARACTERIZATION OF TOI-1080

As the physical properties of the planetary system depend on the stellar properties, here we present the careful stellar characterization of the host TOI-1080 based on optical (Section 2.1) and near-infrared (Section 2.2) medium-resolution spectra, and its spectral energy distribution (Section 2.3). The star's coordinates, magnitudes, and properties are summarized in Table 1. We note that although TOI-1080 was observed with high-resolution spectrographs (see Section 6), these observations do not have a sufficiently high SNR to derive accurate atmospheric parameters for the host star and are used only for radial velocities.

2.1 Spectral type from near-infrared FIRE spectrum

We observed TOI-1080 with the Folded port InfraRed Echellette (FIRE) spectrograph (R. A. Simcoe et al. 2008) on the 6.5-m Magellan Baade Telescope on UT 2021-10-15 under clear conditions with 0.7 arcsec seeing. The observations were taken in the high-resolution echellette mode using the 0.6 arcsec slit, providing wavelength coverage of 0.82–2.51 μm at a resolving power of $R \approx 6000$. We gathered four exposures of 200.8 s at an average airmass of 1.2. For flux and telluric calibration, we observed the A0 V star HD 163761 ($V = 6.7$) in two 5-s exposures at a similar airmass. ThAr arc lamp frames were collected at each pointing, along with daytime internal quartz and dome flats for pixel and

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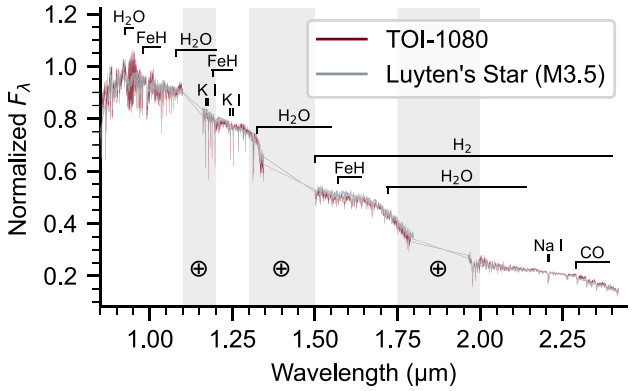


Figure 1. FIRE spectrum of TOI-1080. The target spectrum (red) is shown alongside that of the SPEX SXD spectrum of the M3.5 V standard Luyten's Star (GJ 273; grey). The higher spectral resolution of the FIRE spectrum gives it a more jagged appearance. Strong M-dwarf spectral features and spectral regions with strong tellurics are indicated. The figure shows the normalized flux of the planet host TOI-1080 as a function of wavelength, between 0.9 to 2.35 microns.

blaze corrections. The data were reduced with the FIREHOSE pipeline.[1] The resulting spectrum has a median SNR per resolution element of 360.

The FIRE spectrum of TOI-1080 is shown in Fig. 1. Comparison with single-star standards in the IRTF Spectral Library (M. C. Cushing, J. T. Rayner & W. D. Vacca 2005; J. T. Rayner, M. C. Cushing & W. D. Vacca 2009) using the SpeX Prism Library Analysis Toolkit (SPLAT; A. J. Burgasser & Splat Development Team 2017) indicates the closest match to the M3.5 V standard Luyten's Star (GJ 273), and we adopt a near-infrared spectral type of M3.5 ±0.5 accordingly. From the *K*-band Na I and Ca I features and the H_2O–K2 index (B. Rojas-Ayala et al. 2012), we estimate $[Fe/H] = -0.09 \pm 0.11$ using the A. W. Mann et al. (2013) relation. Furthermore, contrary to what is observed in giant stars, a visual inspection of the spectra of TOI-1080 reveals clear absorption by both the Na I doublet at ∼2.21 μm and weak CO bands near 2.3 μm. This provides further support on the cool main-sequence nature of TOI-1080.

2.2 Spectral type from optical MagE spectrum

We observed TOI-1080 on UT 2023-05-22 with the Magellan Echellette spectrograph (MagE; J. L. Marshall et al. 2008) on the 6.5-m Magellan Baade Telescope. The night was clear, with seeing of 1.3 arcsec. We used the 0.7 arcsec × 10 arcsec slit, which provides a resolving power of $\lambda/\Delta\lambda \approx 6000$ over the 4000–9000 Å range. Two 200-s exposures were obtained at an airmass of 1.1. For flux calibration, we observed the spectrophotometric standard EG 274 (M. Hamuy et al. 1992, 1994) in two 10-s integrations. Wavelength calibration and flat-field corrections were performed using ThAr arc frames at each pointing and the standard set of afternoon calibrations (biases, Xe flats, and internal flats). Data reduction followed standard procedures using PYPEIT (J. Prochaska et al. 2020; J. X. Prochaska et al. 2020). The final spectrum has an SNR of ∼100 per pixel at 8000 Å.

The reduced spectrum is shown in Fig. 2, and displays features typical of mid-type M dwarfs, with a red spectral slope; strong

[1] https://github.com/rasimcoe/FIREHOSE

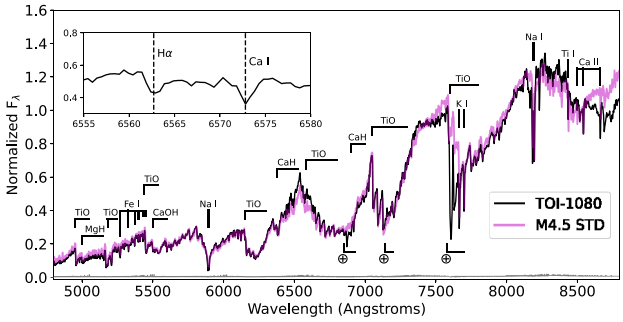


Figure 2. MagE optical spectrum of TOI-1080 (black line) normalized at 7500 Å compared to a combined M4 + M5 optical spectral template from J. J. Bochanski et al. (2007, magneta line). Standard M dwarf atomic and molecular features are labelled, as well as regions of uncorrected telluric absorption (⊕). The inset box shows the 6555–6580 Å region highlight absorption from Hα and Ca I.

TiO and CaH molecular features; and line absorption from K I, Na I, Ca I, Ca II, Ti I, and Fe I. The spectrum appears to be intermediate between the M4 and M5 dwarf spectral templates from J. J. Bochanski et al. (2007), as confirmed by index-based classifications from I. N. Reid, S. L. Hawley & J. E. Gizis (1995), J. E. Gizis (1997), S. Lépine, R. M. Rich & M. M. Shara (2003), and F. C. Riddick, P. F. Roche & P. W. Lucas (2007) which span M4–M4.5. We adopt an optical classification of M4.5 ± 0.5, slightly later but formally consistent with the near-infrared classification. We detect the 6563 Å H*α* feature in absorption with an equivalent width EW = 0.33 ± 0.02 Å, indicating weak or absent magnetic emission and hence an activity age ≳5–7 Gyr (A. A. West et al. 2008). The metallicity index $\zeta = 0.909 \pm 0.002$ (S. Lépine, R. M. Rich & M. M. Shara 2007; S. Lépine et al. 2013) indicates a near-solar metallicity of [Fe/H] = −0.1±0.2 based on the calibration of A. W. Mann et al. (2013), similar to that inferred from the near-infrared spectrum.

2.3 Spectral energy distribution for TOI-1080

As an independent determination of the basic stellar parameters, we performed an analysis of the broad-band spectral energy distribution (SED) of the star together with the *Gaia* DR3 parallax (with no systematic offset applied; see e.g. K. G. Stassun & G. Torres 2021), in order to determine an empirical measurement of the stellar radius, following the procedures described in K. G. Stassun & G. Torres (2016), K. G. Stassun, K. A. Collins & B. S. Gaudi (2017), and K. G. Stassun et al. (2018). We pulled the JHK_S magnitudes from 2MASS, the $G_{BP}G_{RP}$ magnitudes from *Gaia*, the *gri* magnitudes from the AAVSO Photometric All-Sky Survey, and the *W*1–*W*3 magnitudes from *WISE*. Together, the available photometry spans the full stellar SED over the wavelength range 0.4–10 μm (see Fig. 3).

