

Breaking news from the HST: The central star of the Stingray Nebula is now returning towards the AGB

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ABSTRACT

SAO 244567 is a rare example of a star that allows us to witness stellar evolution in real time. Between 1971 and 1990 it changed from a B-type star into the hot central star of the Stingray Nebula. This observed rapid heating has been a mystery for decades, since it is in strong contradiction with the low mass of the star and canonical post-asymptotic giant branch (AGB) evolution. We speculated that SAO 244567 might have suffered from a late thermal pulse (LTP) and obtained new observations with HST/COS to follow the evolution of the surface properties of SAO 244567 and to verify the LTP hypothesis. Our non-LTE spectral analysis reveals that the star cooled significantly since 2002 and that its envelope is now expanding. Therefore, we conclude that SAO 244567 is currently on its way back towards the AGB, which strongly supports the LTP hypothesis. A comparison with state-of-the-art LTP evolutionary calculations shows that these models cannot fully reproduce the evolution of all surface parameters simultaneously, pointing out possible shortcomings of stellar evolution models. Thereby, SAO 244567 keeps on challenging stellar evolution theory and we highly encourage further investigations.

Key words: stars: AGB and post-AGB – stars: atmospheres – stars: evolution

1 INTRODUCTION

Revealing the evolution of stars is a complex task, since their evolutionary history can usually only be reconstructed indirectly due to the long evolutionary times scales. In a few cases, however, the evolution of the surface properties of a star occurs on a time scale shorter than a human life-time. Such events allow us to witness stellar evolution in real time and provide a unique way to gain a direct knowledge of stellar evolution.

SAO 244567, the ionization source of the Stingray Nebula (NGC 331.3–12.1), is one such rare example of an unusually fast evolving star. Rapid changes of its observable properties were at first noticed by Parthasarathy et al. (1993, 1995). Based on the optical spectrum of SAO 244567 obtained in 1971 and its *UBV* colors, they concluded that SAO 244567 was a B-type star with an effective temperature of $T_{\text{eff}} \approx 21$ kK at that time. However, the optical spectra from 1990 and 1992 as well as IUE spectra from 1992 on display many nebular emission lines, which let them conclude that SAO 244567 has ionized its surrounding nebula within only two decades. In Reindl et al. (2014, Paper I), we presented first quantitative spectral analyses of all available spectra taken from 1988 to 2006 with various space-based telescopes. We found that the central star had steadily increased its T_{eff} from 38 kK in 1988 to a peak value of 60 kK in 2002. During the same time, the star

was contracting, as concluded from an increase in surface gravity from $\log(g/\text{cm s}^{-2}) = 4.8$ to 6.0 and a drop in luminosity. Simultaneously, the mass-loss rate declined from $\log(\dot{M}/M_{\odot} \text{ yr}^{-1}) = -9.0$ to -11.6 and the terminal wind velocity increased from $v_{\infty} = 1800 \text{ km s}^{-1}$ to 2800 km s^{-1} .

The existence of the planetary nebula (PN) suggests that SAO 244567 is likely a post-AGB star because – in general – only such a star is expected to eject a PN. On the other hand, its low mass is in strong contradiction with canonical post-AGB evolution (Parthasarathy et al. 1995; Bobrowsky et al. 1998). A comparison of the position of SAO 244567 in the $\log T_{\text{eff}} - \log g$ plane with latest post-AGB stellar evolutionary calculations (Miller Bertolami 2016) indicates a mass below $0.53 M_{\odot}$, while the rapid heating rate (dT_{eff}/dt) would require a central star mass of $0.7 M_{\odot}$.

Under certain circumstances, however, the classical picture of the evolution of post-AGB stars is altered by the occurrence of a late He-shell flash. These can occur either during the blue-ward excursion of the post-AGB star (late thermal pulse, LTP), or during its early white dwarf cooling phase (very late thermal pulse, VLTP). The release of nuclear energy by the flashing He-shell forces the already very compact star to expand back to giant dimensions – the so-called born-again scenario. This scenario was first explored in detail by Schönberner (1979) and by Iben et al. (1983) and is of great importance for the explanation of H-deficient post-AGB stars, which make up about a quarter of the post-AGB stars. The evolutionary time scales of stars, which are considered to have un-

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dergone such a late He-shell flash, are very short (decades) and thus, the detection and the repeated observation of such quickly evolving objects make it possible to record the temporal evolution of their stellar parameters and to establish constraints for stellar evolution theory. Well known examples for “born-again” stars are V605 Aql (e.g., Clayton et al. 2006), V4334 Sgr (Sakurai’s object, e.g. Hajduk et al. 2005) and FG Sge (e.g., Jeffery & Schönberner 2006). V605 Aql and V4334 Sgr are considered to have undergone a VLTP, which produces a H-free stellar surface already during the flash. FG Sge, on the other hand, had a slower cooling rate and turned H-deficient only when it had returned back to the AGB. Therefore FG Sge must instead have experienced a LTP (see Schönberner 2008 for a review about these objects).

