

SN 2011fu: A type IIb Supernova with a luminous double-peaked light curve

A. Morales-Garoffolo¹*, N. Elias-Rosa², M. Bersten^{3,4,5}, A. Jerkstrand⁶, S. Taubenberger^{7,8}, S. Benetti², E. Cappellaro², R. Kotak⁶, A. Pastorello², F. Bufano⁹, R. M. Domínguez¹⁰, M. Ergon¹¹, M. Fraser¹², X. Gao¹³, E. García^{14,15,16}, D. A. Howell^{17,18}, J. Isern¹, S. J. Smartt⁶, L. Tomasella², S. Valenti^{17,18}

¹*Institut de Ciències de l'Espai (CSIC-IEEC), Campus UAB, Camí de Can Magrans S/N, 08193 Cerdanyola (Barcelona), Spain*

²*INAF Osservatorio Astronomico di Padova, Vicolo dell'Osservatorio 5, 35122 Padova, Italy*

³*Instituto de Astrofísica La Plata, IALP (CCT La Plata), CON ICET-UNLP, Paseo del Bosque s/n, 1900 La Plata, Argentina*

⁴*Facultad de Ciencias Astronómicas y Geofísicas, Universidad Nacional de La Plata, Paseo del Bosque s/n, 1900 La Plata, Argentina*

⁵*Kavli Institute for the Physics and Mathematics of the Universe (WPI), The University of Tokyo, Kashiwa, Chiba 277-8583, Japan*

⁶*Astrophysics Research Centre, School of Mathematics and Physics, Queen's University Belfast, Belfast BT7 1NN, UK*

⁷*Max-Planck-Institut für Astrophysik, Karl-Schwarzschild-Str. 1, 85741 Garching bei München, Germany*

⁸*European Southern Observatory, Karl-Schwarzschild-Str 2, 85748 Garching bei München, Germany*

⁹*INAF-Osservatorio Astrofisico di Catania, Via S.Sofia 78, 95123, Catania, Italy*

¹⁰*Instituto de Física de Cantabria (IFCA), Observatorio Astronómico de Cantabria, Av. de los Castros, 39005 Santander (Cantabria), Spain*

¹¹*The Oskar Klein Centre, Department of Astronomy, Stockholm University, Albanova, 10691 Stockholm, Sweden*

¹²*Institute of Astronomy, University of Cambridge, Madingley Road, Cambridge CB3 0HA, United Kingdom*

¹³*No.1 Senior High School, Urumqi, Xinjiang, China*

¹⁴*Departamento de Física Aplicada I, E.T.S. Ingeniería, Universidad del País Vasco, 9 Alameda Urquijo s/n, 48013 Bilbao, Spain*

¹⁵*Esteve Duran Observatory, Foundation Observatori Esteve Duran. Avda. Montseny 46, Seva 08553, Spain*

¹⁶*Institut d'Estudis Espacials de Catalunya (IEEC), Edif. Nexus, C/ Gran Capità 2-4, 08034 Barcelona, Spain*

¹⁷*Las Cumbres Observatory Global Telescope Network, 6740 Cortona Dr., Suite 102, Goleta, CA 93117, USA*

¹⁸*Department of Physics, University of California, Santa Barbara, Broida Hall, Mail Code 9530, Santa Barbara, CA 93196-9530, USA*

ABSTRACT

We present optical and near infrared observations of the type IIb supernova (SN) 2011fu from a few days to ~ 300 d after explosion. The SN presents a double-peaked light curve (LC) similar to that of SN 1993J, although more luminous and with a longer cooling phase after the primary peak. The spectral evolution is also similar to SN 1993J's, with hydrogen dominating the spectra to ~ 40 d, then helium gaining strength, and nebular emission lines appearing from ~ 60 d post-explosion. The velocities derived from the P-Cygni absorptions are overall similar to those of other type IIb SNe. We have found a strong similarity between the oxygen and magnesium line profiles at late times, which suggests that these lines are forming at the same location within the ejecta. The hydrodynamical modelling of the pseudo-bolometric LC and the observed photospheric velocities suggest that SN 2011fu was the explosion of an extended star ($R \sim 450 R_{\odot}$), in which 1.3×10^{51} erg of kinetic energy were released and $0.15 M_{\odot}$ of ^{56}Ni were synthesised. In addition, a better reproduction of the observed early pseudo-bolometric LC is achieved if a more massive H-rich-envelope than for other type IIb SNe is considered ($0.3 M_{\odot}$). The hydrodynamical modelling of the LC and the comparison of our late-time spectra with nebular spectral models for type IIb SNe, point to a progenitor for SN 2011fu with a ZAMS mass of 13–18 M_{\odot} .

Key words: supernovae: general, supernovae: individual: SN 2011fu.

1 INTRODUCTION

Core-Collapse-Supernovae (CC-SNe) are believed to be the explosions that mark the end-point in the evolution of mas-

* amg@ice.csic.es

sive stars ($M_{ZAMS} > 8 M_{\odot}$; see e.g. Heger et al. 2003). They are spectroscopically divided in different groups according mainly to the presence of the H and He lines in their optical spectra. While type II SNe show prominent H features, type I do not, and can be further subdivided, according to the presence or absence of He, as types Ib and Ic respectively. There are also hybrid objects which undergo a transition from being H dominated at early phases of their spectral evolution, to He dominated at later times. These are termed type I Ib SNe. The first SN to have shown this spectral transition was SN 1987K, and its progenitor was proposed to be a star that had lost most of its hydrogen envelope before exploding (Filippenko et al. 1988). The best studied type I Ib SN to date is SN 1993J (e.g. Nomoto et al. 1993; Podsiadlowski et al. 1993; Woosley et al. 1994), that is considered the prototype of the subclass. Its LC showed an early peak, followed by a rapid decline thought to be the consequence of the cooling of the progenitor’s stellar envelope after shock breakout. Subsequently it presented a secondary maximum attributed to input from the radioactive decay of ^{56}Ni . The star that exploded as SN 1993J was detected in archival images (Aldering et al. 1994), and its disappearance was confirmed by Maund & Smartt (2009). Two mechanisms have been proposed by which type I Ib SN progenitors lose part of their H envelope before exploding, namely stripping by a close companion after Roche Lobe overflow, and wind-driven stellar mass loss in a single star. The former scenario has gained strength over the past years, for example with the detection of signatures of the possible hot blue companion star of SN 1993J (Maund et al. 2004; see also Fox et al. 2014), and the modelling of a progenitor binary system and likely detection of the stellar companion of SN 2011dh (Benvenuto et al. 2013; Folatelli et al. 2014). Moreover, although a Wolf-Rayet (WR) wind-like spectrum was obtained for the type I Ib SN 2013cu a few hours after explosion (Gal-Yam et al. 2014), recent work has shown that the progenitor star for SN 2013cu was unlikely a WR (Groh 2014; Shivvers et al. 2014; Smith et al. 2015).

An important question is which of the observed properties of SNe I Ib can be connected with the characteristics of their progenitor systems. For example, Chevalier & Soderberg (2010) proposed that bright early luminosity from the shock-heated progenitor stellar envelope, low radio shell velocities, and thermal X-ray emission were characteristic of extended progenitors with $R \sim 150 R_{\odot}$, while compact progenitors with $R \sim 1 R_{\odot}$ have faint early optical LCs, have high radio shell velocities and non-thermal X-ray emission. However, this link is not straightforward and needs to be revised since the type I Ib SNe 2011dh and 2011hs present fast expanding radio shells but the hydrodynamical modelling of their LC (Bersten et al. 2012; Bufano et al. 2014), and in the case of SN 2011dh the detection of its progenitor in archival images (Maund et al. 2011; Van Dyk et al. 2011) and its disappearance (Van Dyk et al. 2013), point to the explosion of extended stars. To further complicate the picture, SNe I Ib show a variety of continuum flux excess in early UV spectra (Ben-Ami et al. 2014). Strong UV-excess suggest significant interaction between the SN ejecta and circumstellar material.

To date, few type I Ib SNe have been thoroughly studied since they are relatively rare events. Li et al. (2011) estimated the fraction of type I Ib over a volume limited

sample of 81 type II SNe to be $11.9\%_{-3.6}^{+3.9}$. In this paper we present optical and near infrared (NIR) data for the type I Ib SN 2011fu. SN 2011fu, with coordinates $\alpha = 02^{\text{h}}08^{\text{m}}21.26^{\text{s}}$ and $\delta = 41^{\circ}29'9.9''$ (J2000), was discovered in a spiral arm of the galaxy UGC 1626 by F. Ciabattari and E. Mazzoni of the Italian Supernovae Search Project (ISSP)¹, on 2011 September 21.04 UT and classified by Tomasella et al. (2011) as a young type II SN 2011 September 23.84 UT. A previous analysis of optical data of SN 2011fu was presented by Kumar et al. (2013), which confirmed it was a type I Ib SN. In this manuscript we contribute with data probing the whole SN evolution. In Section 2 we discuss the distance, reddening, and explosion epoch of SN 2011fu. In Section 3 we summarize our observations and the reduction process of the data. In Section 4 we present the optical and NIR LCs of the object, while in Section 5 we present and discuss the spectral evolution. In Section 6 we discuss the results obtained from the hydrodynamical modelling of the pseudo-bolometric LC and we compare our nebular spectra with published models for late time type I Ib SN spectra. Finally, in Section 7, we present the conclusions of our work.

2 DISTANCE, REDDENING, AND EXPLOSION EPOCH

The rise to a first maximum in the V and R LCs (see the inset of Figure 2) implies that SN 2011fu was discovered very soon after explosion. In fact, models show that the evolution of the SN LC during this rise should be very fast, and lasts ~ 1 d for a number of configurations (e.g. Blinnikov et al. 1998; Bersten et al. 2012). By adopting $\text{JD}_{\text{expl}} = 2455824.5 \pm 0.7$ as the explosion date of SN 2011fu, we obtain the best fit of the early phase of the pseudo-bolometric LC with the hydrodynamical models presented in Section 6.1. For this reason we have chosen $\text{JD}_{\text{expl}} = 2455824.5 \pm 0.7$ as the explosion date of SN 2011fu, which is in agreement with the discovery epoch, the last non detection of the object which was 2011 August 10 (JD = 2455783.5), the pre-discovery image taken by Xing Gao 2011 September 20 (JD = 2455825.2), and the classification as a young object. Note that this assumed explosion date also agrees, within the uncertainties, with the explosion epoch adopted by Kumar et al. (2013). All phases in the rest of this manuscript are given with respect to $\text{JD}_{\text{expl}} = 2455824.5 \pm 0.7$.

UGC 1626 is a SAB(rs)c type galaxy whose recessional velocity given by the NASA/IPAC extragalactic database (NED) is $5543 \pm 11 \text{ km s}^{-1}$. We have performed measurements of the SN redshift from the narrow H α emissions in its spectra and have obtained the same redshift as for its host galaxy, which we adopt in the rest of this paper. The redshift derived distance to UGC 1626, also provided by NED assuming $H_0 = 73 \pm 5 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.27$, $\Omega_v = 0.73$, and accounting for the Virgo, Great Attractor and Shapley infalls is $74.5 \pm 5.2 \text{ Mpc}$, i.e. $\mu = 34.36 \pm 0.15 \text{ mag}$, (Mould et al. 2000).

The reddening in the line of sight of UGC 1626 due to the Milky Way is $E(B-V)_{\text{MW}} = 0.068 \pm 0.002 \text{ mag}$ (Schlafly & Finkbeiner 2011). We detect a narrow absorption likely

¹ <http://italiansupernovae.org/en.html>

due Na I D $\lambda\lambda 5890, 5896$ at the host galaxy redshift. The equivalent width of this line is $0.33 \pm 0.03 \text{ \AA}$ that using the relations for unresolved Na I D given by Poznanski et al. (2012), provides an estimate of the reddening in the host galaxy as: $E(B-V)_{\text{host}} = 0.035^{+0.042}_{-0.029}$ mag. In fact, a similar value of reddening due to the host galaxy is obtained using the relations of Turatto et al. (2003). In the rest of this manuscript we adopt $E(B-V)_{\text{total}} = 0.10^{+0.04}_{-0.03}$ mag as the total reddening towards SN 2011fu. We note that in Kumar et al. (2013) the relations of Munari & Zwitter (1997) were used to estimate the contribution to the reddening due to the host galaxy, and they obtained $E(B-V)_{\text{total}} = 0.22 \pm 0.11$, which is consistent with the value we have adopted within the uncertainties.

3 OBSERVATIONAL DATA, REDUCTION, AND CALIBRATION PROCESS

The observational follow-up data of SN 2011fu presented in this paper covers the period between a few days post discovery, until the SN disappeared behind the Sun (end of February 2012). After that, one more spectrum and one epoch of *gri* photometric data were obtained well into the nebular phase of the SN, 2012 July 20 UT. Thanks to a large collaboration between many European institutions, optical photometry and spectroscopy, as well as some NIR photometric data of SN 2011fu, have been collected at several sites. Some amateur data taken by F.Ciabattari and E. Mazzoni from Italy, and Xing Gao from China were also included in our analysis. A summary of the characteristics of the instruments used to acquire our photometric data can be seen in Table A1.

The list with our spectroscopic observations and the instrumental configuration used to acquire the data is presented in Table A2.