We performed a fit using PHOENIX stellar atmosphere models (T. O. Husser et al. 2013), with the free parameters being the effective temperature ($T_{\rm eff}$) and metallicity ([Fe/H]). The extinction, A_V, was fixed at zero due to the proximity of the system. The resulting fit (Fig. 3) has a reduced χ^2 of 1.6, with a best-fit $T_{\rm eff} = 3065 \pm 50$ K, and $[Fe/H] = -0.25 \pm 0.25$. Integrating the model SED gives the bolometric flux at Earth, $F_{\rm bol} = 1.589 \pm 0.056 \times 10^{-10}$ erg s^{-1} cm^{-2}. Taking the $F_{\rm bol}$ together with the *Gaia* parallax directly yields the stellar radius via the Stefan–Boltzmann relation, giving $R_\star = 0.2019 \pm 0.0075$ R$_\odot$. In addition,

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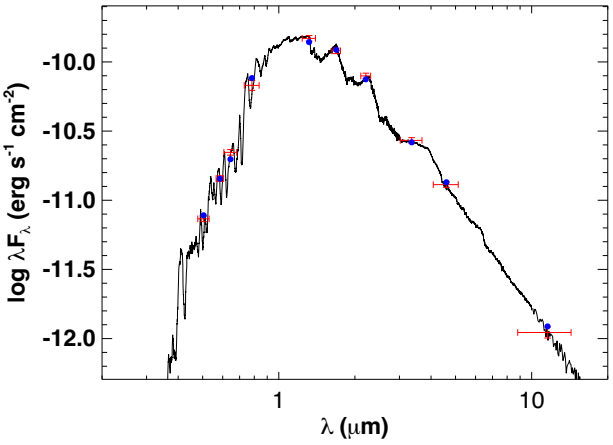


Figure 3. Spectral energy distribution of TOI-1080. The symbols with error bars represent the observed photometric measurements, where the horizontal bars represent the effective width of the passband. The square symbols are the model fluxes from the best-fitting PHOENIX atmosphere model (continuous line).

we can estimate the stellar mass from the empirical relations of A. W. Mann et al. (2019), giving $M_\star = 0.1667 \pm 0.0041$ M$_\odot$.

2.4 Search for blends

2.4.1 SOAR speckle imaging

High-angular resolution imaging is needed to search for nearby sources that can contaminate the *TESS* photometry, resulting in an underestimated planetary radius, or be the source of astrophysical false positives, such as background eclipsing binaries. We searched for stellar companions to TOI-1080 with speckle imaging on the 4.1-m Southern Astrophysical Research (SOAR) telescope (A. Tokovinin 2018) on UT 2021-07-14, observing in Cousins *I* band, a similar visible bandpass as *TESS*. This observation was sensitive with 5σ detection to a 5.0-mag fainter star at an angular distance of 1 arcsec from the target. More details of the observations within the SOAR *TESS* survey are available in C. Ziegler et al. (2020). The 5σ detection sensitivity and speckle autocorrelation functions from the observations are shown in Fig. 4. No nearby stars were firmly detected within 3 arcsec of TOI-1080 in the SOAR observations.

2.4.2 Archival images of TOI-1080

Fig.5 shows the archival images for the TOI-1080 field. We have centred TOI-1080 in the SPECULOOS-Southern Observatory (SSO) Io image taken in 2021 and note that the position is not significantly different in the SSO Europa image taken in 2025. Over the span of over 45 yr that the archival images encompass, there is no other star in the position of the SSO Io observations of TOI-1080. Thus, it is unlikely that there is contamination by background source that is bright enough to affect our photometry and results.

2.5 Stellar activity

We analysed the *TESS* and SPECULOOS light curves, described in the next section, of TOI-1080 to search for evidence of stellar rotational modulation or flaring. We could not identify any flares

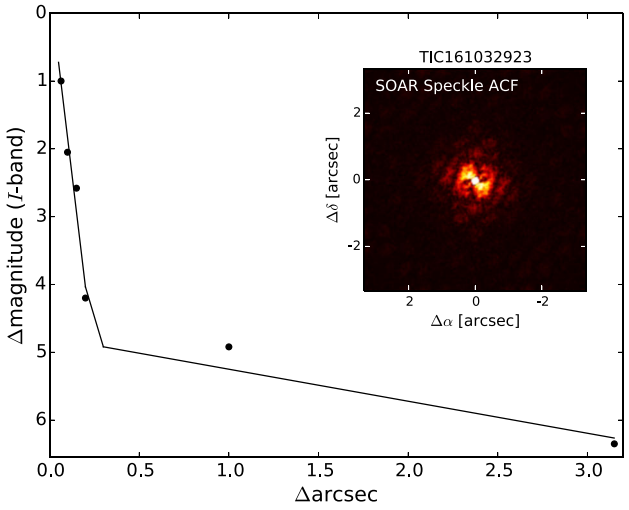


Figure 4. Detection sensitivity and autocorrelation (inset) functions for the TOI-1080 SOAR speckle imaging. We do not identify any nearby stars within 3 arcsec of TOI-1080.

in the light curves or any rotationally induced photometric variability. A frequency analysis of the *TESS* photometry (i.e. both PDCSAP and SAP fluxes), with in-transit data masked, was performed using a generalized Lomb–Scargle periodogram (GLS; M. Zechmeister & M. Kürster 2009) to search for the stellar rotation period. When analysed independently, each of the three sectors yields a different period and the phased light curves do not show clear rotational modulation, suggesting that the signals are not significant. The *TESS* SAP light curves present some variability that if it is due to rotation, it has a much longer time-scale than any given sector. Because the sectors are too far apart, the modulation is not coherent on any particular period. We could not derive a robust rotation period from the *TESS* data set, indicating that any spot modulation of the light curves has a small amplitude and/or a long rotation period below our detection sensitivity. We further analysed the ground-based photometry from SPECULOOS, which comprises 243 h of observations spanning ∼1364 d. The resulting periodogram is noisy, with the amplitude at the period of maximum power remaining below 0.4 ppt. The low amplitude of photometric variability and absence of flares in both datasets indicate that TOI-1080 is a quiet star.

Unfortunately, the available high-resolution spectra of TOI-1080 do not have sufficient signal in the Ca II H&K lines (SNR <1) nor in Hα (SNR ∼5) to assess chromospheric activity via spectroscopic indicators.

3 TRANSIT PHOTOMETRY: OBSERVATIONS AND DATA REDUCTION

In this section, we describe the observations and data reduction used in the identification of the transit events in the *TESS* light curves and the characterization of the transiting planet TOI-1080 b with additional ground-based observations.

3.1 *TESS* light curve and planet candidate status

TOI-1080 was observed by *TESS* in Sector 13 (2019-04-22 to 2019-05-21), 66 (2023-06-02 to 2023-07-01), and 93 (2025-06-03 to 2025-06-29) with a 2-min cadence. The Science Processing Operations

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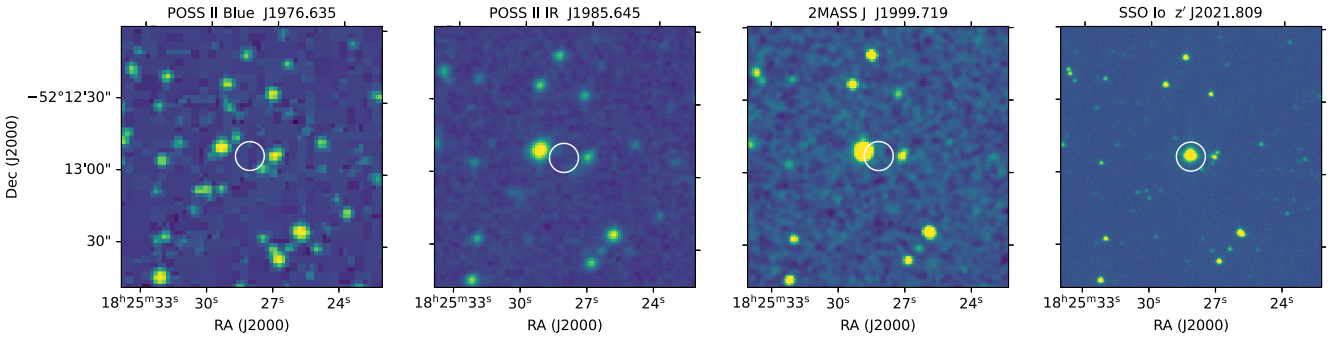


Figure 5. Archival images for TOI-1080 from the Digital Sky Survey (DSS) and 2MASS compared to the SSO Io observations taken in 2021. The position of TOI-1080 in the 2021 SSO Io stack image is marked in all panels by the circle and the coordinates are given in J2000. All images are the same size of $0.03^\circ \times 0.03^\circ$, and are centred on TOI-1080's position in the 2021 SSO Io observations. The colour scale for each frame has been optimized to better show the stars.