In Paper I, we speculated that SAO 244567 might also have experienced a late He-shell flash shortly after leaving the AGB. This scenario predicts that SAO 244567 would eventually become a cool supergiant, i.e., that the star will become cooler and expand. To follow the evolution of the surface properties of SAO 244567 and to verify the LTP hypothesis, we performed further observations with the Cosmic Origins Spectrograph (COS) on the Hubble Space Telescope (HST).

2 OBSERVATIONS

We obtained HST/COS spectra using the PSA aperture (total exposure time 13.4 ks, proposal ID 13708). The observations were performed on the 9th of August, 2015, more than nine years after the star was previously observed with an ultraviolet (UV) telescope. Far-UV medium and low resolution spectra were obtained using gratings G130M and G140L at central wavelength settings 1222 Å and 1280 Å, respectively. In addition, we acquired near-UV low resolution spectra with the G230L grating at central wavelength settings 2635 Å and 3360 Å. The latter are unfortunately blurred by the strong [C II] λ 2326 Å (nebular) and [O II] λ 2471 Å (nebular and geocoronal) emission lines and, hence, provide neither reliable fluxes nor line profiles. This is due to the very compact PN ($\phi \approx 2''$) and the much larger aperture used for COS ($2''.5$) observations compared to the apertures used for the former STIS ($52'' \times 0''.05$) and FOS ($0''.3$) observations. Comparing the new HST/COS observation to the FUSE observations in 2002 and 2006, we found that the flux decreased slightly by a factor of 1.6 and 1.4, respectively.

3 SPECTRAL ANALYSIS

The spectral analysis was carried out analogously to Paper I by modelling both the photospheric and the interstellar line-absorption spectrum. Again, we employ the OWENS program to model the ISM line absorption. The interstellar reddening was derived as outlined in Paper I. Using the LMC reddening law of Howarth (1983), we found $E(B-V) = 0.18 \pm 0.02$. For the quantitative spectral analysis we used the Tübingen non-LTE Model-Atmosphere Package (TMAP, Werner et al. 2003, 2012; Rauch & Deetjen 2003), which allows the computation of fully metal-line blanketed model atmospheres in radiative and hydrostatic equilibrium (Sect. 3.1). In order to constrain the presence of a possible weak stellar wind we employed PoWR (Potsdam Wolf-Rayet model-atmosphere code, Gräfener et al. 2002; Hamann & Gräfener 2003, 2004, Sect. 3.2).

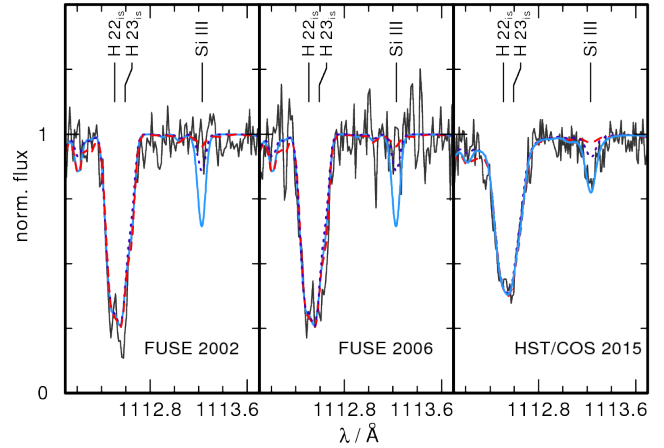


Figure 1. Emergence of Si III λ 1113.2 Å in the HST/COS observation taken in 2015 (right). This line was not visible in the FUSE observation in 2002/2006 (left/middle) when the star was about 10/5 kK hotter. The dotted / dashed line corresponds to TMAP models with $T_{\text{eff}} = 60/55$ kK (best fit for 2002/2006), while the solid line represents a model with $T_{\text{eff}} = 50$ kK (best fit for 2015). Note that the differences in the line profiles are caused by the instruments’ line spread functions.

3.1 TMAP model atmosphere analysis

As a first step we compared the new observations with the best fit model atmosphere for 2006 (year of the last UV observations, see Paper I), which included the elements H, He, C, N, O, Si, P, S, Fe, and Ni to identify several interstellar and photospheric spectral lines. Thanks to the better S/N of the COS G130M spectrum and the extent towards larger wavelengths we were able to identify numerous lines of Fe, and for the first time also Ni. By comparison with the grid calculated in Paper I, we found $\text{Ni} = 6.8 \times 10^{-5}$ (mass fraction, which is solar according to Scott et al. 2015b). Surprisingly, we discovered the emergence of the Si III λ 1113.2 Å line in the new HST/COS spectra (Fig. 1). This line was not visible in the previous FUSE observations, which cover this wavelength range as well and have about the same resolving power ($R \approx 20000$).