All the data were corrected for overscan, bias, and flat-field within the IRAF² environment, except for data obtained at the Liverpool Telescope (LT) which were automatically reduced with the instrument specific pipeline. The instrumental magnitudes of the SN were derived via Point Spread Function (PSF) fitting which was done with SNOOPY (Cappellaro 2014). In the case of the *U* and *u* photometry, the measurements were done after performing template subtraction. For all other bands, PSF photometry at late time was verified by performing template subtractions and no major differences were found³. A set of fifteen local stars in the SN field (Figure 1) were used to trace the photometric calibration. The stellar sequence was calibrated with the zero points and colour terms derived for each site thanks to observations of Landolt standards (Tables B4 and B5) and the SN magnitudes were calculated relative to these. The errors of the magnitudes of the stellar sequence were obtained as the rms deviation with respect to their mean values. Given

² Image reduction and Analysis Facility, a software system distributed by the National Optical Astronomy Observatories (NOAO)

³ The *UBVRI* band templates were obtained on 2012 October 23 with AFOSC at the 1.82 m telescope in the Asiago Observatory, while the *uri* band templates were obtained on 2014 October 25 at the LT with IO:O in the *Roque de los Muchachos* Observatory.

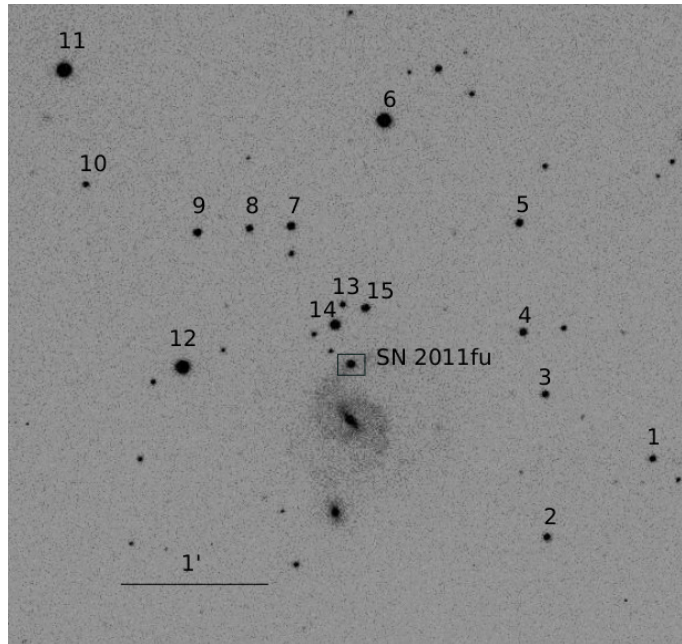


Figure 1. *V* band image of UGC 1626 taken with the 2.2m Calar Alto Telescope + CAFOS on 2011 October 16. The stars used for the photometric calibration of SN 2011fu are labelled. North is to the top and East to the left.

the fact that several instruments with different passbands were used to acquire our data, we performed S-corrections of our *BVRriI* photometry to the Bessell system (Stritzinger et al. 2002; Pignata et al. 2004). In order to derive the S-correction terms we used our spectral sequence of the SN. Note that we were not able to perform S-corrections to the *U*, *u*, and *z* band photometry given the fact that our spectra do not cover completely these passbands. The SN magnitudes together with their errors, calculated by adding in quadrature the uncertainties associated with the calibration and the PSF fit⁴, are reported in Tables B1, B2, and B3.

The unfiltered photometry provided by F. Ciabattari was rescaled to the *R* band considering that the detector sensitivity peaks at this wavelength. With the same argument the data provided by X. Gao were rescaled to *V*.

Note that stars 2, 6, 9, 12, and 14 from our sequence coincide with stars 2, 1, 4, 3 and 5 used for photometric calibration by Kumar et al. (2013). The greatest difference between the two sequences is 0.1 mag found for star 2 in the *U* band. For the rest of the bands differences are found to be less than 0.1 mag.

The NIR data was reduced within IRAF with the help of CAINDR, a package developed by José A. Acosta Pulido and Rafael Barrena for the instrument CAIN3 at the TCS. The images of each specific filter were corrected for flat field, and bad pixel corrections were applied. Then background images were constructed and subtracted from the science frames, and finally the images were combined taking into account the dithering pattern. Similarly to the optical data,

⁴ Obtained by placing artificial stars with magnitudes similar to those of the SN at positions near the SN and calculating the standard deviation of their recovered magnitudes.

we measured the instrumental SN magnitudes by PSF fitting. The SN calibration was done relative to 2-MASS stars in the field (Skrutskie et al. 2006; see Table B4).

The spectra were corrected for overscan, bias, and flat-field within IRAF. Next they were variance-weighted extracted and wavelength calibrated with the use of arc lamp spectra. The wavelength calibration was cross-checked with respect to sky emission lines. Finally a flux calibration was applied using sensitivity functions obtained from the observation of spectrophotometric standards. The flux calibration was cross-checked with the photometry of the SN at the nearest epoch. Telluric lines were corrected dividing the SN spectra by telluric line spectra obtained from the spectrophotometric standards.

4 PHOTOMETRY

4.1 Light curves

The optical-NIR LCs of SN 2011fu are presented in Figure 2. In addition to the data presented in the figure we obtained two epochs of Sloan *g* band data, which are listed in Table B2. In the optical, the LCs clearly present two peaks in all bands, but not in the NIR since the data do not cover the early phases. The rise to primary peak in the *V* and *R* band can be seen thanks to the early time data provided by amateur astronomers (see inset in the top-right of the figure).

By making low order polynomial fits to the optical LCs we have estimated the phases and magnitudes at which the minima after first peak and secondary maxima take place. The results are presented in Table 1. Note that we obtain different absolute magnitudes at secondary peak than Kumar et al. (2013) due to the fact that we have adopted different distance and reddening towards the host, however, our values agree within the uncertainties. We also fitted low order polynomials to the first peaks in the *R* and *V* LCs. In the *V* band the first peak at 16.3 mag is reached ~ 2.3 d after explosion while in the *R* band the 16.0 mag first peak is obtained ~ 2.8 d post explosion.

We have also estimated the decline rates of the tails of the *BVRIZ* LCs (Table 1). These rates are quite similar to those estimated for SN 2013df by Morales-Garoffolo et al. (2014), and are steeper than expected from ^{56}Co decay. This is a common characteristic to stripped envelope SNe (e.g., SN 2008ax, Taubenberger et al. 2011) and is possibly due to increasing transparency for γ rays in their lower mass ejecta.

4.1.1 Colour curves

Figure 3 shows the intrinsic colour evolution of SN 2011fu and for comparison the data for other type IIb SNe. such as SN 1993J (Lewis et al. 1994, Barbon et al. 1995, Richmond et al. 1994, Matthews et al. 2002; assumed explosion date $\text{JD}_{\text{exp}} = 2449074.0$), SN 2008ax (Pastorello et al. 2008, Taubenberger et al. 2011, Tsvetkov et al. 2009; assumed explosion date $\text{JD}_{\text{exp}} = 2454528.8$), SN 2011dh (Ergon et al. 2015; assumed explosion date $\text{JD}_{\text{exp}} = 2455713.0$), SN 2013df (Morales-Garoffolo et al. 2014; assumed explosion date $\text{JD}_{\text{exp}} = 2456450.0$). The adopted extinctions along the line of site of the SNe plus their distance are the same as

those given in Table 7 of Morales-Garoffolo et al. (2014). All colour indices of SN 2011fu show a smooth rise from explosion up to ~ 40 days, that is ~ 15 d after secondary peak, and afterwards they have a blue-ward trend.

4.2 Pseudo-bolometric LC

To obtain the pseudo bolometric optical-NIR LC, first we converted the apparent magnitudes of SN 2011fu (corrected for extinction) to effective fluxes. At the phases at which there were no data, the missing points were obtained by interpolation, or as in the cases of the *U* and *u* bands at more than ~ 90 d, and NIR prior and past secondary maximum, by extrapolation assuming a constant colour from the data at the nearest epoch. The fluxes were integrated over wavelength following a trapezoidal rule, and finally converted to luminosities taking into account the adopted distance to the SN.

In Figure 4 we present the pseudo-bolometric NIR-optical LC of SN 2011fu as well as those of SNe 1993J, 2008ax, 2011dh, and 2013df, which we calculated in a similar way. In comparison to SNe 1993J and 2013df, SN 2011fu presents a longer cooling phase after the first peak, and a longer rise time to secondary peak than the rest of the objects, which makes SN 2011fu the one with the highest ^{56}Ni mass synthesised in the explosion, in accordance with Kumar et al. (2013) and our modelling of the LC, which we present in Section 6.1. The slope of the LC tail starting at 40 d and up to ~ 300 d is a factor 1.1 steeper for SN 1993J, indicating as well that SN 2011fu ejected more mass in its explosion.

5 SPECTROSCOPY

5.1 Spectral evolution of SN 2011fu

In Figure 5 we show the spectral evolution of SN 2011fu ranging from ~ 4 d to ~ 304 d after explosion. The early spectra show a blue continuum and shallow features. Starting at around 6 d $\text{H}\alpha$, $\text{H}\beta$ (with some possible Fe II contamination), He I $\lambda 5876$, Ca II H & K, Ca II NIR $\lambda\lambda 8498, 8542, 8662$, and Fe II $\lambda 5169$ appear to grow. The $\text{H}\alpha$ P-Cygni absorption component at 11.1 d shows a small trough at around 6300 Å in addition to a broader and deeper one centred at approximately 6190 Å. These double $\text{H}\alpha$ absorption features have been seen before in some type IIb SNe (e.g. SN 2011hs; Bufano et al. 2014) and have been claimed to be either due to Si II (Hachinger et al. 2012) or the presence of a double density distribution of hydrogen in the ejecta of the SNe as claimed for some type II SNe (e.g. Inserra et al. 2013). Note that in the first spectrum and the one at 14 d, there is a narrow emission line at approximately the rest wavelength of $\text{H}\alpha$. This line is also seen in some of our subsequent spectra, however, as we will explain below, we believe it is not associated with the SN or its circumstellar medium (CSM). After the LC's secondary maximum, at 34.1 d, two new lines with fairly narrow absorptions at approximately 4840 Å and 4890 Å are discernible, which we believe are associated with Fe II $\lambda 4924$ and Fe II $\lambda 5018$ + He I $\lambda 5015$ respectively. In addition there is an absorption that could be associated with O I $\lambda 7774$. A major change occurs at ~ 40 d,

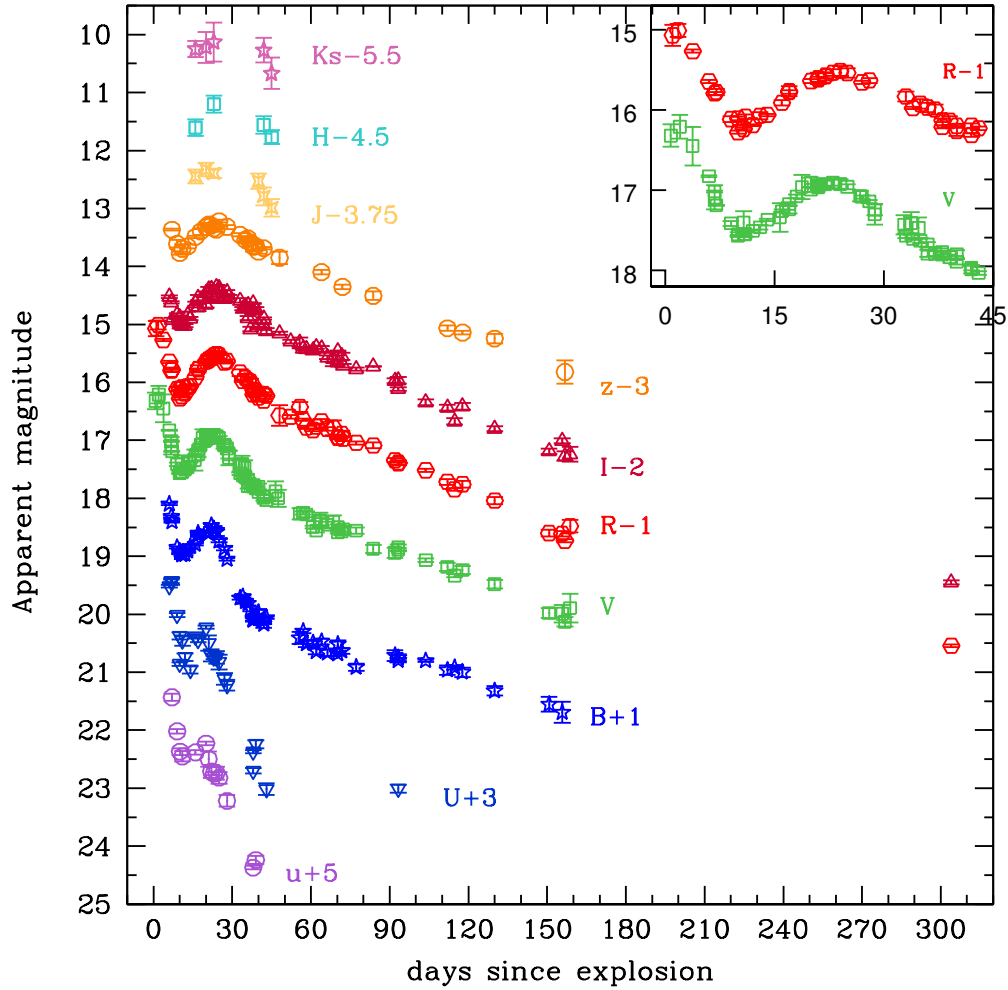


Figure 2. Optical-NIR light curves of SN 2011fu. The assumed explosion epoch is $\text{JD} = 2455824.5 \pm 0.7$. The LCs have been shifted for clarity by the values indicated in the figure. The inset in the top right corner of the figure is a zoom of the V and R LCs up to 45 d. A colour version of this figure can be found in the online journal.

Table 1. Optical and NIR magnitudes of the minimum and secondary maximum of SN 2011fu, the corresponding times at which they occurred, and tail decline rates in the *BVRIZ* LCs.