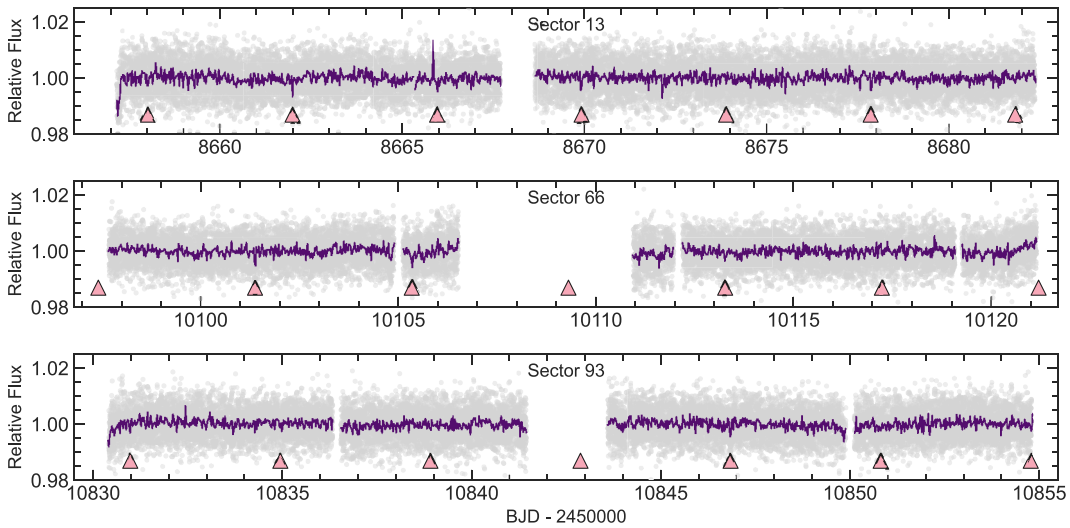


Figure 6. Normalized *TESS* light curve. The 45 967 individual photometric points are shown as individual points (in grey). The continuous line (in purple) shows the 30 min binned flux. The transits of TOI-1080 b are marked by the pink triangles.

Center (SPOC; J. M. Jenkins et al. 2016) NASA Ames Research Center conducted a transit search in Sector 13 using a wavelet-based, noise-compensating matched filter (J. M. Jenkins 2002; J. M. Jenkins et al. 2010, 2020) and detected a threshold crossing event. The signature passed all the tests in the Data Validation (DV) report (J. D. Twicken et al. 2018; J. Li et al. 2019), including the difference image centroid test, which constrained the location of the transit source to within 4.3 ± 5.0 arcsec of the catalogue location of TOI-1080. The *TESS* Science Office reviewed the data validation reports and issued an alert for TOI-1080 b on UT 2019 August 16 (N. M. Guerrero et al. 2021). The transit signature was detected in a subsequent search of sectors 13 and 66 where the difference image centroid test constrained the transit source to within 5.3 ± 3.0 arcsec. The three-sector *TESS* light curve is composed of 45 967 individual photometric points and it spans ∼2198d.

SPOC data products are publicly available as simple aperture photometry (SAP; J. D. Twicken et al. 2010; R. L. Morris et al. 2020) or presearch data conditioning simple aperture photometry (PDCSAP; J. C. Smith et al. 2012; M. C. Stumpe et al. 2012, 2014); the latter is corrected for instrumental systematics and contamination of known nearby sources. We downloaded the PDCSAP lightcurve for TOI-1080 from the NASA Milkulski Archive for Space Telescopes via LIGHTKURVE (Lightkurve Collaboration 2018). The *TESS* light curve for all three sectors is presented in Fig. 6.

With the python package TPFPLOTTER (A. Aller et al. 2020), we plotted the field of view and aperture of the three *TESS* sectors to assess contamination from nearby stars using the *Gaia* DR3 catalogue (Gaia Collaboration 2023) and to discard that the source of the transit event itself in the *TESS* light curve is another star in the field. Fig. 7 shows the target pixel files for each *TESS* sector from which we can conclude the dilution is minimal in the *TESS* light curves from nearby stars.

3.2 Follow-up, ground-based light curves

To better constrain the planet radius and ephemeris, we obtained ground-based, time-series photometry for the transit of TOI-1080 b. A summary of the follow-up, time-series observations is presented in Table 2.

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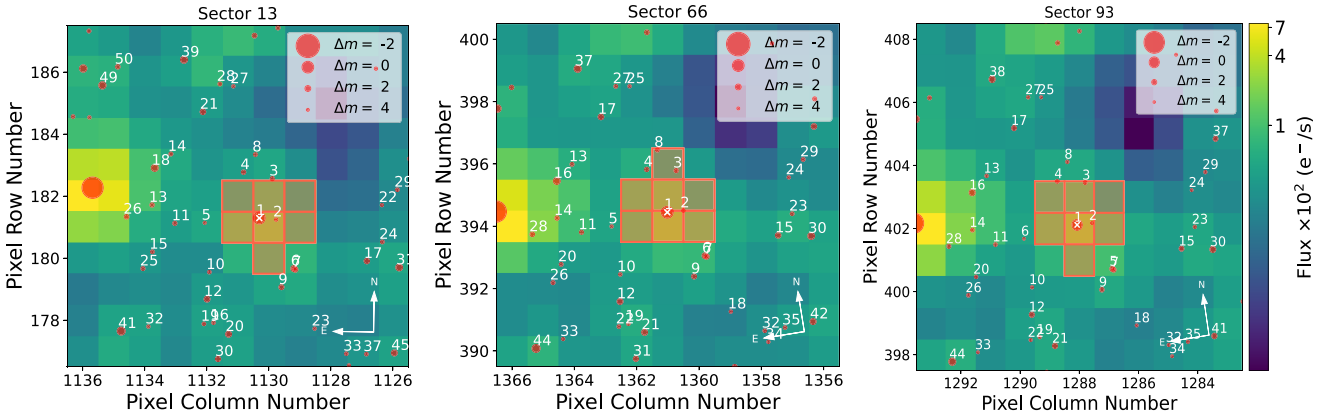


Figure 7. *TESS* Target pixel files and apertures (shaded). Although, the TOI-1080 field is relatively crowded, the brightest star is not affecting the *TESS* photometry and nearest stars are >4 mag fainter than the target in the *TESS* passband.

Table 2. Ground-based time-series photometric follow-up observations of TOI-1080. The $\Delta \ln Z$ column is the Bayesian evidence for the transit, as described in Section 5. Only light curves with $\Delta \ln Z > 5$ are used in the global analysis. *a*: The ExTrA light curves were treated as a single long time series for each telescope as described in Section 5; we therefore do not report a Bayes Factor for each individual light curve.

Date (UT)	Facility	Filter	$t_{\rm exp}(s)$	$\Delta \ln Z$	Notes
2019-09-02	LCO-SAAO−1.0m	*Sloan*−i'	100	8.32	full
2020-05-09	LCO-SSO−1.0m	Pan-STARRS z_s	150	0.79	full
2020-05-21	PEST	$R_{\rm C}$	60	−3.04	partial
2020-08-01	LCO-SAAO−1.0m	*Sloan*−r'	150	6.48	full
2020-08-16	LCO-SAAO−1.0m	Pan-STARRS z_s	150	11.89	full
2020-08-16	LCO-SAAO−1.0m	V	150	0.72	full
2020-08-24	LCO-SAAO−1.0m	Pan-STARRS z_s	150	3.04	full
2021-04-19	LCO-SSO−1.0m	*Sloan*−i'	100	229.59	full
2021-05-24	ExTrA 2 & 3[a]	1.21 µm	60	–	full
2021-06-01	ExTrA 2 & 3	1.21 µm	60	–	full
2021-06-05	TRAPPIST-South	$I+z$	65	20.78	full
2021-06-05	ExTrA 2 & 3	1.21 µm	60	–	full
2021-06-09	ExTrA 2 & 3	1.21 µm	60	–	full
2021-10-18	ExTrA 2 & 3	1.21 µm	60	–	full
2021-10-22	SPECULOOS Io	*Sloan*−z'	10	0.42	full
2021-10-22	ExTrA 2 & 3	1.21 µm	60	–	full
2021-10-26	TRAPPIST-South	$I+z$	65	−1.88	partial
2022-05-08	ExTrA 2 & 3	1.21 µm	60	–	full
2022-05-12	ExTrA 2 & 3	1.21 µm	60	–	full
2023-04-14	SPECULOOS Io	*Sloan*−z'	10	25.81	full
2023-08-23	ExTrA 2 & 3	1.21 µm	60	–	full
2023-08-27	ExTrA 2	1.21 µm	60	–	full
2024-08-14	ExTrA 2	1.21 µm	60	–	full
2025-06-27	SPECULOOS Europa	*Sloan*–z'	10	31.046	full