Some lines, however remained unidentified. In order to investigate the origin of those lines we proceeded as follows. To check for the presence of lines from iron-group elements, besides Fe and Ni, we calculated TMAP models including H and He, plus Ca, Sc, Ti, V, Cr, or Mn (at abundances 0.1, 1, 10, 100 \times solar, according to Scott et al. 2015a,b), respectively. These were then compared to the G130M observation in order to derive the abundances or at least upper limits. The model atoms were calculated via the Tübingen iron-group opacity interface TIRO (Müller-Ringat 2013). No lines of Sc are predicted in the G130M wavelength range and the strongest line of calcium (Ca IV λ 1096.7 Å) is blended by interstellar lines of H₂₃ and Fe II. Therefore no upper limits are given for these elements. The strongest predicted lines of titanium, namely Ti IV $\lambda\lambda$ 1183.6, 1195.2 Å and Ti V $\lambda\lambda$ 1153.3, 1168.1, 1230.4 Å are not visible in the observations. We derive an upper limit of $\text{Ti} \leq 3.0 \times 10^{-4}$ (100 \times solar). Also the strongest predicted lines of vanadium (V V $\lambda\lambda$ 1142.7, 1157.6 Å), as well as manganese (Mn VI $\lambda\lambda$ 1285.1, 1333.9 Å and Mn V λ 1359.2 Å) are not detected. We derive upper limits of $\text{V} \leq 2.9 \times 10^{-7}$ (solar) and $\text{Mn} \leq 1.0 \times 10^{-5}$ (solar), respectively. For the first time we discovered Cr V $\lambda\lambda$ 1121.1, 1127.6 Å and derived $\text{Cr} = 1.6 \times 10^{-5}$ (solar).

As a next step, we increased the numbers of non-LTE levels for O, Si, P, and S for our model atmosphere calculations. We then calculated a model grid spanning $T_{\text{eff}} = 45 - 60$ kK (step size 2.5 K) and

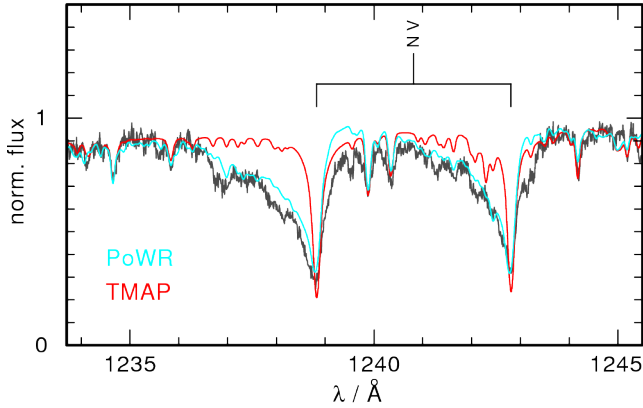


Figure 2. The N v resonance doublet as observed in 2015 with HST/COS (grey). It is not possible to reproduce the asymmetrical line profiles with the static TMAP model (red, $T_{\text{eff}} = 50$ kK, $\log g = 5.5$). In the PoWR model (blue, $10^{-12} M_{\odot} \text{ yr}^{-1}$, $v_{\infty} = 500 \text{ km s}^{-1}$), however, the blue-shifted absorption wings are reproduced well with $v_{\infty} = 500 \text{ km s}^{-1}$. The emission peak of the P-Cygni profiles vanishes when \dot{M} falls below $10^{-12} M_{\odot} \text{ yr}^{-1}$.

including the elements H, He, C, N, O, Si, P, S, Cr, Fe, and Ni (at abundances as derived above and in Paper I). Because the near-UV COS spectrum, which we originally obtained to derive $\log g$ via the He II lines ($n \rightarrow n' = 3 \rightarrow 6, \dots, 14$), is not suited for the spectral analysis (see Sect. 2), we assumed initially $\log g = 6.0$, i.e., that the surface gravity has not changed significantly. Note that the surface gravity is further investigated below, but the next step only focused on the T_{eff} determination by comparing our grid to the new observations.