Band	t_{\min}^a (d)	Apparent magnitude at minimum (mag)	t_{\max}^a (days)	Apparent magnitude at secondary maximum (mag)	Absolute magnitude at secondary maximum (mag)	Decline tail [mag (100d) ⁻¹]*
<i>U</i>	13.6 ± 1.1	17.90 ± 0.01	18.8 ± 0.8	17.29 ± 0.01	-17.58 ± 0.16	–
<i>B</i>	11.8 ± 1.0	17.96 ± 0.01	21.8 ± 1.6	17.52 ± 0.01	-17.26 ± 0.15	1.25 ± 0.07
<i>V</i>	10.6 ± 0.7	17.55 ± 0.01	22.9 ± 0.7	16.92 ± 0.01	-17.76 ± 0.15	1.78 ± 0.04
<i>R</i>	11.1 ± 2.1	17.18 ± 0.01	24.2 ± 1.2	16.53 ± 0.01	-18.08 ± 0.15	2.04 ± 0.04
<i>I</i>	11.4 ± 1.5	16.95 ± 0.01	24.1 ± 1.2	16.42 ± 0.01	-18.12 ± 0.15	1.97 ± 0.05
<i>z</i>	10.5 ± 0.7	17.29 ± 0.01	24.1 ± 1.0	16.29 ± 0.01	-18.22 ± 0.15	1.82 ± 0.04
<i>J</i>	–	–	24.5 ± 2.6	16.08 ± 0.01	-18.38 ± 0.15	–
<i>H</i>	–	–	29.6 ± 0.7	15.62 ± 0.01	-18.80 ± 0.15	–
<i>K_s</i>	–	–	29.2 ± 4.5	15.56 ± 0.01	-18.84 ± 0.15	–

^a t_{\min} and t_{\max} are calculated with respect to our adopted explosion date $\text{JD} = 2455824.5 \pm 0.7$. The errors in the NIR are large due to less photometric coverage of the maxima.

*Considering the interval starting ~ 40 days after explosion to ~ 160 days

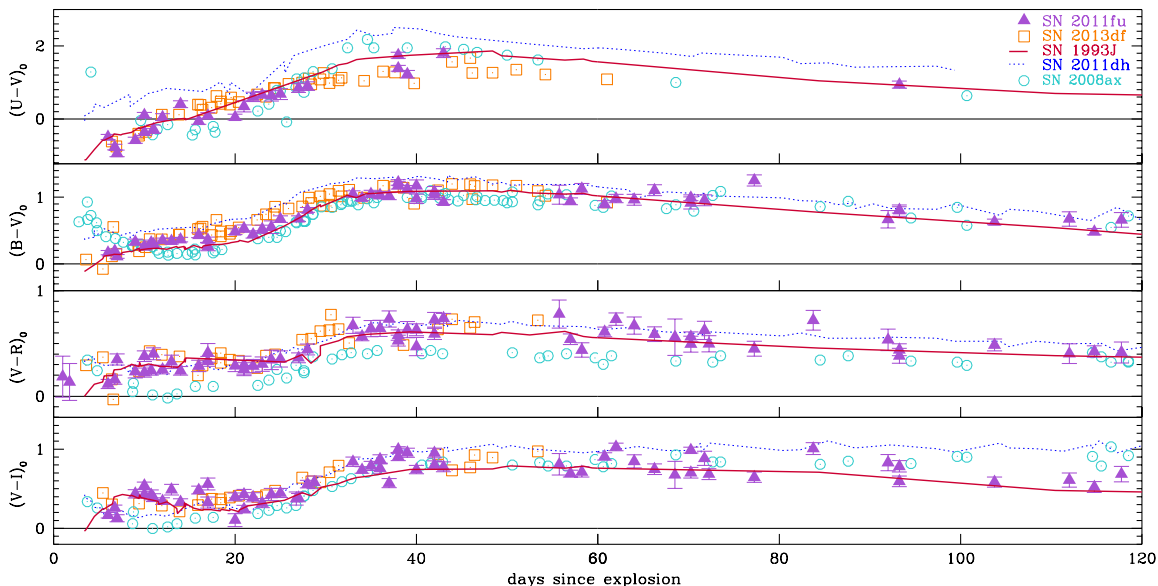


Figure 3. Comparison of the $(U - V)_0$, $(B - V)_0$, $(V - R)_0$, $(V - I)_0$ colours of type IIb SNe 1993J, 2008ax, 2011dh, 2011fu and 2013df. The colour of SN 2011fu has been corrected for the assumed reddening $E(B - V)_{\text{Total}} = 0.10^{+0.04}_{-0.03}$ mag, while the data and explosion epochs for the comparison supernovae were taken from the literature. A colour version of this figure can be found in the online journal.

when He I $\lambda 6678$ and He I $\lambda 7065$ appear and together with He I $\lambda 5876$ become progressively stronger. At 62 d a double-peaked emission associated with [O I] $\lambda\lambda 6300, 6364$ with components at approximately 6240 and 6300 \AA is observed, as well as an emission which has increased in intensity with respect to previous spectra, and is possibly associated with [O I] $\lambda 5577$. Starting at 103.8 d, [N II] $\lambda\lambda 6548, 6583$ produces an emission between ~ 6400 and ~ 6700 \AA (Jerkstrand et al. 2015). At phase 155.8 d we notice a decrease in the intensity of Na I around 5890 \AA (which is possibly contaminated by residual He I $\lambda 5876$; Jerkstrand et al. 2015). In addition we identify [Ca II] $\lambda\lambda 7291, 7324$ and two strong emission lines, a narrow one at approximately the rest position of H α (on top of N II $\lambda\lambda 6572, 6583$) and the other at ~ 6723 \AA . In the last spectrum of our sequence, other than [O I] $\lambda\lambda 6300, 6364$, [Ca II] $\lambda\lambda 7291, 7324$, O I $\lambda 7774$, we also distinguish Mg I $\lambda 4571$, [Fe II] $\lambda 7155$. The Ca II NIR line has diminished significantly from our previous spectra and is now blended with [C I] $\lambda 8727$. The narrow emissions at H α and the one at ~ 6700 \AA are still present although the first seems to have diminished in intensity while the second has become broader and now is centred at approximately 6735 \AA . The unresolved narrow H α emission line seen in our two last spectra, and also observed in some of the earlier-time spectra, is detected mainly on nights in which the seeing was not good or the SN was faint, so it is probably due to contamination from a nearby H II region.

Concerning the line at ~ 6700 \AA , we note that in the nebular spectra of SN 2013df (Morales-Garoffolo et al. 2014) there was also an emission line at a similar wavelength, and it was interpreted as a red-shifted component of the SN H α feature due to asymmetrical ejecta-circumstellar material interaction. However, since the line is detected in two different SNe it may be linked to a specific transition. In the models presented by Jerkstrand et al. (2015) there is an emission associated with [S II] $\lambda\lambda 6716, 6731$ about

an order of magnitude weaker than the observed line. However, uncertainties in the temperature and ionization state might cause it to be stronger for SN 2011fu. Another possibility is due to contamination by the nearby H II region which we believe is causing the narrow H α emission. In fact, in the 156 d spectrum both the narrow H α and the ~ 6700 \AA feature have similar FWHM and are unresolved suggesting a common origin not related to the SN. Similar narrow H α and [S II] $\lambda\lambda 6716, 6731$ emissions caused by nearby H II regions are detected e.g. in the spectra of the type IIb SNe 1987K (Filippenko et al. 1988) and 2008ax (Milisavljevic et al. 2010). However, in H II regions, [S II] $\lambda\lambda 6716, 6731$ lines are less intense than H α and this is not the case of the 304 d spectrum of SN 2011fu. In addition, the line at ~ 6700 \AA at this phase is clearly resolved, for these reasons, a [S II] $\lambda\lambda 6716, 6731$ emission associated with the SN is the favoured possibility at 304 d.

We have estimated the black-body temperatures of the SN by fitting a black-body function to the whole spectrum at each epoch. As can be seen in Figure 6, the black-body temperature evolves with a steep decline up to approximately 40 days since explosion and after that, it roughly remains constant. The time evolution of the temperatures is in agreement with the colour evolution of the SN and with that derived by Kumar et al. (2013). As seen in Section 4.1.1, the colours redden during approximately the first 40 days and after that they are practically constant with a slight trend towards the blue. In the first ~ 10 d, the trend of the black body temperature is opposite to SN 2011dh (Ergon et al. 2014). Note that during this phase the LC of SN 2011fu is in the adiabatically-powered declining phase to the minimum after the first peak while the pseudo bolometric LC of SN 2011dh is radioactively-powered.

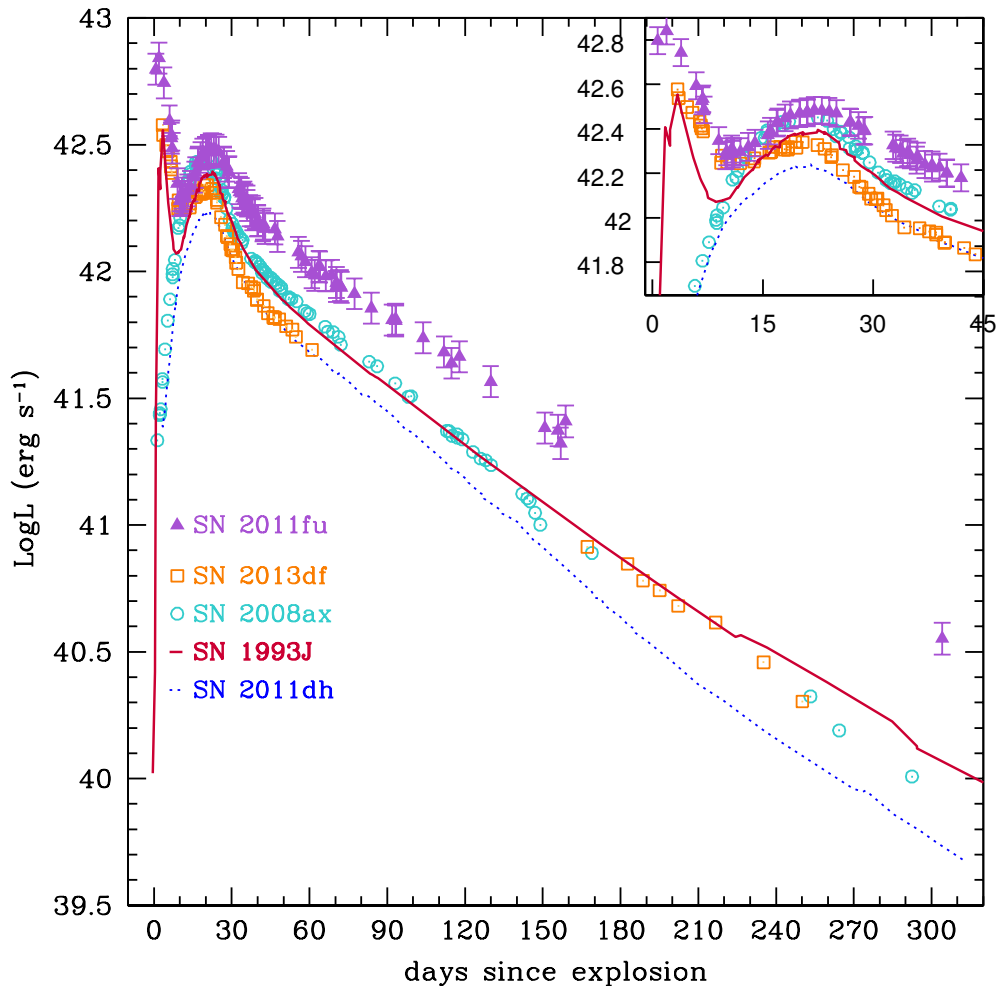


Figure 4. Pseudo bolometric optical-NIR LC of SN 2011fu compared to type IIb SNe 2013df, 1993J, 2008ax, and 2011dh. The data and explosion epochs for the comparison supernovae were taken from the literature. A colour version of this figure can be found in the online journal.

5.2 Velocities

From the minima of the P-Cygni absorptions of $H\alpha$ and $He\text{I } \lambda 5876$, we have an estimate of the velocities of the ejected material. Our results are presented in Figure 7 together with the velocities for SNe 1993J (Barbon et al. 1995; Taubenberger et al. 2011), 2008ax (Taubenberger et al. 2011), 2011dh (Ergon et al. 2014), and 2013df (Morales-Garoffolo et al. 2014). Note that the first two $He\text{I } \lambda 5876$ velocities reported in the figure were obtained from shallow, low contrast profiles, and in consequence are uncertain.

As we can see in Figure 7, both the $H\alpha$ and $He\text{I } \lambda 5876$ velocities we have derived are overall similar to those for the comparison SNe. Between 12 d and 28 d we note that the $H\alpha$ velocities for SN 2011fu are around 2000 km s^{-1} higher than those obtained with SYNOW by Kumar et al. (2013). From ~ 40 d on, the velocities remain constant and similar to those of SN 1993J but lower than the ones derived for SN 2011dh. We would like to point out that around a month past explosion the $He\text{I } \lambda 5876$ absorption component has a complex profile with a double trough, however, also

in these cases the values reported in the figure are those obtained by adjusting a single Gaussian to the whole profile. Thus, the velocities measured for $He\text{I } \lambda 5876$, may have an added uncertainty due to the possible contamination by $\text{Na I } \lambda\lambda 5890, 5896$, especially at later times. Noteworthy is the fact that for SN 2011dh, the hydrogen lines were never seen below 11000 km s^{-1} , whereas the helium lines were always constrained to lower velocities. In fact the modelling of the SN data suggested that 11000 km s^{-1} marked the interface between the H-rich envelope and the He core. For SN 2011fu, however, we cannot establish an analogous boundary based solely on the observed velocities, since in the velocity space $H\alpha$ is not separated from He.