3.2.1 SPECULOOS photometry

We made use of two of the four telescopes of the SSO (L. Delrez et al. 2018; M. Gillon 2018; D. Sebastian et al. 2021) to obtain transit light curves of TOI-1080 b. SSO is located at ESO Paranal Observatory, and is made up of four identical 1 m-class telescopes named after the Galilean moons (Io, Europa, Ganymede, and Callisto). They are equipped with a 2K × 2K Andor iKon-L cameras covering a field of view of 12 arcmin × 12 arcmin for a pixel scale of 0.35 arcsec. We obtained three full transits of TOI-1080.01: two with the Io telescope on UT 2021-10-22 and UT 2023-04-14, and one the Europa telescope on UT 2025-06-28 in the *Sloan-z'* filter and with 10 s exposures. We reduced the data and performed differential photometry using a custom pipeline built with the PROSE package (L. J. Garcia et al. 2022).

In addition to transit observations, we obtained 27 nights of continuous monitoring of the star between UT 2023-06-26 and UT 2023-07-26. The goal was to search for hints of other possible candidates in the system, and assess our performance and the detectability of such candidates with the SPECULOOS telescopes. These observations were carried out in the *Sloan-z'* filter with exposures of 10 s, and they make a total of 243 h.

3.2.2 LCO-SAAO and LCO-SSO photometry

We used the Las Cumbres Observatory Global Telescope (LCOGT; T. M. Brown et al. 2013) 1.0 m network to observe a total of eight full transits of TOI-1080 b. Seven transits were observed

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with the LCOGT 1.0 telescope at the South Africa Astronomical Observatory (LCO-SAAO) and one was observed with the LCOGT 1.0 telescope at Siding Spring Observatory (LCO-SSO) in Australia. The telescopes are equipped with a 4096 × 4096 SINISTRO camera having an image scale of 0.389 arcsec per pixel and a field of view of 26 arcmin × 26 arcmin.

The observations were carried out in the Sloan-r', Sloan-i', Johnson-V, and Pan-STARRS z_s filters with exposure times of 150, 100, 150, and 150s, respectively. The observation dates are given in Table 2. The data calibration was performed using the standard LCOGT BANZAI pipeline (C. McCully et al. 2018). The photometry extraction was performed in an uncontaminated 3.1–5.2 arcmin target apertures using ASTROIMAGEJ (K. A. Collins et al. 2017).

3.2.3 *TRAPPIST-South photometry*

TRAPPIST-South (TRAnsiting Planets and PlanetesImals Small Telescope; M. Gillon et al. 2011; E. Jehin et al. 2011) is a 60 cm Ritchey–Chrétien robotic telescope located at La Silla Observatory in Chile since 2010. The telescope is equipped with a 2K×2K FLI Proline CCD camera with a FOV of 22 arcmin × 22 arcmin and a pixel scale of 0.65 arcsec pixel^{-1}. Two full transits were observed on UT 2021-06-05 and UT 2021-10-26 in the $I + z$ filter with exposure time of 65 s. The data reduction and photometric extraction were performed using the PROSE pipeline (L. J. Garcia et al. 2022).

3.2.4 *ExTrA photometry*

ExTrA is an instrument composed of three 60-cm telescopes located at La Silla Observatory in Chile that is capable of obtaining near-infrared photometry (X. Bonfils et al. 2015). We have observed TOI-1080 during 12 nights using one or two of the ExTrA telescopes between UT 2021-05-24 and UT 2024-08-14. We used 8 arcsec aperture fibers and the lowest-resolution mode (R∼20) of the spectrograph with an exposure time of 60 s. The details of the observations are found in Table 2. To process the data, we followed the procedure described in detail in M. Cointepas et al. (2021) to subtract the sky flux and normalize by the flat, and obtained the differential photometry for TOI-1080 using the comparison stars.

3.2.5 *PEST photometry*

The Perth Exoplanet Survey Telescope (PEST) has a 0.3-m diameter telescope and is located near Perth, Australia. PEST observed the partial transit of TOI-1080 b on UT 2020-05-21 with 60 s exposures in the R_C. A total of 315 images over 401 min were obtained with a 1530 × 1020 SBIG ST-8XME camera with an image scale of 1.2 arcsec pixel^{-1} resulting in a 31 arcmin × 21 arcmin field of view. A custom pipeline based on C-MUNIPACK[2] was used to calibrate the images and extract the differential photometry.

4 PLANET VALIDATION

We made use of the statistical validation software TRICERATOPS (S. Giacalone & C. D. Dressing 2020; S. Giacalone et al. 2021) to test the probability of the transiting planet hypothesis for the TOI-1080.01 transit signal. TRICERATOPS calculates the flux contribution from nearby stars to assess whether they could be responsible for the transit signal. It then calculates probabilities for a variety of transiting planet and eclipsing binary scenarios on the target and nearby stars by fitting lightcurve models, and finally provides a false positive probability (FPP). The typical threshold for statistical validation of a planet is an FPP of 0.015.

In the first instance, we calculated the FPP using the phase-folded *TESS* photometry, including all three sectors, and the contrast curve from our high-resolution imaging described in Section 2.4.1; the resulting value for the FPP is 0.01626 which is just above the threshold.

We then used TRICERATOPS' functionality that allows a user to discard certain false positive scenarios if ground-based observations have already ruled them out. We dropped NTP (nearby transiting planet), NEB (nearby eclipsing binary), and NEB × 2P (nearby eclipsing binary at twice the period) for 13 nearby stars, as the small pixel scale of our ground-based photometry conclusively rules out these stars are the source of the transit event. This reduced the FPP to 0.001948, placing it confidently below the threshold for statistical validation.

We therefore consider the planet to be validated and refer to it as TOI-1080 b.

5 LIGHT CURVE MODELLING

We carried out a global analysis of the complete photometric data set described in Section 3 using ALLESFITTER (M. N. Günther & T. Daylan 2019, 2021), a flexible PYTHON inference package. ALLESFITTER generates light-curve models using ELLC (P. F. L. Maxted 2016), and best-fitting models can be chosen using either nested sampling or Markov chain Monte Carlo algorithms, implemented via DYNESTY (J. S. Speagle 2020) and EMCEE (D. Foreman-Mackey et al. 2013), respectively. For our analysis we used the nested sampling algorithm as this allows effective model comparison by calculating the Bayes Factor (R. E. Kass & A. E. Raftery 1995), enabling us to choose the model with the highest statistical evidence.

We adopted the signal parameters reported by SPOC[3] as uniform priors, and the stellar parameters adopted in Section 2 as normal priors. Additionally, we used PYLDTK (H. Parviainen & S. Aigrain 2015) and the Phoenix stellar atmosphere models (T. O. Husser et al.) to calculate quadratic limb-darkening coefficient priors for each photometric band used in our follow-up campaign. We reparametrized these in the parametrization of D. M. Kipping (2013) and adopted them as normal priors, following e.g. G. Dransfield et al. (2022, 2024), A. H. M. J. Triaud et al. (2023), and M. G. Scott et al. (2025).

Several of the transits described in Section 3.2 are dominated by red noise. To assess which light curves could contribute to a refinement of the transit parameters, we follow the procedure thoroughly described in F. J. Pozuelos et al. (2023) and subsequently applied in S. Zúñiga-Fernández et al. (2025). In brief, we fit two models to each individual ground-based light curve; the first is a pure noise model with fixed transit parameters, including $R_p/R_\star = 0$, and the second is a noise + transit model. For

[2] http://c-munipack.sourceforge.net

[3] Available at https://exofop.ipac.caltech.edu/tess/target.php?id = 161032923

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the noise model we use ALLESFITTER's hybrid spline functionality. We then compare the Bayesian evidence between the pure noise and the transit + noise models, and select only those with $\Delta \ln Z > 5$ to include in the final model from which the planet's physical properties are derived. The derived value of $\Delta \ln Z$ for each lightcurve is presented in the penultimate column of Table 2.