We found that Si III $\lambda 1113.2 \text{ Å}$ was not visible in the FUSE observation in 2002, because the star was about 10 kK hotter in 2002 compared to 2015 where we find the best fit at $T_{\text{eff}} = 50$ kK. The strengths of Si IV $\lambda\lambda 1128.3, 1154.6 \text{ Å}$ do not show a strong dependence on T_{eff} in the temperature region studied, thus, the Si abundance derived in Paper I is reliable. This value is also confirmed by evaluating additionally the ionization equilibria of C III/C IV, N III/N IV, O III/O IV/O V, S IV/S V/S VI, Fe V/Fe VI, and Ni IV/Ni V. The good S/N allows the error on T_{eff} to be reduced to ± 2.5 kK. For $T_{\text{eff}} \geq 47.5$ kK the lines of N III, O III, Si III, and Si IV are too strong in the models, whereas S VI $\lambda 1117.8 \text{ Å}$ becomes too weak. In contrast, for $T_{\text{eff}} \geq 52.5$ kK the lines of C III, N III, O III, Si III, and Si IV are too weak, while the lines of N IV, S VI, and Fe VI are too strong in our models. The line profiles of C IV, Si IV, S V, Ni IV and Ni V show only mild changes in the range $T_{\text{eff}} = 50 \pm 5$ kK. The same holds for O IV $\lambda\lambda 1338.6 - 1343.5 \text{ Å}$ while the line strength of O IV $\lambda\lambda 1164.3, 1164.5 \text{ Å}$ increases significantly for $T_{\text{eff}} \geq 50$ kK. The latter two lines were not included in the model of Paper I, and therefore not identified. However, these lines are clearly detected in the FUSE observations from 2002 and 2006.

Finally, we investigated a possible impact of $\log g$ on the derived T_{eff} by calculating a second model grid with $\log g = 5.5$. We found that the metal lines hardly differ from the $\log g = 6.0$ model and that the ionization equilibria are also reproduced best at $T_{\text{eff}} = 50 \pm 2.5$ kK. Furthermore, we find that He II $\lambda 1640.4 \text{ Å}$ (Fig. A2) is best reproduced at $\log g = 5.5 \pm 0.5$. For larger/lower values of $\log g$, the line wings of our models become too broad/narrow compared to the observations.

3.2 PoWR model atmosphere analysis

Conclusions about the properties of the stellar wind can be drawn by studying prominent wind lines located in the far-UV. The S/N

of the G140L spectrum is poor for $\lambda < 1050 \text{ Å}$ and cannot be improved sufficiently by data binning, which prevented investigations of the O VI resonance doublet. The C IV resonance doublet, also recorded with the G140L grating, does not show any signature of a stellar wind (Fig. A2). N v $\lambda\lambda 1238.8, 1242.8 \text{ Å}$, however, show asymmetrical line profiles, i.e., they show blue-shifted absorption wings while an emission, which would clearly point to a P-Cygni profile, cannot be detected (Fig. 2). To investigate the origin of the asymmetry, we calculated non-LTE models for expanding model atmospheres with PoWR. We assumed $M = 0.5 M_{\odot}$, $\log(L/L_{\odot}) = 2.5$, $T_{\text{eff}} = 50$ kK and elemental abundances as found in the TMAP analysis. In the absence of diagnostic lines, e.g., recombination emission lines vs. electron scattering line wings, we assume a smooth wind for the calculation of the population numbers. In the formal integral we introduce optically thick inhomogeneities (see Oskinova et al. (2007)) to achieve a better fit of the N v resonance doublet line cores, which otherwise appear too strong in the PoWR model. As the line opacity in the wind is already very low, the absorption wings of the P-Cygni profiles are less affected by the porosity effect.

We determine an upper limit for the mass loss rate of $\dot{M} = 10^{-12} M_{\odot} \text{ yr}^{-1}$. Below this value the emission peaks of the N v lines vanish (Fig. 2). The terminal wind velocity was measured from the blue edge of the absorption component of N v $\lambda 1238.8 \text{ Å}$. We found $v_{\infty} = 500 \text{ km s}^{-1}$, which is significantly less than what was found for the previous years ($v_{\infty} = 2800 \text{ km s}^{-1}$ in the year 2006). This drastic change implies that also the surface gravity must have decreased by half an order of magnitude presuming $v_{\infty} \propto v_{\text{esc}} \propto g$. Therefore a change from $\log g = 6.0$ in 2006 to $\log g = 5.5$ in 2015 is perfectly supported.

Note that the PoWR model (with $\dot{M} = 10^{-12} M_{\odot} \text{ yr}^{-1}$, $v_{\infty} = 500 \text{ km s}^{-1}$) does not predict P-Cygni profiles for other prominent wind lines (e.g., the O VI and C IV resonance doublets). This is in line with previous IUE observations in from 1994 to 1996 in which also only the resonance doublet of N v exhibited P-Cygni profiles and SAO 244567 had $T_{\text{eff}} = 50$ kK as well.

4 RESULTS AND DISCUSSION

The new HST/COS spectra allows us to identify lines of Cr and Ni and to measure their abundances for the first time. We confirm the abundance values of all the other elements derived in Paper I, i.e., no hint of a change in the chemical abundances is found. Figure 3 shows the elemental abundances of SAO 244567 as derived in this work and Paper I. The largely solar surface composition is in line with the thin disk nature of SAO 244567 as concluded from its space velocities. These were calculated using $v_{\text{rad}} = 14 \pm 12 \text{ km/s}$ (measured from the G130M spectrum), proper motions and statistical parallax from Fresneau et al. (2007). We found $(U, V, W) = (32 \pm 18, 27 \pm 13, 4 \pm 6) \text{ km/s}$, which is typical for thin disk stars (Kordopatis et al. 2011). We note, that the height of SAO 244567 above the Galactic plane ($z = 150 \text{ pc}$ assuming $d = 826 \text{ pc}$ from Fresneau et al. 2007) also lies within the scale height of the thin disk.