The photospheric velocities are expected to be similar to those derived from the $Fe\text{II } \lambda 5169$ line (see e.g. figure 14 of Dessart & Hillier 2005), these are shown in the right panel of Figure 7. As can be seen in this figure the velocities of $Fe\text{II}$ for SN 2011fu are overall comparable with those of our selected sample of type IIb SNe. We notice again that up to

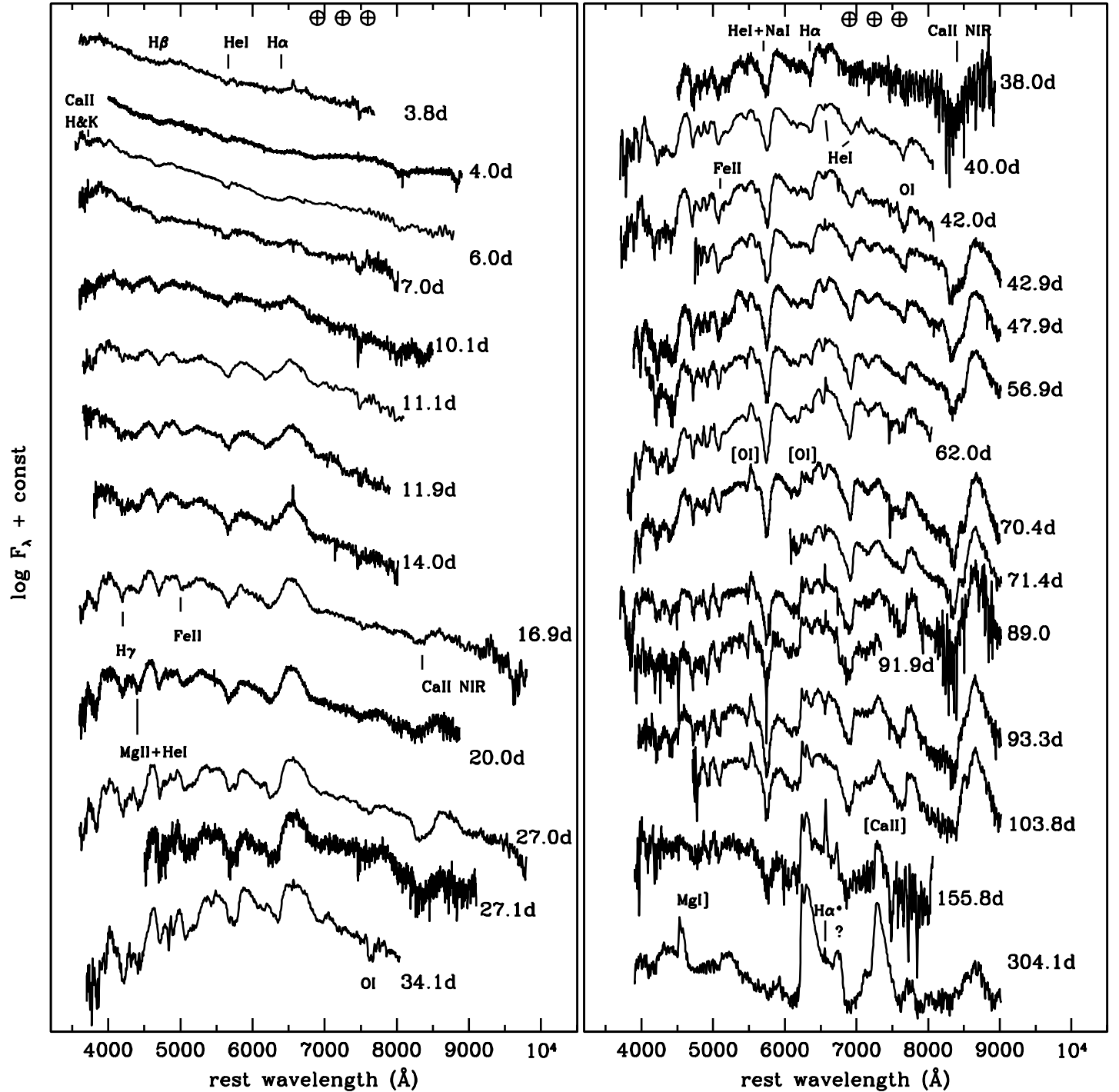


Figure 5. Optical spectral evolution of SN 2011fu, where the most relevant features in the spectra are indicated. The wavelengths at which there are residual telluric features have been marked with \oplus . The spectra have been corrected for the host galaxy redshift. Epochs indicated in the plot are with respect to our assumed explosion date of $\text{JD} = 2455824.5 \pm 0.7$. Spectra have been shifted vertically for clarity. $H\alpha^*$ is the narrow $H\alpha$ emission which we believe is associated to the H II region. The question mark at ~ 6700 \AA marks the line we discussed in Section 5.1.

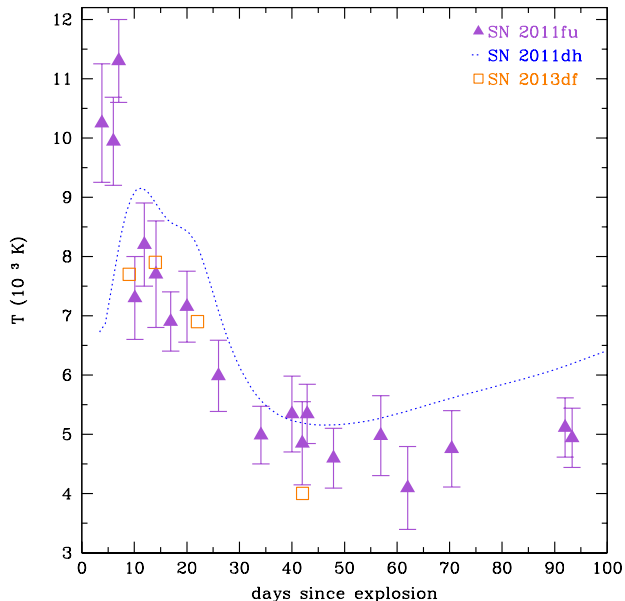


Figure 6. Black body temperature evolution of SN 2011fu with respect to our assumed explosion date, $JD = 2455824.5 \pm 0.7$, compared to those of SNe 2011dh and 2013df. A colour version of this article can be found in the online journal.

30 d they are lower than the ones derived by Kumar et al. (2013) by about 2000 km s^{-1} .

In summary, the expansion velocities of SN 2011fu are similar to those of other type I Ib SNe.

5.3 Comparison to other type I Ib SNe

In Figure 8 we present a comparison of spectra of SN 2011fu at three different phases in its evolution along with those of other type I Ib SNe downloaded from WISEREP⁵ (Yaron et al. 2012). As can be seen in the top panel of the figure, the spectra of SNe 1993J, 2011fu and 2013df exhibit a blue almost featureless continuum with shallow hydrogen and helium lines (note that the spectrum of SN 2013df corresponds to a later phase). The coeval spectra of SNe 2008ax and 2011dh show much stronger lines and redder continuum, reflecting the absence of the first peak of the LC. In the middle panel of Figure 8 we have depicted a spectral comparison at an intermediate phase (~ 40 d). Although there is an overall resemblance in the spectra, the intensity of the lines varies from one SN to the other, indicating differences in the temperature of their ejecta. Specifically, SN 2011fu has the weakest $H\alpha$ and $\text{He I } \lambda 5876$ absorptions at this phase. Although not shown in the figure, we have also compared a spectrum of SN 2011fu at 22 d with the other type I Ib SNe at similar phases. Again, we have found the intensity of the $H\alpha$ and $\text{He I } \lambda 5876$ lines to be lower than for the comparison SNe. For the nebular spectra (bottom panel of Figure 8), in the case of SN 2011fu, the $[\text{O I}] \lambda\lambda 6300, 6364$ emission profile is double-peaked, as it is for SN 2008ax. The $\text{Mg I } \lambda 4571$ profile for SN 2011fu is also

double-peaked and again more similar to that of SN 2008ax. Interestingly and as noted above, SN 2011fu shows an emission line at $\sim 6700 \text{ \AA}$ similarly to SN 2013df, although the line is broader in SN 2011fu.

5.4 Line profiles

In the left panel of Figure 9 we present the evolution of SN 2011fu’s $[\text{O I}] \lambda\lambda 6300, 6364$ nebular profile in velocity space. As seen in the figure, the profile shows two peaks, one at approximately 0 km s^{-1} and the other at $\sim -2800 \text{ km s}^{-1}$. In the middle panel of Figure 9 we show the nebular profiles of $[\text{O I}] \lambda 5577$ and $\text{Mg I } \lambda 4571$. Finally, in the right panel of the figure we have artificially added a component with $1/3$ the intensity of the original $[\text{O I}] \lambda 5577$ and $\text{Mg I } \lambda 4571$ profiles and redshifted by 3000 km s^{-1} , in order to compare them with the $[\text{O I}] \lambda\lambda 6300, 6364$ doublet. As can be seen in this last panel, the shapes of the oxygen and magnesium profiles are quite similar to one another. Taubenberger et al. (2009) already found a great similarity between the oxygen and magnesium profiles in nebular spectra of some stripped envelope SNe, supporting the idea that Mg and O have similar spatial distributions within their ejecta, which is expected from the models (Maeda et al. 2006). Moreover, in the case of the type I Ib SN 2011dh the components of the small scale fluctuations of the $[\text{O I}] \lambda\lambda 6300, 6364$ and $\text{Mg I } \lambda 4571$ in its late-time spectra coincide, which in addition to the results of Jerkstrand et al. (2015) imply that the oxygen lines are mostly formed in the O/Ne/Mg zone.

One possibility to explain the shape of the profiles of SN 2011fu is that the bulk of oxygen and magnesium is distributed in spherically symmetric expanding ejecta but that there is a clump with emission from both these elements expanding at $\sim 2800 \text{ km s}^{-1}$ towards the observer. A second possibility is that the radioactivity exciting the lines is distributed asymmetrically. Given the similarity of the $[\text{O I}] \lambda\lambda 6300, 6364$ and the $[\text{O I}] \lambda 5577$ and $\text{Mg I } \lambda 4571$ when an artificial component is added to the last two, an $H\alpha$ high-velocity absorption, as found by Maurer et al. (2010) for some type I Ib SNe, is not likely the principal cause of the double-peaked $[\text{O I}] \lambda\lambda 6300, 6364$ line, although it could be contributing to the profile. On the other hand, the narrow blue-shifted component of the $[\text{O I}] \lambda\lambda 6300$ carries a significant fraction of the flux even at 300 d. Residual opacity in the core or line blocking explains blue shifted emission lines of stripped envelope and specifically type I Ib SNe (Taubenberger et al. 2009; Jerkstrand et al. 2015). But in the case of SN 2011fu the persistence of the blue shifted emission from the $[\text{O I}] \lambda\lambda 6300, 6364$ line over time suggests some degree of asymmetry in the SN ejecta. So, in principle, either clumping or an asymmetrical distribution of the radioactive material seem the most likely explanations for the profiles. Intriguingly, we have found two other SNe in the literature that have blue shifted double-peaked $[\text{O I}] \lambda\lambda 6300, 6364$ profiles similar to SN 2011fu: Types I Ib SN 1996aq (Taubenberger et al. 2009) and SN 2005bf (Modjaz et al. 2008; Milisavljevic et al. 2010) at 216d. For SN 1996aq Taubenberger et al. (2009) speculated that the most likely explanation for its profile shape is that a clump is moving in the line of sight at high velocity, and this is probably also the case for SNe 2011fu and 2005bf.

⁵ <http://wiserep.weizmann.ac.il/spectra/list>

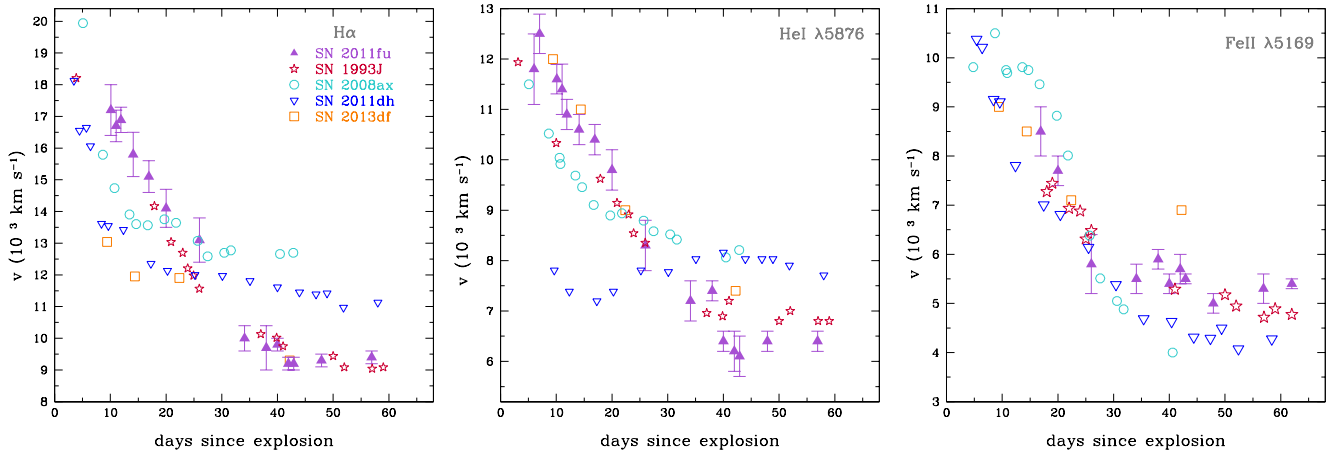


Figure 7. Velocity evolution of H α , He I λ 5876, and Fe II λ 5169 for SN 2011fu compared to those of SNe 1993J, 2008ax, 2011dh, and 2013df.

Finally, we note that the shape of SN 2011fu’s [Ca II] λ 7291, 7324 line (not shown in the figure) is not similar to the oxygen and magnesium profiles, implying that, as expected, they are forming at different locations within the ejecta.