In the case of the light curves obtained with ExTrA, we stitched together the light curves from each of the two telescopes used. In both cases, the transit + noise model was significantly preferred; however, the photometry has very large systematics and are difficult to detrend using the hybrid spline. We therefore use a Gaussian Process (GP) kernel to model the red noise only in the case of the ExTrA light curves. We select the simple harmonic oscillator (SHO) kernel as implemented by CELERITE (D. Foreman-Mackey et al. 2017). When detrending with a GP, the main challenge is to avoid overfitting; to avoid this we place a uniform prior on the GP frequency (ω_0) such that only low frequency noise is removed. We also remove three individual transits which have high amplitude trends that cannot be removed with our frequency upper limit.

We then jointly fit the *TESS* photometry and the highest quality ground-based light curves, fitting in the first instance a circular model and one with free eccentricity. In both cases, we fit for all transit parameters [$R_p/R_\star$, $(R_\star + R_p)/a$, cos i, T_0, and P]. In the free eccentricity model, we parametrize the eccentricity and argument of periastron as $\sqrt{e_b}\cos\omega_b$ and $\sqrt{e_b}\sin\omega_b$.

The eccentricity derived from the more complex model is $e_b = 0.133^{+0.23}_{-0.095}$; this value is consistent with zero at the 2σ level. This small eccentricity has a 37.5 per cent probability of being spurious, as defined by L. B. Lucy & M. A. Sweeney (1971). Additionally, we compare the statistical evidence for the circular and free eccentricity models by calculating the log Bayes factor, and find that the circular model is very marginally preferred ($\Delta \ln Z \approx 0.3$). We therefore adopt the circular model and we report all fitted and derived parameters in Table 3. In Fig. 8, we present all fitted transit lightcurves along with best-fitting models.

6 UPPER LIMIT ON THE PLANET MASS

We obtained six spectra of TOI-1080 with NIRPS (F. Bouchy et al. 2025) spread over three nights: UT 2023-04-14, 2023-04-17, and 2023-04-19. The observations were taken under the NIRPS Guaranteed Time Observation (GTO) programme (ID: 111.254T.001), with the initial objective of constraining the mass of TOI-1080 b. The target was shortly thereafter deemed too challenging, given the small expected semi-amplitude of ∼2–3 m s^{-1} and the photon-noise uncertainties of NIRPS RVs for such a relatively faint target, which was still being established at the start of the GTO in April 2023. Here, we use these spectra as a reconnaissance radial velocity (RV) of the star to put an upper limit on the mass of the transiting object. These RV data discard the scenario that the transit signal is caused by stellar or brown-dwarf companion orbiting the host.

NIRPS is a high-precision, fiber-fed near-infrared spectrograph (970–1900 nm) mounted on the European Southern Observatory (ESO) 3.6-m telescope at La Silla Observatory and capable of operating simultaneously with the High Accuracy Radial velocity Planet Searcher (HARPS; M. Mayor et al. 2003). Each night, we obtained two consecutive 900 s NIRPS exposures, while HARPS collected a single 1800 s integration. Both instruments were set to their high-efficiency modes to maximize throughput, the High-Efficiency mode for NIRPS ($R \approx 75\,000$, 0.9 arcsec fiber) and the EGGS mode for HARPS ($R \approx 80\,000$, 1.4 arcsec fiber). The NIRPS

Table 3. Adopted priors, and fitted and derived parameters from our fit. Uniform priors are indicated as $\mathcal{U}$(lower bound, upper bound), and normal priors are indicated as $\mathcal{N}(\mu, \sigma)$. *a*: ALLESFITTER shifts the epoch to the centre of the data.

Priors and fitted parameters		
$R_b/R_\star$	$\mathcal{U}(0.04, 0.07)$	0.0548 ± 0.0012
$(R_\star + R_b)/a_b$	$\mathcal{U}(0.02, 0.07)$	$0.0360^{+0.0020}_{-0.0016}$
cos i_b	$\mathcal{U}(0.0, 0.03)$	$0.0129^{+0.0047}_{-0.0058}$
$T_{0;b}$ [a], (BJD)	$\mathcal{U}(2458653.5, 2458654.5)$	$2459756.40002^{+0.00029}_{-0.00032}$
P_b, (d)	$\mathcal{U}(3.8, 4.2)$	$3.9652482^{+0.0000014}_{-0.0000015}$
$q_{1;TESS}$	$\mathcal{N}(0.327, 0.05)$	0.328 ± 0.048
$q_{2;TESS}$	$\mathcal{N}(0.290, 0.05)$	0.290 ± 0.049
$q_{1;\mathrm{Sloan-z'}}$	$\mathcal{N}(0.185, 0.05)$	0.188 ± 0.046
$q_{2;\mathrm{Sloan-z'}}$	$\mathcal{N}(0.281, 0.05)$	0.282 ± 0.049
$q_{1;\mathrm{Sloan-i'}}$	$\mathcal{N}(0.391, 0.05)$	0.415 ± 0.048
$q_{2;\mathrm{Sloan-i'}}$	$\mathcal{N}(0.302, 0.05)$	0.317 ± 0.048
$q_{1;\mathrm{Sloan-r'}}$	$\mathcal{N}(0.688, 0.05)$	0.690 ± 0.049
$q_{2;\mathrm{Sloan-r'}}$	$\mathcal{N}(0.368, 0.05)$	0.371 ± 0.048
$q_{1;\mathrm{I+z}}$	$\mathcal{N}(0.229, 0.05)$	0.216 ± 0.048
$q_{2;\mathrm{I+z}}$	$\mathcal{N}(0.290, 0.05)$	0.292 ± 0.049
$q_{1;1.21\mu\mathrm{m}}$	$\mathcal{N}(0.185, 0.05)$	0.182 ± 0.045
$q_{2;1.21\mu\mathrm{m}}$	$\mathcal{N}(0.281, 0.05)$	0.282 ± 0.045
gp ln S_0(ExTrA2)	$\mathcal{U}(-12, 3)$	$1.79^{+0.34}_{-0.30}$
gp ln Q(ExTrA2)	$\mathcal{U}(-4, -2)$	$-3.933^{+0.11}_{-0.050}$
gp ln ω_0(ExTrA2)	$\mathcal{U}(-4, -1)$	$-1.052^{+0.038}_{-0.075}$
gp ln S_0(ExTrA3)	$\mathcal{U}(-12, 3)$	$2.49^{+0.28}_{-0.31}$
gp ln Q(ExTrA3)	$\mathcal{U}(-4, -2)$	$-3.947^{+0.080}_{-0.039}$
gp ln ω_0(ExTrA3)	$\mathcal{U}(-4, -1)$	$-1.044^{+0.033}_{-0.056}$
Derived parameters		
R_b ($R_\oplus$)		1.200 ± 0.058
$a_b/R_\star$		$29.3^{+1.4}_{-1.6}$
a_b (au)		$0.0272^{+0.0015}_{-0.0016}$
i_b (deg)		$89.26^{+0.33}_{-0.27}$
b_b		$0.38^{+0.11}_{-0.16}$
$T_{\mathrm{tot;b}}$ (h)		1.020 ± 0.017
$T_{\mathrm{full;b}}$ (h)		$0.896^{+0.020}_{-0.021}$
$T_{\mathrm{eq;b}}$ (K)		368^{+12}_{-10}
S_b ($S_\oplus$)		4.35 ± 1.32
$u_{1;TESS}$		0.330 ± 0.062
$u_{2;TESS}$		$0.239^{+0.061}_{-0.056}$
$u_{1;\mathrm{Sloan-z'}}$		$0.240^{+0.054}_{-0.050}$
$u_{2;\mathrm{Sloan-z'}}$		$0.185^{+0.051}_{-0.048}$
$u_{1;\mathrm{Sloan-i'}}$		0.407 ± 0.067
$u_{2;\mathrm{Sloan-i'}}$		0.235 ± 0.063
$u_{1;\mathrm{Sloan-r'}}$		0.617 ± 0.082
$u_{2;\mathrm{Sloan-r'}}$		0.213 ± 0.081
$u_{1;\mathrm{I+z}}$		$0.268^{+0.056}_{-0.053}$
$u_{2;\mathrm{I+z}}$		$0.191^{+0.051}_{-0.049}$
$u_{1;1.21\mu\mathrm{m}}$		0.237 ± 0.048
$u_{2;1.21\mu\mathrm{m}}$		0.184 ± 0.047