The solar abundances of He and N excluded the occurrence of the second dredge-up and hot bottom burning. This agrees with the fact that SAO 244567 must be a low mass star since the second dredge-up and hot bottom burning occurs only in stars with initial mass $M > 3 - 4 M_{\odot}$ (García-Hernández et al. 2013). In addition the sub-solar C abundance indicates that SAO 244567 has not experienced a third dredge-up (Paper I), which implies that in

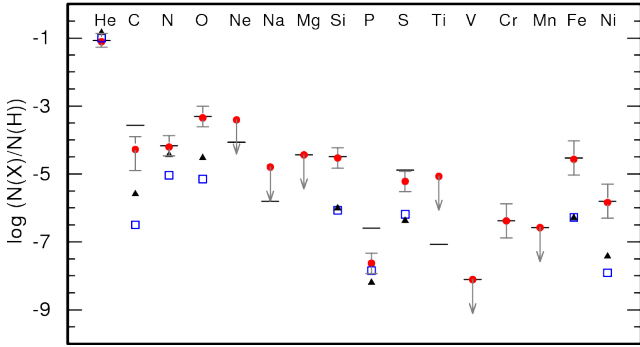


Figure 3. Photospheric abundances of SAO 244567 relative to H (number ratios, red circles). Solar abundance values are indicated by black lines. The abundances of two other lower metallicity low-mass post-AGB stars (ROB 162, blue squares, Chayer in prep., and vZ 1128, black triangles, (Chayer et al. 2015)) are also shown.

case SAO 244567 was once a AGB star, it must have had a low initial mass ($M \lesssim 1.5 M_{\odot}$, Cristallo et al. 2015; Girardi & Marigo 2007). Two well studied low-mass post-AGB stars that also avoided the third dredge-up are vZ 1128 (Chayer et al. 2015) and ROB 162 (Chayer et al. in prep¹), whose abundances are shown in Fig. 3, too. They display very similar abundance patterns to SAO 244567, although the metal abundances are systematically lower by about one order of magnitude. This can, however, be simply understood in terms of the lower initial metallicities of vZ 1128 and ROB 162, since they are found in globular clusters (NGC 5272 and NGC 6397, respectively). Chayer et al. (2015, in prep.) demonstrated that their abundances coincide with the abundances of red giant branch (RGB) and horizontal branch (HB) stars in these clusters, i.e., that the abundances of these stars have not changed since the RGB. The positions of vZ 1128 and ROB 162 in the $\log T_{\text{eff}} - \log g$ plane agree with a $0.53 M_{\odot}$ post-AGB evolutionary track of Miller Bertolami (2016, see Fig. 5 in Reindl et al. 2016), while SAO 244567 seems to have an even lower mass (Fig 4). In Paper I, we discussed the possibilities of SAO 244567 being a post-RGB or post-extreme HB star and we will reconsider these later on. From the star's surface abundances, kinematics, height above the Galactic plane and the position in the $\log T_{\text{eff}} - \log g$ plane, we can, for now, maintain that SAO 244567 is a low-mass, population I star.

The most striking result that we obtained from the new HST observations is the fact that SAO 244567 cooled significantly since 2002. While in 2002 the star had $T_{\text{eff}} = 60 \pm 5$ kK, we find that in 2015 the ionization equilibria of several species are best reproduced at $T_{\text{eff}} = 50 \pm 2.5$ kK. Moreover, we found evidence, that SAO 244567 must have been expanding since 2006. This is concluded from the decrease of the terminal wind velocity by about factor of five since 2006 ($v_{\infty} = 500 \text{ km s}^{-1}$ in 2015, $v_{\infty} = 2800 \text{ km s}^{-1}$ in 2006). This is in line with a decrease of the surface gravity from $\log g = 6.0$ to $\log g = 5.5$. Finally, we found that the mass loss rate of SAO 244567 has further declined ($\dot{M} \leq 10^{-12} M_{\odot} \text{ yr}^{-1}$, compared to $\dot{M} = 10^{-11.6} M_{\odot} \text{ yr}^{-1}$ in 2006).