6 DISCUSSION

6.1 Hydrodynamical modelling of the pseudo-bolometric LC

One approach to derive physical parameters of SNe is to compare LCs and expansion velocities derived from hydrodynamical models with observations. From this comparison it is possible to constrain the explosion energy, the ejecta and the nickel masses as well as the distribution of the radioactive material. These parameters can be estimated by focusing the modelling around the “main peak” of the LC (i.e. the nickel powered phase). Information on the size of the progenitor is given by the post shock breakout cooling phase.

We calculated a set of hydrodynamical models using as initial structures those derived from stellar evolutionary calculations. A one-dimensional Lagrangian LTE radiation hydrodynamics code (Bersten et al. 2011) is used to explode the initial configuration. The code allows to calculate consistently the shock wave propagation in the stellar interior, the shock breakout and the posterior phases of the LC evolution until the SN becomes nebular. An update of the micro physics (equation of state and opacities) appropriate for the study of stripped-envelope SNe were incorporated in the code by Bersten et al. (2012). As pre-SN structure we adopted helium (He) stars with a thin hydrogen envelope ($< 1 M_{\odot}$) which have successfully reproduced the LCs and the spectral features of other SNe I Ib (see e.g. Shigeyama et al. 1994, Blinnikov et al. 1998, and more recently Bersten et al. 2012). The He core models used here were calculated by Nomoto & Hashimoto (1988) following the stellar evolution until the collapse of the core. We have smoothly attached a low-mass hydrogen-rich envelope in hydrostatic and

thermal equilibrium to the He core to take into account the thin H-envelope required for a SN I Ib.

First we focus our analysis on the modelling of the LC and the photospheric velocities around the secondary peak without taking into account the cooling part of the LC. Figure 10 shows the results of the LC modelling for three different He core masses, $4 M_{\odot}$ (He4), $5 M_{\odot}$ (He5), and $8 M_{\odot}$ (He8), which correspond to the stellar evolution of single stars with main-sequence masses of 15, 18, and $25 M_{\odot}$, respectively. For each model different explosion energies and ^{56}Ni masses were explored. Here we show the best set of parameters for each of the models. Specifically, an explosion energy of 1, 1.3 and 2 foe ($1 \text{ foe} = 1 \times 10^{51} \text{ erg}$) and a ^{56}Ni mass of $0.15 M_{\odot}$ ⁶ were found for model He4 (He4E1Ni15), He5 (He5E13Ni15) and He8 (He8E2Ni15) respectively. As can be seen in the figure, model He5 provides the best representation of the observed data. The He8 model is too massive to reproduce the LC unless we assume a more energetic explosion but this would fail to fit the expansion velocities. On the other hand a lower mass model, He4, gives a worse fit of the LC and underestimates the early photospheric velocities.

Our favoured model, He5, is a H-free object with an initial radius of $\approx 2 R_{\odot}$. Although, this model gives a good representation of the data around the secondary peak, it fails to reproduce the cooling phase due to the compact structure of the progenitor. Figure 11 shows the fits to the LC for models with different progenitor radius. The size of the progenitor was modified by attaching thin H-rich envelopes to the core of the He5 model (solid line). Note that the presence of a thin envelope mostly affects the shape of the LC during the cooling phase. Figure 11 shows models with radii of 400, 450, and $500 R_{\odot}$ and envelope masses (M_{env}) of $\approx 0.3 M_{\odot}$. The model with $R = 450 R_{\odot}$ provides the closest match to the data. However, all the models give an initial peak brighter than the observations. The differences

⁶ We have estimated the uncertainty of the ^{56}Ni mass by considering only the photometric errors to be $0.02 M_{\odot}$

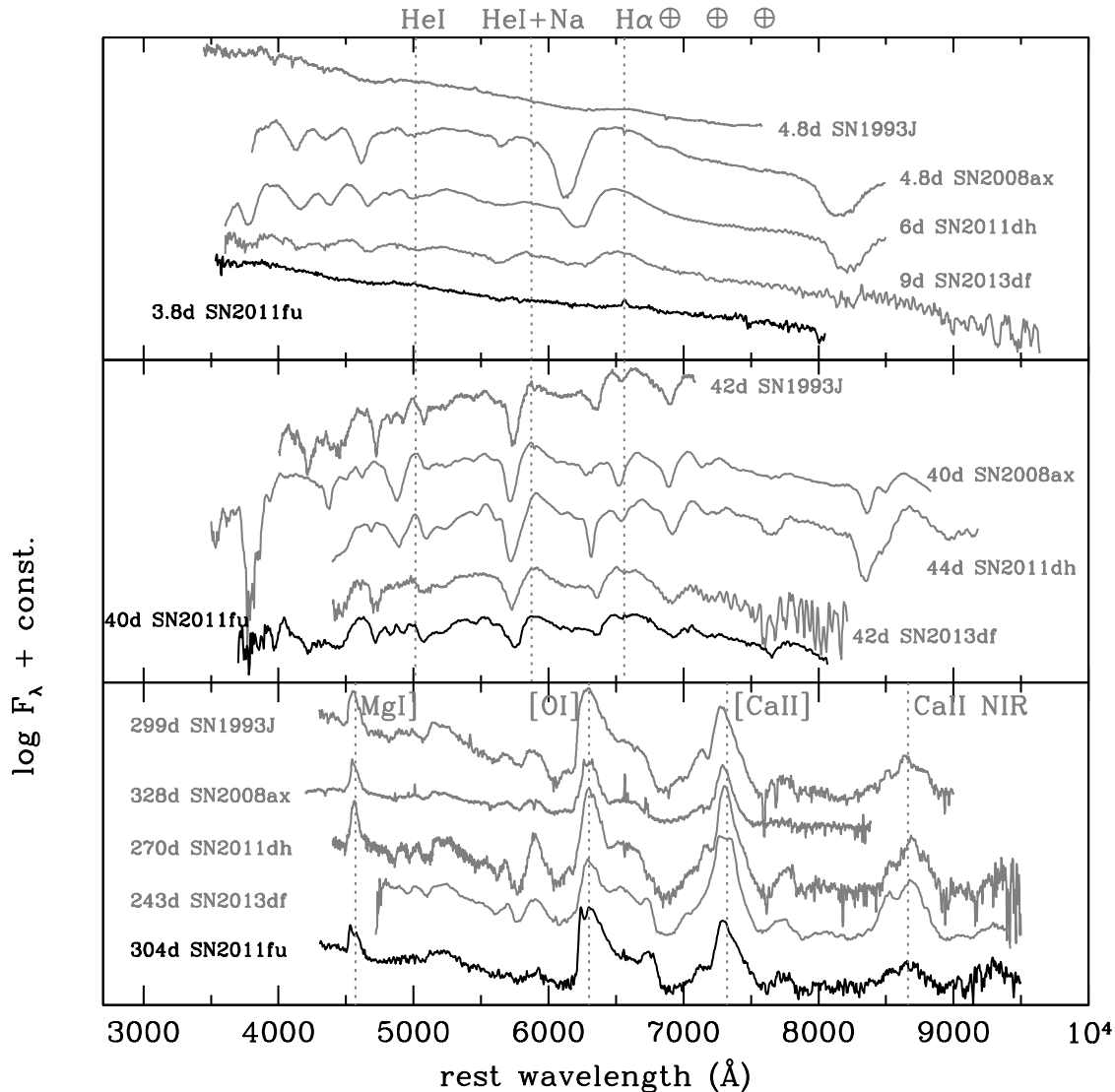


Figure 8. Comparison of early (5–9 days since explosion), intermediate (40–44 days after explosion), and late (299–398 days after explosion) spectra of SN 2011fu with those of type IIb SNe 1993J, 2008ax, 2011dh, and 2013df. The original references for the spectra of the comparison SNe are: Barbon et al. 1995 for SN 1993J at 4.8 and 42 d, and an unpublished spectrum taken at the 1.82-m telescope in Asiago (Italy); Taubenberger et al. 2011 for SN 2008ax at 4.8 and 40 d, Modjaz et al. 2014 for SN 2008ax at 328 d; Ergon et al. 2014 for SN 2011dh at 6 and 44 d, Shivvers et al. 2013 for SN 2011dh at 270 d, and Morales-Garoffolo et al. 2014 for SN 2013df’s spectra. The spectra have been dereddened, redshift corrected and shifted vertically for clarity.

may be due to uncertainties in the bolometric calculation (note that we have no UV data for the SN, which are important at these early phases) or/and limitations of the model such as the LTE approximation or variation of the density profile of the outermost layers, among other. It is noteworthy that to reproduce the early phase of the LC it was not only needed to assume an extended envelope but it was also necessary to adopt a slightly more massive envelope than in previous modelling of SNe IIb, where e.g. $M_{\text{env}} = 0.1 M_\odot$ was required for SN 2011dh and SN 1993J (Bersten et al. 2012; see also Nakar & Piro 2014). The need for a more massive envelope was mainly due to the high luminosity of the minimum after primary maximum of the LC. For lower M_{env} the cooling occurs faster and the

minimum occurs at a lower luminosity even for large radius. A summary of all the parameters obtained from the best core and envelope model that fit the observed LC is shown in Table 2. Similar to Kumar et al. (2013) we have obtained a kinetic energy which is relatively higher than that obtained for other type IIb SNe but still lower than those derived for SN 2011ei (Milisavljevic et al. 2013) and the type IIb hypernova SN 2003bg (Hamuy et al. 2009). The comparison to the pseudo-bolometric LCs of other type IIb SNe in Section 4.2 indicates that SN 2011fu synthesised more ^{56}Ni in its explosion than SNe 1993J, 2008ax, 2011dh and 2013df (cf table 7 of Morales-Garoffolo et al. 2014). The modelling of the LC corroborates this result. Note that we obtain a lower ^{56}Ni mass than Kumar et al.

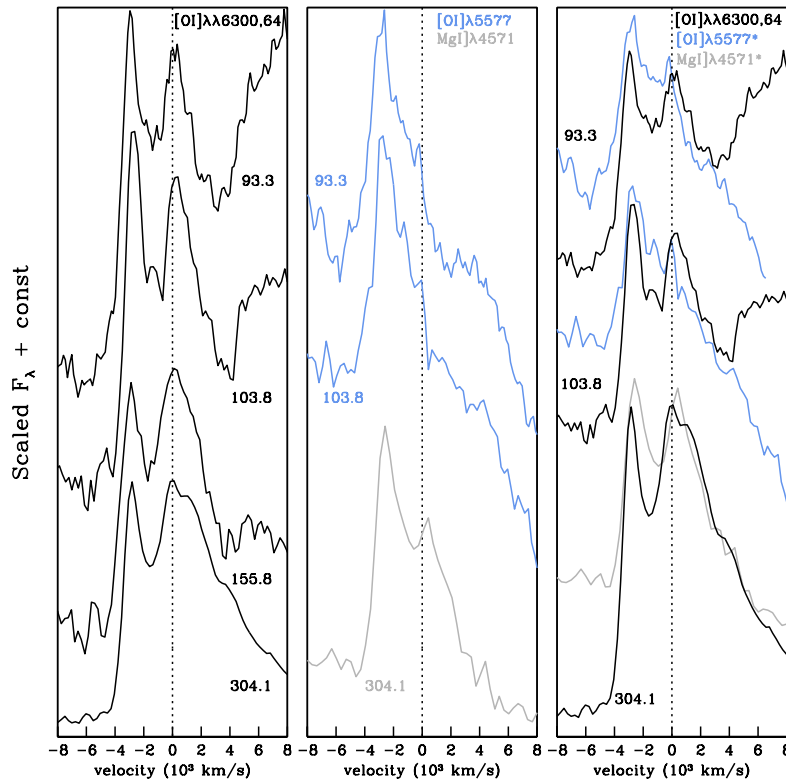


Figure 9. Left panel: evolution of the late-time profiles of [O I] $\lambda\lambda 6300, 6364$ in velocity space. The 0 km s^{-1} vertical line corresponds to 6300 \AA . Middle panel: [O I] $\lambda 5577$ and Mg I $\lambda 4571$ profiles between 93 and 304 d since explosion. Right panel: Comparison of the [O I] $\lambda\lambda 6300, 6364$ profiles with [O I] $\lambda 5577^*$ and Mg I $\lambda 4571^*$, which are the original profiles plus an artificial component, scaled to $1/3$ the intensity and redshifted by 3000 km s^{-1} , added to account for the doublet nature of [O I] $\lambda\lambda 6300, 6364$.

(2013) ($0.21 M_{\odot}$) possibly due to the different distance and extinction to SN 2011fu that we have adopted while our estimate of the ejected mass is larger than theirs ($1.1 M_{\odot}$).

Concerning the progenitor radius, core mass, and hydrogen envelope, we have obtained overall larger values than those obtained by Kumar et al. (2013) with analytical models. Specifically they obtained a progenitor radius of $\sim 150 R_{\odot}$, a He core mass of $1 M_{\odot}$, and a hydrogen envelope mass of $0.1 M_{\odot}$. But as they noted, their results should be considered only order of magnitude estimates. With hydrodynamical modelling, we obtained a radius that is 3 times larger and consistent with that of an extended supergiant similarly to the progenitors of SNe 1993J ($\sim 600 R_{\odot}$, Maund et al. 2004; Van Dyk et al. 2013) and 2013df ($\sim 550 R_{\odot}$, Van Dyk et al. 2014). All in all, our calculations show that the progenitor of SN 2011fu was not a Wolf Rayet (WR) star but a supergiant.

6.2 Comparison with late time spectral models

With the objective of better understanding the nature of the progenitor of SN 2011fu and its ejecta, we have compared the last three spectra of our sequence (104, 156, 304 d) with the late time spectral models presented in Jerkstrand et al. (2015).