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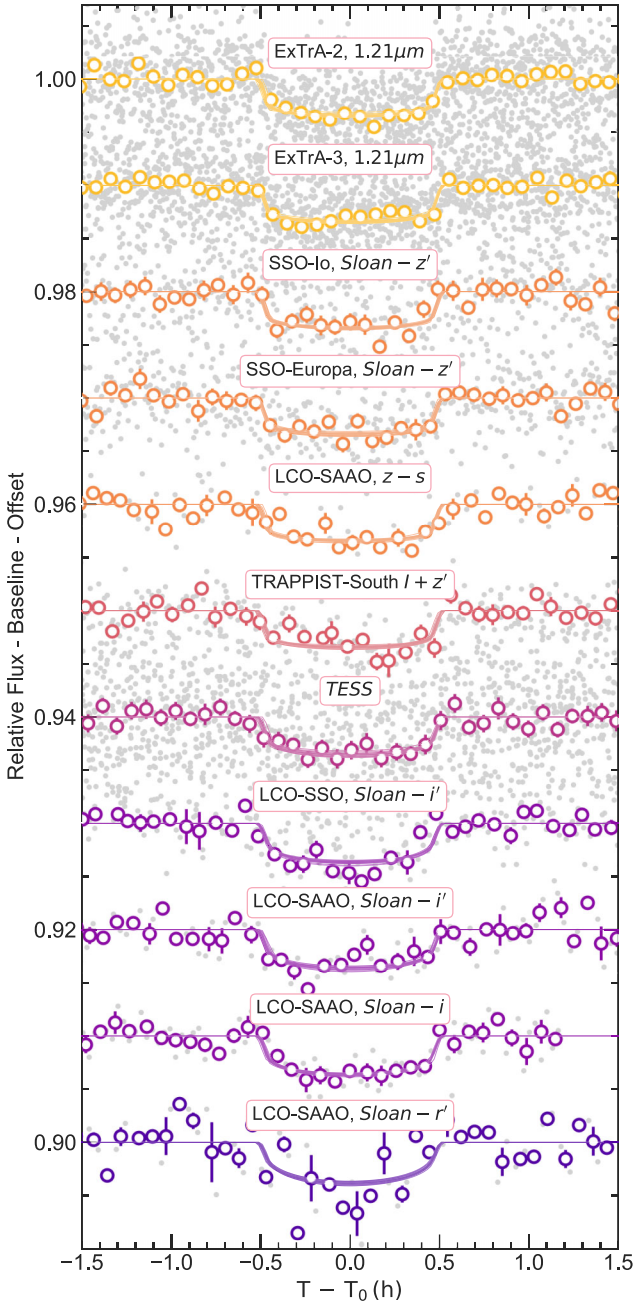


Figure 8. Transit photometry of TOI-1080 b along with best-fitting models. Grey points are the raw flux while white circles are binned to 15 min, coloured according to the filter used in the observation. Photometry from *TESS* and ExTrA are phase folded. The model lines are 20 random draws from the posterior.

spectra have a SNR of about 30 per pixel in the middle of the *H* band and the co-added spectrum has a peak SNR of about 60. For comparison, the reddest HARPS order near 690 nm reached an SNR per pixel of about 5. Although, these data are sufficient for RVs, they cannot be used for accurate stellar characterization of the host.

The NIRPS data were reduced with the APERO pipeline (N. J. Cook et al. 2022), which delivers telluric-corrected spectra. The HARPS observations were processed using version 3.5 of the original HARPS data reduction software (C. Lovis & F. Pepe

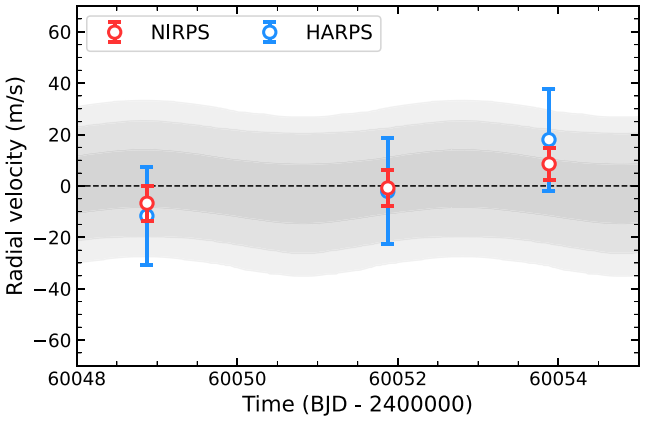


Figure 9. Reconnaissance radial velocity of TOI-1080 with simultaneous NIRPS and HARPS observations. The grey regions show 1, 2, and 3σ confidence envelopes for a fixed ephemeris circular, Keplerian fit with per-instrument offsets (3 free parameters). We derive a 3σ upper limit of 10.7 $M_\oplus$ for the mass of TOI-1080 b.

2007). The RVs were extracted with the line-by-line code (LBL, v0.65.012) described in É. Artigau et al. (2022). To improve RV precision, we used high-SNR templates of GJ 643 (M4V) for NIRPS and GJ 3090 (M2V) for HARPS, both observed as part of the NIRPS-GTO.

We obtain RV precisions of 7 m s^{-1} with NIRPS (nightly bin) and 20 m s^{-1} with HARPS. We show the resulting RVs in Fig. 9, from which we can derive a 3σ upper limit of 10.7 $M_\oplus$ for the mass of TOI-1080 b. This upper limit was obtained by fitting a sinusoid with the period and phase fixed to the transit solution (Section 5) and allowing separate instrumental RV offsets (3 degrees of freedom).

7 DISCUSSION

TOI-1080 is a nearby (∼25 pc) mid-type M dwarf which hosts at least one transiting planet, with a radius of $R_p = 1.200 \pm 0.058 R_\oplus$ and an equilibrium temperature of $T_{eq} = 368^{+12}_{-10}$ K. The planet receives an instellation flux of $S_b = 4.35 \pm 1.32 S_\oplus$, placing it inside the 'temperate zone', defined by M. G. Scott et al. (2026) as $0.1 S_\oplus \leq S_p \leq 5 S_\oplus$ and $T_{eq} \leq 400$ K.

At present we do not have a mass measurement for TOI-1080 b, only an upper limit. However, using the mass–radius relations presented in J. Chen & D. Kipping (2017) we calculate an estimated mass of $M_p = 1.75^{+1.25}_{-0.57} M_\oplus$. We also used SPRIGHT (H. Parviainen et al. 2024) to check this value and find it to be consistent. Planets with radii $< 1.6 R_\oplus$ are likely to be rocky (L. A. Rogers 2015; R. Cloutier & K. Menou 2020; H. Parviainen et al. 2024) and are thus exciting candidates for atmospheric investigations. Those orbiting bright enough stars are particularly amenable. In Fig. 10, we present a mass–radius diagram placing TOI-1080 b in the context of other small, rocky planets within 50 pc of the Sun.

In the remainder of this section we discuss prospects for further analysis of this system. We begin by examining whether precise mass and atmospheric constraints are possible, followed by our insights into the detectability of planetary siblings.