From the initial rapid heating and contraction of SAO 244567 (heating rate from 1988 to 2002: $1570 \pm 570 \text{ K/yr}$), which is now followed by a rapid cooling and expansion of the star, we conclude that SAO 244567 has indeed suffered a late He-shell flash and is now on its way back towards the AGB. The cooling rate of

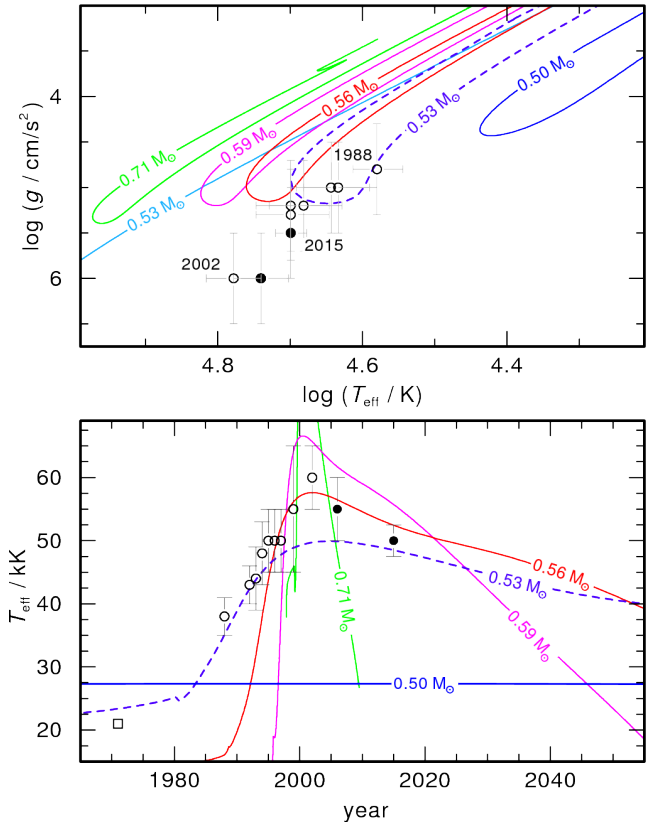


Figure 4. Upper panel: Evolution of SAO 244567 (blue/red-wards evolution is indicated by open/filled circles) in the $\log T_{\text{eff}} - \log g$ compared with different post-AGB evolutionary tracks (labelled with stellar masses, see text for details). The light blue line represents a canonical H-burning post-AGB, while the other models suffer from a LTP. Models that were calculated with $Z = 0.001$ are shown as solid lines, while the dashed line represents a model with $Z = 0.02$. Lower panel: Temporal evolution of T_{eff} as observed and as predicted by the LTP models during the T_{eff} reversal. The open square indicates the T_{eff} estimate when SAO 244567 was still a B-type star (Parthasarathy et al. 1995).

SAO 244567 ($770 \pm 580 \text{ K/yr}$) is similar to that of FG Sge (350 K/yr , Jeffery & Schönberner 2006), the only other hitherto known LTP object. FG Sge was, however, observed only during the final cooling phase, whereas SAO 244567 is now the only known LTP object that was observed during its blue- as well as red-wards evolution. In the upper panel of Fig. 4 we compare the evolution of SAO 244567 in the $\log T_{\text{eff}} - \log g$ plane with five LTP evolutionary tracks that were calculated from the sequences presented in Miller Bertolami (2016) by deliberately removing mass at the end of the AGB to obtain LTPs. While the $0.50 M_{\odot}$ track reproduces the low luminosities ($L \propto T_{\text{eff}}^4/g$) of the star, the model predicts a too low maximum T_{eff} (only about 28 kK, hardly enough to ionize the PN). In addition, the evolutionary speed (lower panel of Fig. 4) is too low. Slightly more massive models do reach ≈ 60 kK, however, at too low surface gravities. The $0.71 M_{\odot}$ track matches the observed gravities, but heats up to $T_{\text{eff}} \approx 100$ kK and evolves too quickly. The heating until 1999 is well reflected by the $0.53 M_{\odot}$ model, while the $0.56 M_{\odot}$ model nicely reproduces the evolutionary speed from 1996 until 2015. Fine-tuning of the LTP model parameters may, however, lead to a better agreement of the observations with theoretical predictions.

Let us now critically examine this LTP scenario in view of other possible evolutionary paths. He-shell flashes may also occur in

¹ http://www-astro.physics.ox.ac.uk/~aelg/SDOB7/proceedings/Chayer_rob162_proceedings.pdf

post-EHB stars (Dorman et al. 1993) or after the merger of two low-mass stars (Zhang & Jeffery 2012). The post-EHB scenario was already ruled out in Paper I, since it is not possible to explain the young PN. In case of a merger event, He-shell flashes occur on too long time scales ($\approx 20,000$ yrs, Zhang & Jeffery 2012) and also the largely solar composition does not support a merger scenario. Finally, let us consider SAO 244567 as a post-RGB star which ended a common-envelope phase (during which the PN was ejected) in thermal non-equilibrium. Firstly, we could not find any hint of relative motion of interstellar and photospheric lines by comparing the HST/COS observations with the FUSE observations in 2002 and 2006. Therefore, there is no proof for a close companion, which would be required for such a scenario. Secondly, and more importantly, in such a scenario the star would directly, although rapidly, evolve to become a white dwarf (Deinzer & von Sengbusch 1970; Hall et al. 2013). Therefore SAO 244567 would be expected to contract further, and not, as shown in this work, expand again. The only way to explain the expansion nevertheless would be, if the star had undergone a H-shell flash after entering the white dwarf cooling sequence. Such H-shell flashes occur, however, only when the star had reached $L \approx 1 L_{\odot}$ and additionally the cooling rate would be orders of magnitude lower (Istrate et al. 2016, 2014; Althaus et al. 2001, 2013) than what is observed for SAO 244567.