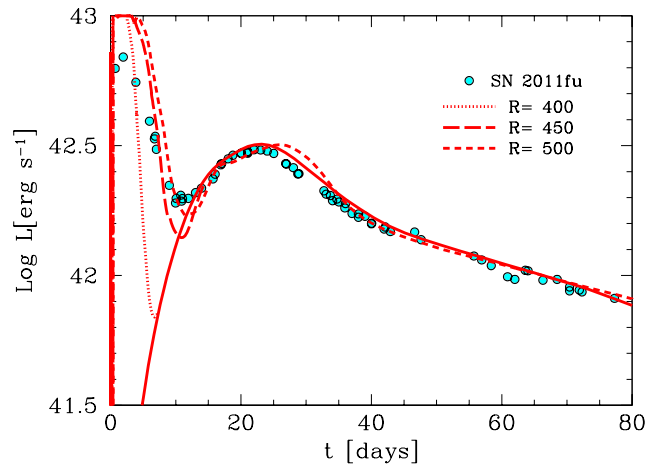


Figure 11. Observed bolometric LC of SN 2011fu (dots) compared with the bolometric LCs for models with the same physical parameters than our preferred model (He5E13Ni15; solid line), but different initial radii. The radius variation is accomplished by attaching thin H-rich envelopes to the He5 model. An extended progenitor with $R \approx 450 R_{\odot}$ (long dashed) is needed to reproduce the cooling phase of SN 2011fu.

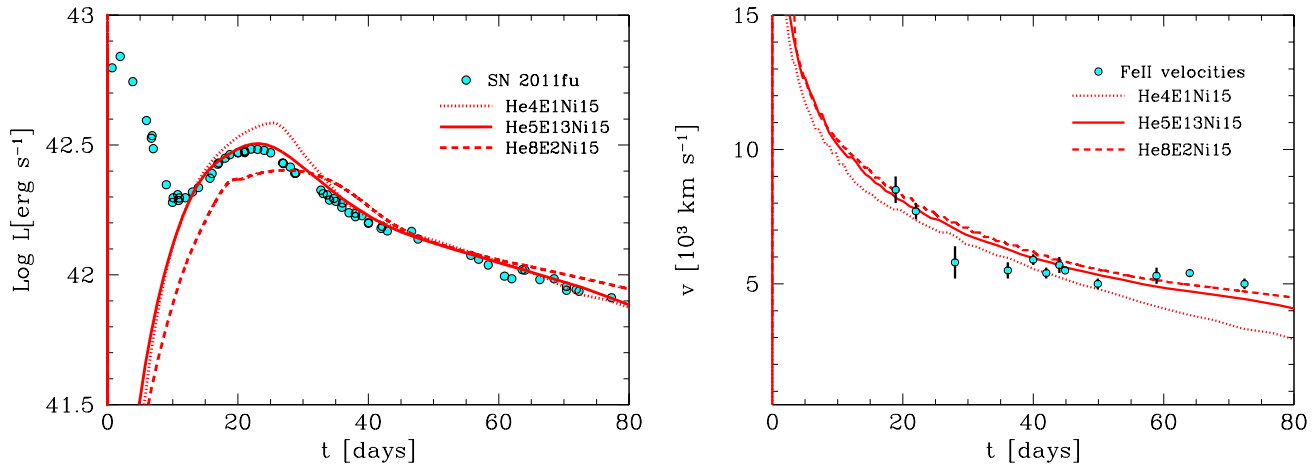


Figure 10. Left panel: Observed bolometric light curve of SN 2011fu (dots) compared with the results of the light curve calculations for models He4E1Ni15 (dotted line), He5E13Ni15 (solid line) and He8E2Ni15 (dashed line) omitting the first peak. Right panel: Evolution of the photospheric velocity for models He4E1Ni15 (dotted line), He5E13Ni15 (solid line) and He8E2Ni15 (dashed line) compared with measured Fe II line velocities of SN 2011fu (dots).

Table 2. Explosion parameters and progenitor properties for the best fitting models to the observed data of SN 2011fu.

Parameters	
$E_{\text{kin}} (10^{51} \text{erg})$	1.3
$^{56}\text{Ni mass} (M_{\odot})$	0.15
$M_{\text{ej}} (M_{\odot})$	3.5*
Progenitor Radius (R_{\odot})	450
Initial mass (M_{\odot})	18
Hydrogen envelope mass (M_{\odot})	0.3

*Assuming that $1.5 M_{\odot}$ collapsed to form a compact remnant and the rest of the mass that formed the core was ejected.

To begin with we compared our spectra with the three models with the same (best-fit to SN 2011dh) settings for mixing, clumping, molecule, and dust, and differing only in progenitor mass (12, 13 and $17 M_{\odot}$). These are models 12C, 13G, and 17A (see table 4 in Jerkstrand et al. 2015). In order to do this, we scaled the models accounting for differences in ^{56}Ni mass synthesised in the explosion, phase (if the difference in phase between models and spectra is Δt , the factor $\exp(-2\Delta t/111.4)$ is used to scale the models), and distances assumed for the models and the SN. The models have been calculated at a distance of 7.8 Mpc and assuming a ^{56}Ni mass of $0.075 M_{\odot}$ (since they were constructed to compare with SN 2011dh), while in Section 6.1 we have estimated $0.15 \pm 0.02 M_{\odot}$ of ^{56}Ni from SN 2011fu. Overall the model spectra are dimmer than the SN spectra. We note that for ^{56}Ni on the higher end of its uncertainty the flux difference between modelled and observed spectra diminishes. Given that the errors in the extinction are not large, it is likely not the cause of the discrepancy between the flux levels. Another possibility is that the ejecta structure of SN 2011fu is different than the one used in the models and the IIB SNe fit with the models. Note that the Jerkstrand et al. (2015) models have an imposed dust extinction of $\tau=0.25$ from 200 d, if no dust is produced in SN 2011fu this introduces a 25% flux error for the last spectrum.

At all phases the overall best fitting model is 13G. The

oxygen lines produced by the 17A model are more intense than those observed, while the 12C model underestimates the oxygen intensities. In Figure 12 we present the three late-time SN spectra compared to the 13G spectral models at coeval phase. We have also indicated in the plot some of the most important features that arise in the late spectral modelling⁷. In the middle panel of the figure we have also plotted model 17A and 12C to reflect the discrepancy between the oxygen lines of the model and those observed of the spectrum. In Figure 13 we present a comparison of the [O I] $\lambda\lambda$ 6300, 6364 late-time luminosities of SN 2011fu and those corresponding to models 12C, 13G and 17A. All in all, these comparisons indicate that the progenitor of SN 2011fu was likely at the lower end of the tested range, with $M_{\text{ZAMS}} \sim 13 M_{\odot}$. The model that best fits the observed LC, however, corresponds to a progenitor star with a $5 M_{\odot}$ He core, i.e. $M_{\text{ZAMS}} = 18 M_{\odot}$. This value is higher but not inconsistent with the estimate we have obtained here. In fact, it is not surprising to have found different values for the progenitor’s initial mass since the initial conditions and the methodology used in both modelling methods are rather different. The hydrodynamical modelling of the LC

⁷ $\text{H}\alpha^*$ is the narrow $\text{H}\alpha$ emission which we believe is associated with an H II region. The question mark at ~ 6700 marks the line we discussed in Section 5.1.

uses stellar evolutionary models with different progenitor M_{ZAMS} and explodes these structures for different values of explosion energies and nickel masses to compare with the data of SN 2011fu. The greatest degeneracy in the model is between mass and energy, this is the reason for modelling both the LC and the expansion velocities. Another source of uncertainty is introduced by the initial stellar evolutionary models. The code treats in a very approximate way the radiation transportation but correctly simulates the explosion dynamics.

The nebular spectral models use ejecta with nucleosynthesis from the evolution and explosion of stars of different M_{ZAMS} . While the dynamic structure of the core is manually arranged to match the observed metal line widths in Type IIb SNe ($\sim 3500 \text{ km s}^{-1}$) as well as to capture effects of macroscopic mixing and clumping seen in 2D simulations and inferred from observed line profiles, the envelope follows a profile obtained in a 1D explosion simulation by Bersten et al. (2012). Uncertainty is introduced by two sources; the dynamic structure of the ejecta, and the margin of error in the spectral modelling. Since the model reproduces well the metal line profiles of SN 2011fu, the core velocity of 3500 km s^{-1} is likely accurate. The envelope absorbs quite little of the gamma-ray energy (see Appendix A in Jerkstrand et al. 2015) so an uncertainty in its density profile introduces only a moderate dispersion for the flux levels. Thus, the model uncertainty is likely dominated by the error in the ^{56}Ni mixing as well as molecule and dust formation. The ^{56}Ni mixing used in 13G is chosen to match both the diffusion phase LC and the nebular spectra of SN 2011dh, but may be somewhat different in SN 2011fu.

It would be interesting to analyse whether the differences between the two modelling approaches persist if the post-explosion density structure of our preferred LC model for SN 2011fu is used as initial condition of the spectral modelling. However, this analysis is beyond the scope of this work. In any case, the mass estimates obtained from the two methods are consistent in the sense that we can discard the possibility of the progenitor being a single Wolf-Rayet star (with $M_{\text{ZAMS}} \geq 25$).

To check for evidence of molecule formation in the ejecta of the SN, we compared models 12C (no molecular cooling; see table 4 of Jerkstrand et al. 2015) and 12D (with molecular cooling) with our late-time spectra. In the bottom panel of Figure 12 the comparison at 300 d is shown. The oxygen lines are the ones that are most sensitive to molecule formation at 100 and 150 days along with [C I] $\lambda 8727$ at 300 d. In our case, the model without molecular cooling is favoured, similarly to the case of SN 2011dh. We also looked at the contrast factor or oxygen zone density, i.e. the density ratio between the metal zones in the core and the Fe/Co/He zone, but the models with different oxygen zone densities do not show palpable differences. However, we do note that the Mg I] $\lambda 4571$ line at 304 d is quite strong, which favours a high oxygen zone density, since the line is quite sensitive to this parameter, and this is in agreement with the results found for SN 2011dh in Jerkstrand et al. (2015).

In summary, the comparisons of our nebular spectra with the models presented in Jerkstrand et al. (2015) point to a progenitor with an initial mass of $\sim 13M_{\odot}$, likely no molecules were formed in the ejecta, and the oxygen zone density was possibly high.

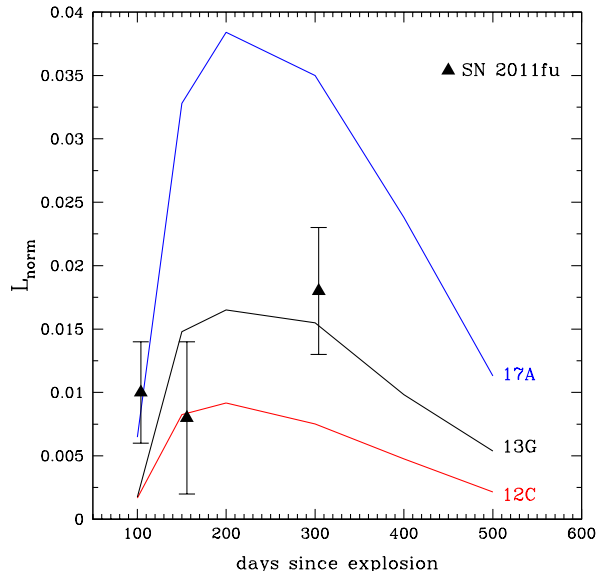


Figure 13. Late-time [O I] $\lambda\lambda 6300, 6364$ luminosities of SN 2011fu (triangles) compared to model tracks 12C, 13G, and 17A of Jerkstrand et al. 2015.

7 CONCLUSIONS

In this paper we have analysed optical and NIR data for the double-peaked type IIb SN 2011fu spanning from a few days to approximately 300 days after explosion. The double-peaked LC is brighter than those of other type IIb SNe such as SNe 1993J, 2008ax, 2011dh, and 2013df. It also presents a longer cooling phase after primary peak than SNe 1993J and 2013df.

SN 2011fu's spectra evolve in a similar fashion to those of SN 1993J. During the first phases after explosion Balmer, He, Ca, and Fe lines are present in the spectra. At around 40 d, He lines dominate the spectra although $H\alpha$ is still present and persists up to later phases. Forbidden oxygen lines are clearly visible at ~ 62 d, which indicates that the SN is entering the nebular phase. We derived $H\alpha$, He I $\lambda 5876$, and Fe II $\lambda 5169$ velocities from the absorption minima of the P-Cygni profiles in SN 2011fu's spectra, which are consistent with the results obtained for other type IIb SNe. At late times the most important features are [O I] $\lambda\lambda 6300, 6364$, and [Ca II] $\lambda\lambda 7291, 7324$. In addition in our latest spectrum taken at 304 d, Mg I] $\lambda 4571$ and a line that could be due to [S II] $\lambda\lambda 6713, 6730$ are detected.

The nebular profiles of oxygen and magnesium indicate that the oxygen/magnesium zone is clumped. Calcium lines do not show similar profiles indicating, as expected, that they are formed at a different location.