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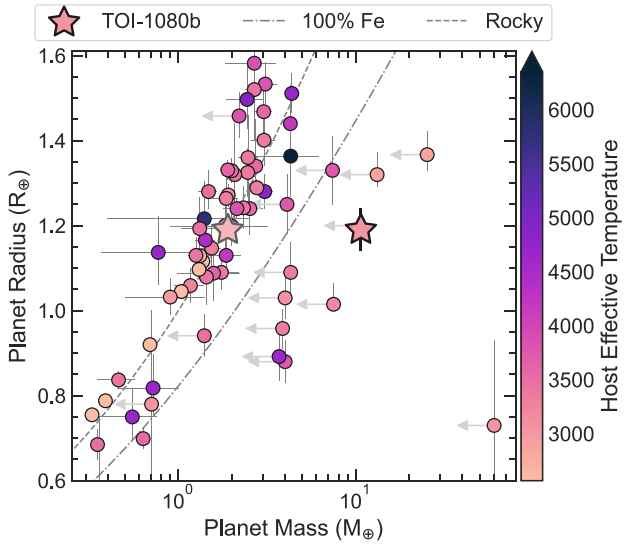


Figure 10. Planets within 50 parsecs with $R_p < 1.6R_\oplus$ from the NASA Exoplanet Archive. Points are coloured according to host star effective temperature and upper limits on masses are indicated with grey arrows. Composition curves for 'rocky' and '100 per cent Fe' are taken from L. Zeng et al. (2019). TOI-1080 b is indicted as a star; the mass value plotted is the upper limit we derive in Section 6; we show as a fainter star the mass derived from the mass–radius relations of J. Chen & D. Kipping (2017) using FORECASTER.

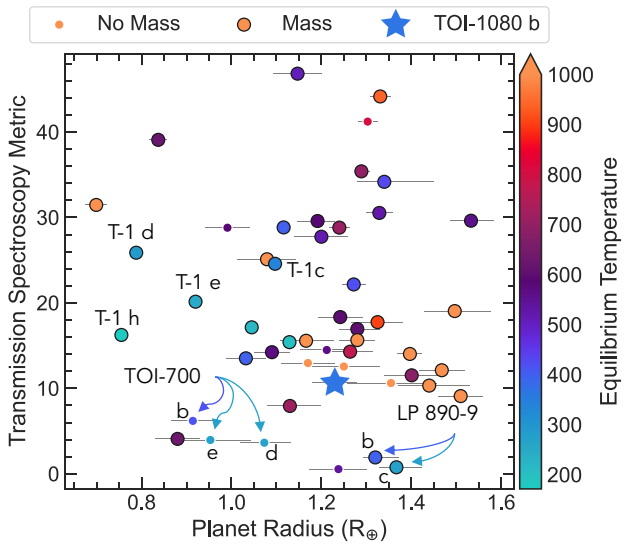


Figure 11. Planets from NASA exoplanet archive with $R < 1.6R_\oplus$ that have already been selected for observations with *JWST* according to TrExoLiSTS (Transiting Exoplanets List of Space Telescope Spectroscopy; N. K. Nikolov et al. 2022). Points are coloured according to planetary equilibrium temperature; small circles have no mass measurement and a mass–radius relation is used to calculate their mass. Larger circles have a mass measurement via radial velocities or transit timing variations.

7.1 Prospects for a precise mass measurement of TOI-1080 b

Based on the the mass estimate calculated above and the stellar mass derived in Section 2, we estimate that TOI-1080 b would produce a radial velocity semi-amplitude of $\sim 2.5\,\mathrm{m\,s^{-1}}$. With NIRPS, the typical precision achieved was $\sim 7\,\mathrm{m\,s^{-1}}$ with 1800 s integrations. Assuming only RV photon-noise without accounting for stellar activity or additional unknown companions, we calculate that it would have taken at least 200 measurements to obtain a 5σ mass measurement. This reduces to 70 spectra to measure the planet mass to 3σ. Therefore, the optimistic estimate of NIRPS telescope time needed would be 100 or 36 h, respectively.

We used the ESPRESSO ETC (exposure time calculator) to estimate the RV precision that could be achieved with this instrument. Given the larger aperture of the VLT, ESPRESSO is well suited to faint stars like TOI-1080. We find that in average observing conditions (airmass ≤ 1.5 and seeing ≤ 1.3 arcsec) we could achieve RV precisions of $5.6\,\mathrm{m\,s^{-1}}$ with 1800 s exposures. With this set-up and not accounting for stellar activity or additional companions, we calculate that 133 or 48 spectra would be needed for 5 or 3σ mass measurements, respectively. Including overheads, this amounts to 85 or 30 h of telescope time, which is a realistic programme for a nearby rocky planet suitable for atmospheric characterization (see Section 7.2). Especially given the scarcity of nearby rocky planets amenable to atmospheric characterisation.

There is a strong precedent for precise mass measurements of small rocky planets orbiting low-mass stars with ESPRESSO. For instance, using 44 ESPRESSO spectra, G. Lacedelli et al. (2024) measured the masses of both planets in the TOI-406 system, consisting of 1.2 and $2.2R_\oplus$ planets with orbital periods of 3.3 and 13 d, respectively. TOI-406 is an M3 host, and about 2 mag brighter than TOI-1080, so naturally we expect more spectra would be needed in our case. In the case that there are additional planets in the system more telescope time would also be needed to resolve both orbits.

7.2 Is TOI-1080 b suitable for atmospheric investigations?

There are several methods used to quantify how suitable a planet is for atmospheric characterisation, the most common being the transmission spectroscopy metric (TSM; E. M.-R. Kempton et al. 2018). We find that TOI-1080 b has a TSM of 10.6, and we contextualise this number in Fig. 11. All planets in this figure have $R_\mathrm{p} < 1.6R_\oplus$ and have already been selected for atmospheric characterization with *James Webb Space Telescope* (*JWST*), according to Transiting Exoplanets List of Space Telescope Spectroscopy (TrExoLiSTS) (N. K. Nikolov, A. Kovacs & C. Martlin 2022). TOI-1080 b has a higher TSM than TOI-700 b, d, and e, (E. A. Gilbert et al. 2020, 2023) and both planets in the LP 890-9 (L. Delrez et al. 2022) system. We also calculate the ESM (emission spectroscopy metric; E. M.-R. Kempton et al. 2018) for TOI-1080 b and find a value of 1.4, comparable to TRAPPIST-1 c which has been observed in emission with *JWST*'s MIRI and had a thick atmosphere ruled out (M. Gillon et al. 2025).

Another way to select suitable targets for atmospheric characterisation is to attempt to determine a priori their likelihood of possessing an atmosphere in the first place. For instance, the 'cosmic shoreline' (K. J. Zahnle & D. C. Catling 2017), which is a theoretical boundary between planets with and without atmospheres, based on instellation flux received and planetary escape velocity. This framework was recently updated by X. Ji et al. (2025), and they use their parametric definition of the CO_2 cosmic shoreline to calculate a priority score for rocky exoplanets. We use equation (7) of this paper to calculate the CO_2 cosmic shoreline for TOI-1080 b, defined as the instellation at which it has 90 per cent

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probability of retaining a CO_2 atmosphere. We find a value of $11.04S_\oplus$, placing TOI-1080 b comfortably below this boundary with its instellation of $S_b = 4.35 \pm 1.32S_\oplus$. This gives a priority score of 0.40; for context, this would place TOI-1080 b 28th on the list of top 50 rocky planets for atmospheric characterization, had it been known at the time X. Ji et al. (2025) was published.

The habitable zone, defined as the region around a star where a rocky planet with atmosphere may retain surface liquid water, is located between 0.047 and 0.124 au for TOI-1080 (R. K. Kopparapu et al. 2013). TOI-1080 b is located outside the habitable zone at 0.0272 ± 0.0016 au. Given that TOI-1080 b receives an instellation flux $> 1.48S_\oplus$, it is expected to start a runway greenhouse, that may result in a massive oxygen (O_2) atmosphere (from a few hundreds to a few thousands bars) after water molecules are photolyzed by ultraviolet (UV) radiation and hydrogen escapes from the planet (R. Luger & R. Barnes 2015; L. Schaefer et al. 2016; E. S. Kite & L. Schaefer 2021; C. Cherubim et al. 2025). Rich abiotic O_2 atmospheres are a robust prediction for short period planets around M dwarfs, like TOI-1080 b, and their characterization could be considered as an intermediate and necessary step towards the search of spectroscopic biosignatures (C. Cherubim et al. 2025).

Thus far, no robust detections of a secondary atmosphere have been made for temperate rocky exoplanets observed in transmission with *JWST* (e.g. N. H. Allen et al. 2025; K. A. Bennett et al. 2025; R. Luque et al. 2025). L. Kreidberg & K. B. Stevenson (2025) show that in most cases, a 5 scale-height atmosphere would still be hidden in the noise for most planets observed, despite the outstanding precision of *JWST* and the large amount of time already invested. Therefore, the Rocky Worlds DDT Programme (S. Redfield et al. 2024; Q. Xue et al. 2025) has selected nine of the most exciting planets to observe in emission, as with this approach a thick atmosphere can be ruled out if there is no evidence of heat redistribution on tidally locked planets. TOI-1080 b is an excellent target for this programme; its priority score as defined by X. Ji et al. (2025) is higher than four out of nine targets selected, and its ESM is higher than two of the selected targets.