5 CONCLUSIONS

In this work we have shown, that after the initial rapid heating and contraction (Paper I), SAO 244567 has cooled significantly since 2002 and is now expanding. This can only be explained with a LTP scenario, and indicates that the star is now on its way back to the AGB. We stress that SAO 244567 is the only LTP object, that was observed during its blue- and red-wards motion through the Hertzsprung-Russell diagram. The evolutionary speed suggests a central star mass between 0.53 and $0.56 M_{\odot}$. However, none of the current LTP models can fully reproduce the evolution of all surface parameters simultaneously. In particular the problem with the apparently high gravity of the star should be pursued further. A high S/N near-UV spectrum would therefore be highly desirable in order to derive $\log g$ more precisely. Refined LTP evolutionary calculations, on the other hand, may not only help to explain the nature of SAO 244567, but could also provide a deeper insight into the evolution of central stars of PNe as well as the formation of H-deficient stars.

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APPENDIX A: IMAGES

In the following we compare our best fit TMAP model to the HST/COS observations.

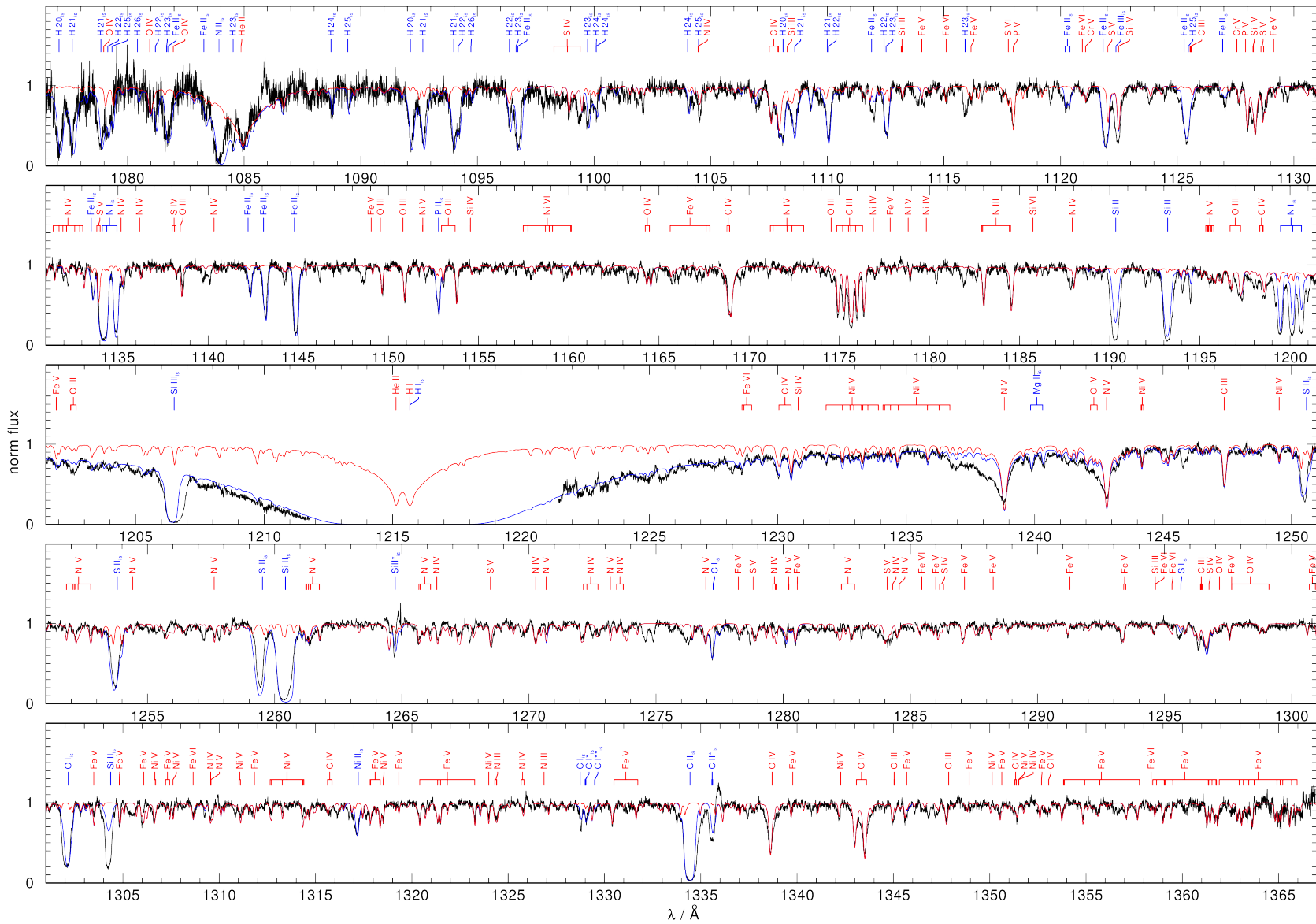


Figure A1. G130M HST/COS observation taken in 2015 overplotted with our best fit TMAP model ($T_{\text{eff}} = 50 \text{ kK}$, $\log g = 5.5$). The red line indicates the pure stellar model spectrum, the blue line additionally includes the interstellar lines. Photospheric and interstellar lines are marked.

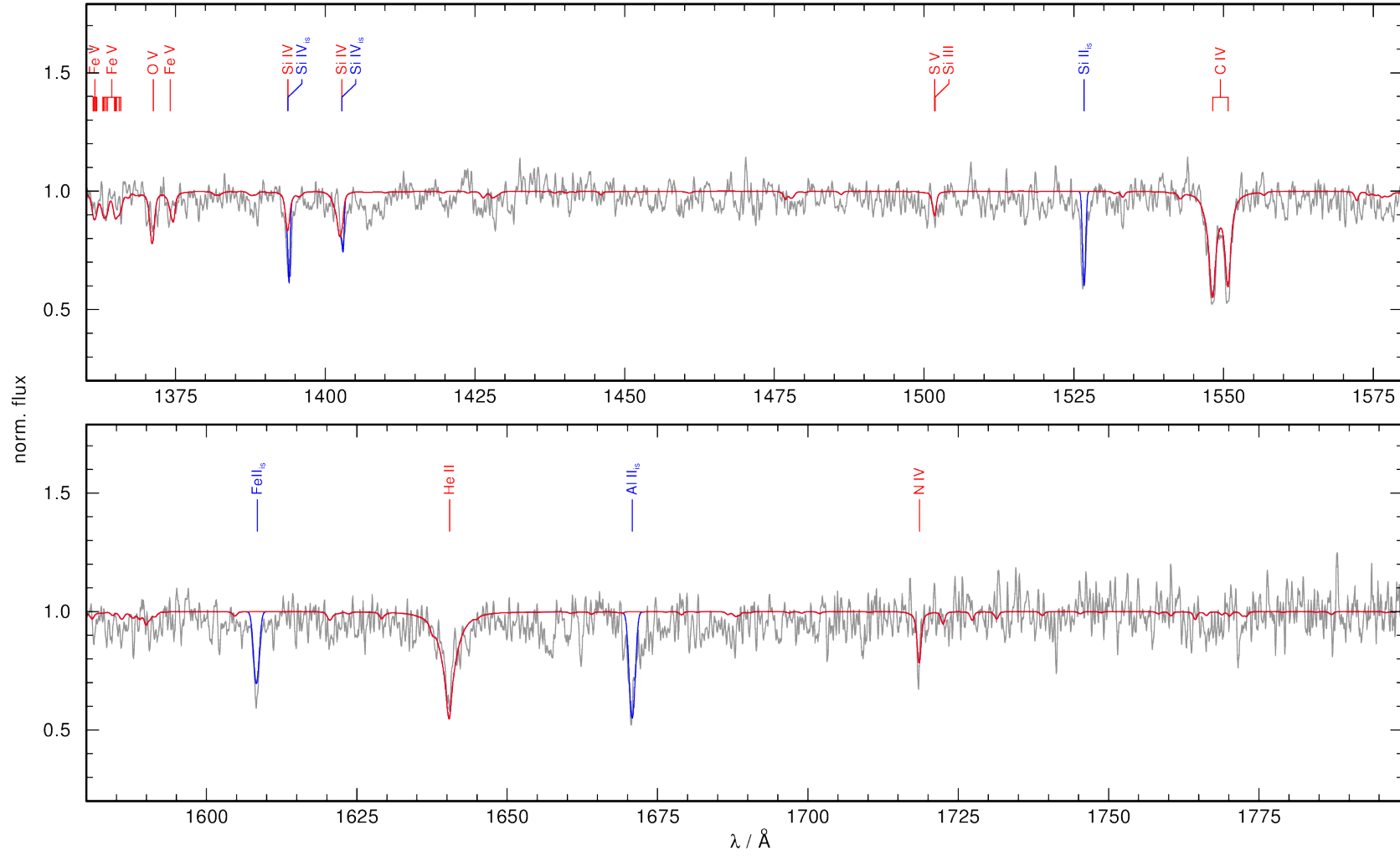


Figure A2. Like Fig. A1 for the G140L spectrum.

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