Thanks to the hydrodynamical modelling of the pseudo-bolometric LC we conclude that SN 2011fu was the explosion of an extended object ($R \approx 450 R_{\odot}$), with a He core mass of $5 M_{\odot}$ ($M_{\text{ZAMS}} \approx 18M_{\odot}$) and an ejecta mass of $3.5 M_{\odot}$, an explosion energy of 1.3×10^{51} erg, and a ^{56}Ni mass of $0.15 M_{\odot}$. In comparison with other SNe IIb, this object seems to be a bit more massive, more energetic and to have produced more ^{56}Ni . The stellar envelope of its progenitor also seems to have been more massive than for other

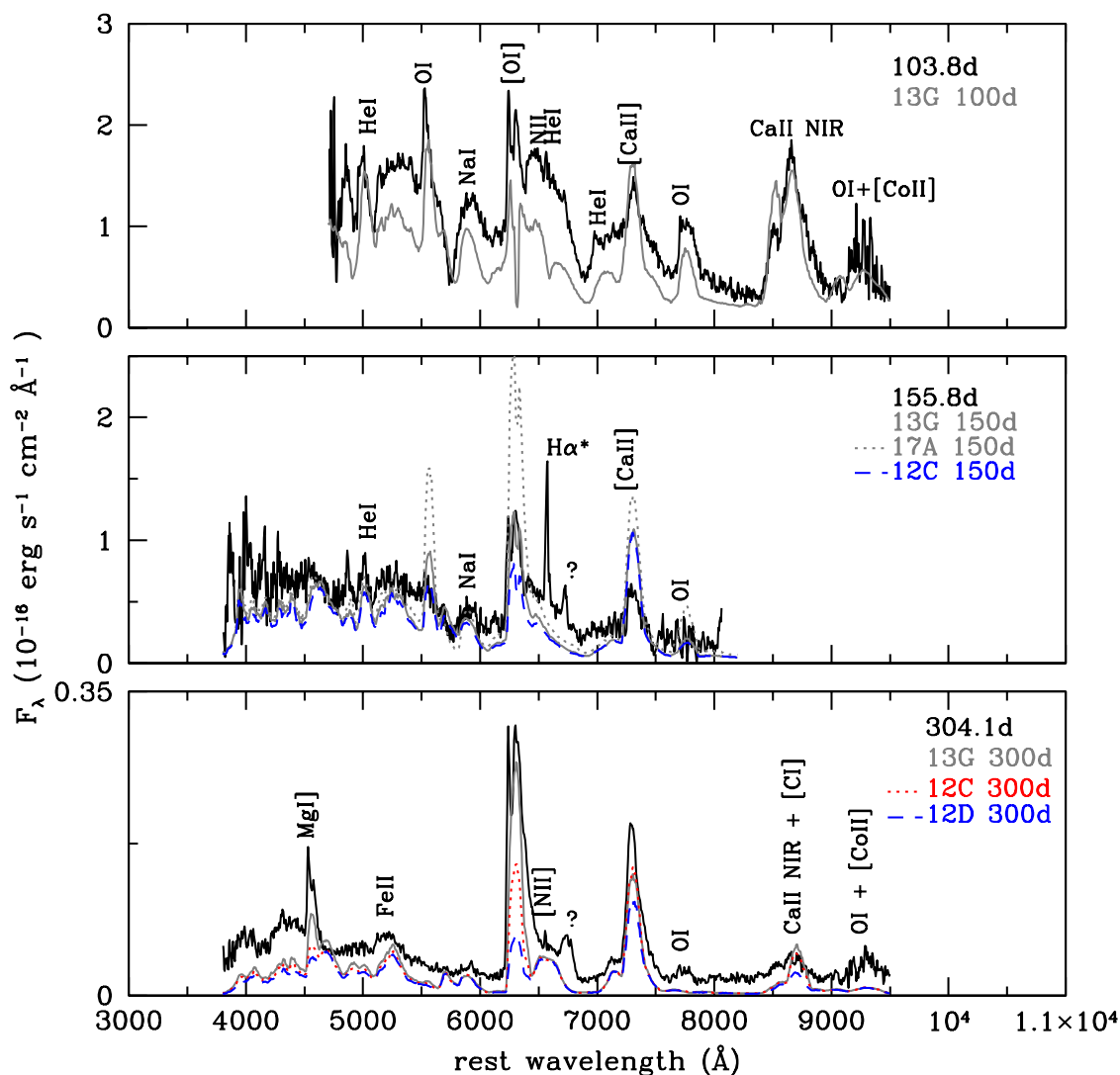


Figure 12. Late-time spectra of SN 2011fu compared to the best fitting spectral model 13G presented in Jerkstrand et al. 2015. In the middle panel we show our spectrum at ~ 156 d and the three models that vary only in progenitor mass. In the bottom panel other than model 13G we show models 12C and 12D, which are the ones used to study molecule formation in the SN ejecta. The spectral models have been scaled accounting for difference in phase with respect to the SN spectra (Δt) by the factor $\exp(-2\Delta t/111.4)$, difference in ^{56}Ni synthesised mass, and distance assumed for the models and the SN.

type IIb SNe ($\sim 0.3 M_{\odot}$), implying the presence of more H in the SN ejecta.

We have compared our late-time spectra (≥ 100 d) with late-time spectral models for type IIb SNe presented in Jerkstrand et al. (2015). From this analysis, the overall best fitting value of the progenitor mass is $13 M_{\odot}$, similar to the M_{ZAMS} values that have previously been derived for SNe 1993J, 2008ax, and 2011dh ($12\text{--}16 M_{\text{ZAMS}}$, Jerkstrand et al. 2015). As for these other IIb SNe, we find also for SN 2011fu that molecule formation is non-existent.

The extensive data-set of SN 2011fu presented in this paper have permitted us to derive some of the characteristics of the ejecta of one more type IIb SN presenting a double-peaked LC, as well as to set constraints on some of

its progenitor’s properties. SN 2011fu is another example of a type IIb SN which shows evidence that the progenitor mass was in the range $13\text{--}18 M_{\odot}$. This is in the higher range of stars that are seen to explode (Smartt 2015), which supports type IIb progenitor detections all being significantly more luminous than the II-P progenitors. Furthermore, as for all previously modelled IIb SNe, we determine a M_{ZAMS} value much lower than what would have been needed for a wind-stripped WR progenitor, and evidence is piling up that most, if not all, type IIb SNe come from binary stripped progenitors.

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APPENDIX A: INSTRUMENTAL SET-UP

Table A1. Characteristics of the instruments used to acquire the photometric data of SN 2011fu.

Telescope	Instrument	Field of View	Pixel Scale ($'' \text{ pix}^{-1}$)	Filters	Key
2-m Liverpool	RatCam	4'.6 x 4'.6	0.140	uBVriz	LT
1.82-m Asiago	AFOSC	8'.1 x 8'.1	0.473	UBVri	A1p82
2.2-m Calar Alto	CAFOS	9' x 9'	0.530	UBVRI	Ca2p2
2.5-m Northern Optical	ALFOSC	6'.4 x 6'.4	0.190	UBVRI	NOT
4.2-m William Herschel	ACAM	9'.0 x 10'.4	0.250	griz	WHT
2-m Faulkes North	fs02	10'.5 x 10'.5	0.304	BVRI	FTN1
2-m Faulkes North	EM01	4'.7 x 4'.7	0.278	BVRI	FTN2
0.6-m Esteve Duran	ST-9XE	11'.7 x 11'.7	1.370	VI	TED
0.4-m Cantabria	ST-8XE	24'.0 x 26'.0	0.900	VRI	OAC
0.36-m Celestron C14	QHY9 CCD	25'.0 x 19'.0	0.450	unfiltered	Xing Gao
0.5-m Newtonian	FLI Proline CCD	20'.4 x 19'.8	2.320	unfiltered	F.Ciabattari/E.Mazzoni
0.92-m Asiago Schmidt	SCAM	48'.7 x 48'.7	1.437	BVRI	ASch
10.4-m Gran Telescopio Canarias	OSIRIS	7'.8 x 8'.5	0.127	gri	GTC
1.52-m Carlos Sánchez	CAIN3	4'.2 x 4'.2	1.000	JHK _s	TCS

Table A2. List of spectroscopic observations of SN 2011fu.

Date	Julian date (+2400000.00)	Phase ^a (d)	Telescope+Instrument	Grism+Filter	Range (Å)	Resolving Power ^b
2011/09/23	55828.3	3.8	AS1p82+AFOSC	g4	3500–8450	300
2011/09/24	55828.5	4.0	WHT+ISIS	R158R+R300B+gg495	3200–10000	1000,1300
2011/09/25	55830.4	5.9	AS1p82+AFOSC	g2+g4	3500–10200	200,300
2011/09/26	55831.3	6.8	AS1p82+AFOSC	g4	3500–8450	300
2011/09/29	55833.6	9.1	NOT+ALFOSC	g4	3200–9100	500
2011/09/30	55835.3	10.8	AS1p82+AFOSC	g2+g4	3500–10200	200,300
2011/10/01	55836.4	11.9	AS1p82+AFOSC	g4	3500–8450	300
2011/10/03	55838.5	14.0	AS1p82+AFOSC	g4	3500–8450	300
2011/10/06	55841.4	16.9	Ca2p2+CAFOS	B200+R200	3200–11000	500,500
2011/10/10	55844.6	20.1	NOT+ALFOSC	g4	3200–9100	500
2011/10/17	55851.5	27.0	Ca2p2+CAFOS	B200+R200	3200–11000	500,500
2011/10/17	55851.6	27.1	WHT+ACAM	V400	3500–9400	500
2011/10/24	55858.5	34.0	AS1p82+AFOSC	g4	3500–8450	300
2011/10/27	55862.4	37.9	NOT+ALFOSC	g4	3200–9100	500
2011/10/30	55864.6	40.1	AS1p82+AFOSC	g4	3500–8450	300
2011/10/31	55866.4	41.9	AS1p82+AFOSC	g4	3500–8450	300
2011/11/01	55867.4	42.9	Ca2p2+CAFOS	G200	3700–9500	500
2011/11/07	55872.5	48.0	WHT+ISIS	R158R+R300B+gg495	3200–10000	1000,1300
2011/11/15	55881.4	56.9	Ca2p2+CAFOS	G200	3700–9500	500
2011/11/20	55886.4	61.9	AS1p82+AFOSC	g4	3500–8450	300
2011/11/29	55894.5	70.0	NOT+ALFOSC	g4	3200–9100	500
2011/11/29	55895.4	70.9	Ca2p2+CAFOS	R200	6300–11000	500
2011/12/17	55913.5	89.0	TNG+DOLORES	LRR+LRB	3000–10000	460,460
2011/12/20	55916.4	91.9	AS1p82+AFOSC	g4	3500–8450	300
2011/12/21	55917.4	92.9	NOT+ALFOSC	g4+g5	3200–10250	500,500
2012/01/01	55928.4	103.9	Ca2p2+CAFOS	G200+gg495	4800–9500	500
2012/02/22	55980.3	155.8	AS1p82+AFOSC	g4	3500–8450	300
2012/07/20	56128.6	304.1	GTC+OSIRIS	R300B+R300R	3600–10000	430

^a Phase with respect to our adopted explosion epoch $\text{JD} = 2455824.5 \pm 0.7$.^b The resolution was measured from the full width at half maximum of the night sky lines. The resolving power values reported here are for a reference wavelength of 6500 Å.**APPENDIX B: SN 2011FU'S PHOTOMETRIC DATA**

Table B1. Optical Bessell photometry of SN 2011fu. The U data are not S-corrected, note that the reported RI LT data are the S-corrected Sloan ri magnitudes.