7.3 Additional transiting planets in the TOI-1080 system

We first assessed the detection limits of additional transiting planets in TOI-1080 by running Transit Least Squares (TLS; M. Hippke & R. Heller 2019) on the residuals of *TESS* light curve to the best-fitting model described in Section 5. We do not find any significant signal. We then did an injection-recovery test on the same *TESS* residuals where we used BATMAN (L. Kreidberg 2015) to generate synthetic transit signals of planets with radii between 0.38 (like Mercury) to 5 $R_\oplus$ and with orbital periods between 0.5 to 30 d. In all the cases, circular orbits and 90 degrees of inclination were assumed while the mid-transit times were randomly selected between the minimum and maximum time in the *TESS* light curve. Once the planetary signal was injected in the residuals, we used TLS to search for the transits. If the period recovered was within 5 per cent of the injected period, we considered a positive detection.

As shown in Fig. 12, we can rule out the existence of transiting planets down to $\sim$0.9 $R_\oplus$ with orbital periods shorter than $\sim$7.7 d in the TOI-1080 system, as in our simulations the planets were recovered in this period range in $\gtrsim$90 per cent of the injections. We are also able to rule out larger planets than 1.4 $R_\oplus$ with the *TESS* data with periods up to 19 d (better than 90 per cent recovery of the injections). Because the three *TESS* sectors are separated by

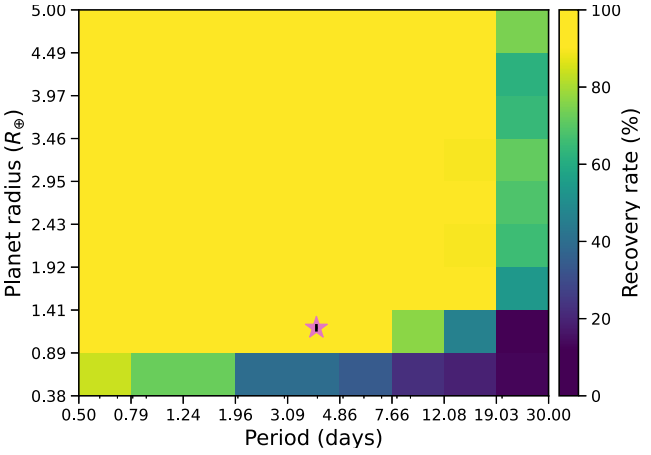


Figure 12. Injection-recovery test performed on the residuals of the *TESS* light curve of TOI-1080. High and low recovery rates are indicated with light and dark colours, respectively. TOI-1080 b is marked by the pink star and the errors on the planet radius are marked by the black vertical line. The errors on its orbital period are too small to be seen in the plot. These injection-recovery tests are able to discard other planets in the system.

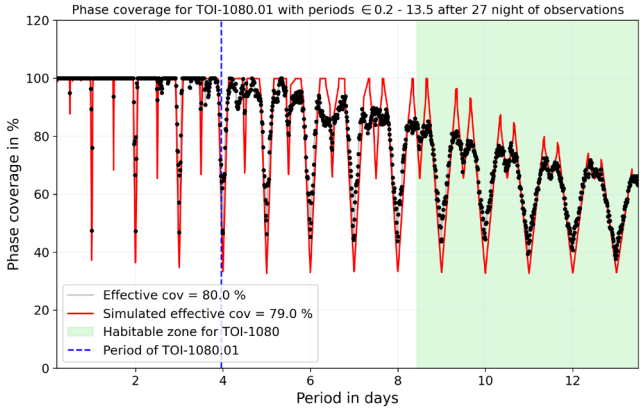


Figure 13. Phase coverage as a function of the orbital period of an hypothetical additional planet around TOI-1080, observed for 243 h in total by SPECULOOS.

2 yr or more, the detection rate starts quickly declining beyond 19-d orbital periods for all planet radii, as expected for individual 28-d long *TESS* sectors. However, planet multiplicity is common for M-dwarf hosts. TOI-1080 is not planned to be observed with PLAnetary Transits and Oscillations of stars (PLATO) based on current LOPS2 field definition, so revealing further planets will need dedicated ground-based monitoring campaign, more *TESS* sectors or alternatively an intensive RV campaign.

Using all the available SPECULOOS observations (243 h), we computed the percentage of phase coverage for all orbital periods from 0.1 to 15 d, following a similar approach to that described in D. Sebastian et al. (2021). We obtained an effective coverage of 80 per cent for $P_{max} = 12.8$ d. This effective coverage metric is defined as the integral of the percentage of phase coverage over the period range from $P = 0.1$ to $P_{max} = 12.8$ d.

Our goal is to estimate the number of additional nights required to achieve an effective phase coverage of 80 per cent over the full optimistic habitable zone (HZ) of TOI-1080, thereby allowing us to significantly detect or rule out HZ planets in this

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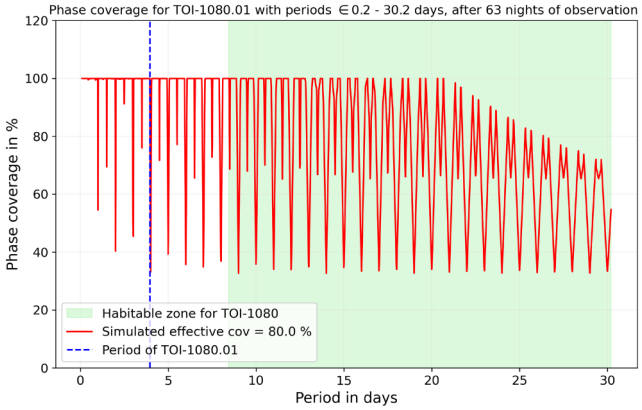


Figure 14. Simulation of the phase coverage as a function of the orbital period for 63 nights to reach effective coverage of 80 per cent for P_{max} set to the period of the outer bound of optimist habitable zone.

system. To this end, we simulated 27 nights of observations – matching the number of nights obtained with SPECULOOS – each lasting 8 h. We verified that the resulting percentage of phase coverage and effective coverage from the simulations are consistent with those derived from the real observations. Fig. 13 confirms this agreement, showing that the simulated and observed coverages are indeed very similar.

We then extended the simulations to N nights, where N is the number of nights required for the effective coverage at $P_{\rm max}$ = Period of the outer HZ limit (computed using R. K. Kopparapu et al. 2014) to reach 80 per cent. We find $N = 67$ nights, implying that 36 additional nights (63–27) of monitoring are needed to fully explore the HZ of TOI-1080 (see Fig. 14).

8 SUMMARY

In this work we have presented the detection and validation of TOI-1080 b, a temperate $R_{\rm p} = 1.200 \pm 0.058 {\rm R}_{\oplus}$, likely rocky planet orbiting an M4 host $\sim$ 25 pc away. We have shown that this planet is suitable for mass measurements with radial velocity with ESPRESSO, as well as atmospheric characterization via transmission spectroscopy, and would be a prime target for the *JWST* + *HST* Rocky Worlds DDT programme. Although inner to the TOI-1080 habitable zone, TOI-1080 b could potentially have a CO_2 or a massive O_2 atmosphere amenable for characterization. We rule out additional transiting planets with radii $> 0.9 R_{\oplus}$ and periods between 0.5 and 7.7 d and for periods up to 19 d we can rule out planets larger than 1.4 $R_{\oplus}$ in the system based on the *TESS* photometry. Given the scarcity of nearby terrestrial planets suitable for in-depth characterization, TOI-1080 b represents an exciting addition to the known sample.

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DATA AVAILABILITY

TESS data products are available via the MAST portal at https://mast.stsci.edu/portal/Mashup/Clients/Mast/Portal.html. Follow-up photometry and high-resolution imaging data for TOI-1080 are available on ExoFOP at https://exofop.ipac.caltech.edu/tess/target.php?id = 161032923. These data are freely accessible to ExoFOP members immediately and are publicly available following a one-year proprietary period. NIRPS and HARPS spectra are available on the ESO Portal following a 1-yr proprietary period.

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