Date	JD (+2400000.00)	Phase ^a (d)	U (mag)	B (mag)	V (mag)	R (mag)	I (mag)	Key
2011/09/20	55825.2	0.7	—	—	16.32 ± 0.14	—	—	Xing Gao
2011/09/21	55825.5	1.0	—	—	—	16.07 ± 0.13	—	F. Ciabattari/E.Mazzoni
2011/09/21	55826.4	1.9	—	—	16.21 ± 0.15	—	—	Xing Gao
2011/09/21	55826.4	1.9	—	—	—	16.01 ± 0.08	—	F. Ciabattari/E.Mazzoni
2011/09/23	55828.3	3.8	—	—	—	16.27 ± 0.02	—	A1p82
2011/09/23	55828.4	3.9	—	—	16.45 ± 0.24	—	—	Xing Gao
2011/09/25	55830.5	6.0	16.52 ± 0.02	17.10 ± 0.02	16.83 ± 0.01	16.65 ± 0.02	16.52 ± 0.02	A1p82
2011/09/26	55831.2	6.7	—	—	17.09 ± 0.15	—	—	Xing Gao
2011/09/26	55831.4	6.9	16.43 ± 0.03	17.31 ± 0.02	17.02 ± 0.01	16.80 ± 0.02	16.62 ± 0.02	A1p82
2011/09/27	55831.5	7.0	—	17.40 ± 0.02	17.19 ± 0.01	16.78 ± 0.04	16.92 ± 0.03	LT
2011/09/29	55833.5	9.0	—	17.85 ± 0.02	17.41 ± 0.03	17.12 ± 0.04	16.84 ± 0.03	LT
2011/09/29	55834.5	10.0	17.86 ± 0.03	17.93 ± 0.02	17.57 ± 0.01	17.28 ± 0.02	17.00 ± 0.02	A1p82
2011/09/29	55834.6	10.1	—	17.91 ± 0.02	17.55 ± 0.02	17.11 ± 0.02	16.87 ± 0.03	LT
2011/09/30	55835.3	10.8	—	—	17.40 ± 0.14	—	—	Xing Gao
2011/09/30	55835.4	10.9	—	17.97 ± 0.02	17.55 ± 0.02	17.24 ± 0.01	16.98 ± 0.05	A1p82
2011/10/01	55835.5	11.0	—	17.95 ± 0.01	17.55 ± 0.03	17.08 ± 0.03	17.02 ± 0.03	LT
2011/10/01	55836.4	11.9	17.74 ± 0.06	17.96 ± 0.02	17.51 ± 0.01	17.19 ± 0.01	17.02 ± 0.03	A1p82
2011/10/03	55837.5	13.0	—	17.91 ± 0.02	17.46 ± 0.03	17.07 ± 0.43	16.83 ± 0.05	LT
2011/10/03	55838.5	14.0	17.96 ± 0.06	17.84 ± 0.02	17.37 ± 0.01	17.06 ± 0.01	16.90 ± 0.03	A1p82
2011/10/05	55840.2	15.7	—	—	17.34 ± 0.18	—	—	Xing Gao
2011/10/06	55840.5	16.0	—	17.79 ± 0.02	17.26 ± 0.02	16.91 ± 0.03	16.64 ± 0.04	LT
2011/10/06	55841.4	16.9	17.45 ± 0.04	17.64 ± 0.03	17.17 ± 0.02	16.78 ± 0.02	16.71 ± 0.06	Ca2p2
2011/10/07	55841.5	17.0	—	17.60 ± 0.03	17.24 ± 0.05	16.77 ± 0.06	16.54 ± 0.05	LT
2011/10/08	55842.5	18.0	—	—	17.08 ± 0.02	—	—	LT
2011/10/08	55843.3	18.8	—	—	16.96 ± 0.15	—	—	Xing Gao
2011/10/10	55844.5	20.0	—	17.58 ± 0.02	16.99 ± 0.03	16.64 ± 0.03	16.46 ± 0.03	LT
2011/10/10	55844.5	20.0	—	—	16.90 ± 0.06	—	16.66 ± 0.05	TED
2011/10/11	55845.5	21.0	—	17.60 ± 0.04	16.97 ± 0.03	16.61 ± 0.07	16.40 ± 0.05	LT
2011/10/11	55845.5	21.0	—	—	16.94 ± 0.04	16.62 ± 0.03	—	OAC
2011/10/11	55845.5	21.0	—	—	16.94 ± 0.02	—	16.56 ± 0.03	TED
2011/10/12	55846.5	22.0	—	17.47 ± 0.03	16.92 ± 0.04	16.57 ± 0.03	16.39 ± 0.03	LT
2011/10/12	55846.5	22.0	—	—	—	16.58 ± 0.05	16.50 ± 0.06	OAC
2011/10/13	55847.5	23.0	—	17.53 ± 0.04	16.90 ± 0.04	16.54 ± 0.04	16.45 ± 0.03	LT
2011/10/14	55848.5	24.0	—	17.58 ± 0.07	16.92 ± 0.06	16.52 ± 0.06	16.36 ± 0.04	LT
2011/10/15	55849.5	25.0	—	17.73 ± 0.03	16.96 ± 0.03	16.54 ± 0.06	16.37 ± 0.04	LT
2011/10/16	55851.4	26.9	—	—	17.07 ± 0.07	—	16.56 ± 0.04	TED
2011/10/16	55851.5	27.0	18.12 ± 0.10	17.87 ± 0.03	17.08 ± 0.02	16.66 ± 0.02	16.56 ± 0.04	Ca2p2
2011/10/18	55852.5	28.0	—	18.05 ± 0.02	17.14 ± 0.02	16.63 ± 0.03	16.44 ± 0.03	LT
2011/10/18	55853.2	28.7	—	—	17.30 ± 0.13	—	—	Xing Gao
2011/10/18	55853.4	28.9	—	—	17.24 ± 0.07	—	16.53 ± 0.04	TED
2011/10/22	55857.2	32.7	—	—	17.43 ± 0.12	—	—	Xing Gao
2011/10/23	55857.6	33.1	—	18.72 ± 0.03	17.57 ± 0.03	16.83 ± 0.06	16.59 ± 0.04	LT
2011/10/23	55858.2	33.7	—	—	17.41 ± 0.14	—	—	Xing Gao
2011/10/24	55858.6	34.1	—	18.70 ± 0.02	17.61 ± 0.01	16.98 ± 0.01	16.73 ± 0.03	A1p82
2011/10/24	55859.2	34.7	—	—	17.47 ± 0.13	—	—	Xing Gao
2011/10/25	55859.5	35.0	—	18.78 ± 0.04	17.63 ± 0.03	16.92 ± 0.06	16.70 ± 0.04	LT
2011/10/25	55860.4	35.9	—	—	17.79 ± 0.05	—	16.89 ± 0.04	TED
2011/10/26	55860.5	36.0	—	18.81 ± 0.04	17.68 ± 0.03	16.97 ± 0.06	16.67 ± 0.04	LT
2011/10/27	55861.5	37.0	—	18.91 ± 0.04	17.79 ± 0.03	16.99 ± 0.06	17.08 ± 0.04	LT
2011/10/28	55862.5	38.0	19.70 ± 0.05	19.10 ± 0.03	17.77 ± 0.02	17.13 ± 0.02	16.64 ± 0.01	NOT
2011/10/28	55862.5	38.0	—	19.08 ± 0.04	17.80 ± 0.03	17.21 ± 0.02	16.76 ± 0.02	LT
2011/10/29	55863.5	39.0	—	19.06 ± 0.04	17.83 ± 0.03	17.14 ± 0.06	16.74 ± 0.04	LT

Table B1 – *continued*

Date	JD (+2400000.00)	Phase ^a (d)	<i>U</i> (mag)	<i>B</i> (mag)	<i>V</i> (mag)	<i>R</i> (mag)	<i>I</i> (mag)	Key
2011/10/30	55864.5	40.0	–	19.07 ± 0.07	17.80 ± 0.05	17.26 ± 0.06	–	LT
2011/10/30	55864.5	40.0	–	18.97 ± 0.06	17.89 ± 0.03	17.20 ± 0.03	17.01 ± 0.03	A1p82
2011/10/31	55866.5	42.0	–	19.10 ± 0.06	17.97 ± 0.03	17.31 ± 0.02	17.03 ± 0.03	A1p82
2011/11/01	55866.6	42.1	–	19.17 ± 0.06	17.99 ± 0.05	17.20 ± 0.03	16.90 ± 0.02	LT
2011/11/01	55867.4	42.9	20.02 ± 0.10	19.06 ± 0.02	18.03 ± 0.02	17.23 ± 0.02	17.12 ± 0.05	Ca2p2
2011/11/05	55871.1	46.6	–	–	17.88 ± 0.17	–	–	Xing Gao
2011/11/06	55872.1	47.6	–	–	18.00 ± 0.15	–	–	Xing Gao
2011/11/07	55872.6	48.1	–	–	–	17.57 ± 0.18	17.15 ± 0.03	WHT
2011/11/11	55876.9	52.4	–	–	–	17.59 ± 0.03	17.30 ± 0.03	FTN1
2011/11/14	55880.2	55.7	–	19.41 ± 0.10	18.27 ± 0.10	17.42 ± 0.09	17.32 ± 0.09	ASch
2011/11/15	55881.4	56.9	–	19.29 ± 0.04	18.25 ± 0.03	17.65 ± 0.02	17.42 ± 0.02	Ca2p2
2011/11/17	55882.9	48.4	–	19.51 ± 0.06	18.28 ± 0.05	17.78 ± 0.03	17.43 ± 0.03	FTN1
2011/11/19	55885.4	60.9	–	19.49 ± 0.04	18.49 ± 0.02	17.82 ± 0.02	17.45 ± 0.03	A1p82
2011/11/21	55886.5	62.0	–	19.64 ± 0.05	18.56 ± 0.03	17.76 ± 0.04	17.39 ± 0.02	A1p82
2011/11/22	55888.1	63.6	–	–	18.38 ± 0.13	–	–	Xing Gao
2011/11/23	55888.5	64.0	–	19.47 ± 0.05	18.40 ± 0.06	17.67 ± 0.04	17.42 ± 0.05	LT
2011/11/25	55890.8	66.3	–	19.66 ± 0.06	18.45 ± 0.05	17.80 ± 0.03	17.57 ± 0.03	FTN1
2011/11/27	55893.0	68.5	–	–	18.41 ± 0.11	17.78 ± 0.13	17.59 ± 0.13	FTN1
2011/11/29	55894.9	70.4	–	19.52 ± 0.07	18.50 ± 0.05	17.94 ± 0.04	17.66 ± 0.03	FTN1
2011/11/29	55894.9	70.4	–	19.68 ± 0.03	18.59 ± 0.02	17.96 ± 0.02	17.46 ± 0.01	NOT
2011/11/30	55896.4	71.9	–	19.63 ± 0.05	18.57 ± 0.06	17.89 ± 0.05	17.55 ± 0.06	LT
2011/12/01	55896.8	72.3	–	–	18.53 ± 0.05	17.97 ± 0.04	17.71 ± 0.03	FTN1
2011/12/06	55901.8	77.3	–	19.91 ± 0.06	18.56 ± 0.05	18.04 ± 0.03	17.77 ± 0.03	FTN2
2011/12/12	55908.4	83.9	–	–	18.87 ± 0.07	18.09 ± 0.05	17.73 ± 0.01	LT
2011/12/20	55916.4	91.9	–	19.71 ± 0.09	18.94 ± 0.09	18.34 ± 0.04	17.97 ± 0.05	A1p82
2011/12/22	55917.8	93.3	20.01 ± 0.07	19.80 ± 0.03	18.88 ± 0.03	18.38 ± 0.03	17.96 ± 0.06	NOT
2011/12/22	55917.8	93.3	–	19.75 ± 0.06	18.85 ± 0.05	18.40 ± 0.03	18.11 ± 0.03	FTN2
2012/01/01	55928.3	103.8	–	19.80 ± 0.02	19.07 ± 0.03	18.52 ± 0.03	18.34 ± 0.05	Ca2p2
2012/01/09	55936.5	112.0	–	19.96 ± 0.08	19.18 ± 0.07	18.72 ± 0.05	18.44 ± 0.05	LT
2012/01/12	55939.4	114.9	–	19.92 ± 0.03	19.33 ± 0.03	18.85 ± 0.02	18.67 ± 0.05	Ca2p2
2012/01/15	55942.3	117.8	–	20.00 ± 0.08	19.24 ± 0.07	18.76 ± 0.06	18.41 ± 0.06	LT
2012/01/27	55954.4	129.9	–	20.32 ± 0.08	19.48 ± 0.07	19.04 ± 0.06	18.80 ± 0.04	LT
2012/02/17	55975.3	150.8	–	20.55 ± 0.13	19.98 ± 0.09	19.60 ± 0.06	19.19 ± 0.05	A1p82
2012/02/22	55980.3	155.8	–	20.69 ± 0.18	19.99 ± 0.15	19.61 ± 0.04	19.01 ± 0.04	A1p82
2012/02/23	55981.4	156.9	–	–	20.13 ± 0.08	19.73 ± 0.03	19.27 ± 0.08	LT
2012/02/25	55983.4	148.9	–	–	19.89 ± 0.25	19.48 ± 0.11	19.24 ± 0.13	LT
2012/07/20	56128.6	304.1	–	–	–	21.54 ± 0.02	21.45 ± 0.04	GTC

^aPhase in days with respect to the adopted explosion date JD = 2455824.5 ± 0.7

Table B2. Optical Sloan photometry of SN 2011fu in the Vega system.

Date	JD (+2400000.00)	Phase ^a (d)	<i>u</i> (mag)	<i>g</i> (mag)	<i>z</i> (mag)	Key
2011/09/27	55831.5	7.0	16.43 ± 0.06	–	16.37 ± 0.03	LT
2011/09/29	55833.5	9.0	17.02 ± 0.04	–	16.61 ± 0.04	LT
2011/09/30	55834.6	10.1	17.37 ± 0.06	–	16.77 ± 0.03	LT
2011/10/01	55835.5	11.0	17.45 ± 0.09	–	16.70 ± 0.03	LT
2011/10/03	55837.5	13.0	–	–	16.66 ± 0.05	LT
2011/10/06	55840.5	16.0	17.38 ± 0.04	–	16.49 ± 0.04	LT
2011/10/08	55842.5	18.0	–	–	16.40 ± 0.06	LT
2011/10/10	55844.5	20.0	17.23 ± 0.03	–	16.32 ± 0.03	LT
2011/10/11	55845.5	21.0	17.5 ± 0.13	–	16.28 ± 0.05	LT
2011/10/12	55846.5	22.0	17.71 ± 0.11	–	16.30 ± 0.03	LT
2011/10/13	55847.5	25.0	17.75 ± 0.03	–	16.34 ± 0.05	LT
2011/10/14	55848.5	24.0	17.74 ± 0.11	–	16.36 ± 0.05	LT
2011/10/15	55849.5	25.0	17.82 ± 0.10	–	16.22 ± 0.05	LT
2011/10/18	55852.5	28.0	18.22 ± 0.10	–	16.32 ± 0.03	LT
2011/10/23	55857.6	33.1	–	–	16.46 ± 0.05	LT
2011/10/25	55859.5	35.0	–	–	16.54 ± 0.05	LT
2011/10/26	55860.5	36.0	–	–	16.52 ± 0.05	LT
2011/10/27	55861.5	37.0	–	–	16.59 ± 0.05	LT
2011/10/28	55862.5	38.0	19.37 ± 0.03	–	16.66 ± 0.02	LT
2011/10/29	55863.5	39.0	19.24 ± 0.07	–	16.65 ± 0.04	LT
2011/10/30	55864.5	40.0	–	–	16.75 ± 0.04	LT
2011/11/01	55866.6	42.1	–	–	16.69 ± 0.02	LT
2011/11/07	55872.6	48.1	–	18.72 ± 0.29	16.85 ± 0.05	WHT
2011/11/23	55888.5	64.0	–	–	17.10 ± 0.06	LT
2011/11/30	55896.4	71.9	–	–	17.35 ± 0.05	LT
2011/12/12	55908.4	83.9	–	–	17.51 ± 0.05	LT
2012/01/09	55936.5	112.0	–	–	18.07 ± 0.07	LT
2012/01/15	55942.3	117.8	–	–	18.14 ± 0.08	LT
2012/01/27	55954.4	129.9	–	–	18.25 ± 0.07	LT
2012/02/23	55981.4	156.9	–	–	18.82 ± 0.16	LT
2012/07/20	56128.6	304.1	–	22.41 ± 0.15	–	GTC

^aPhase in days with respect to the adopted explosion date JD = 2455824.5 ± 0.7**Table B3.** NIR photometry of SN 2011fu.

Date	JD (+2400000.00)	Phase ^a (d)	<i>J</i> (mag)	<i>H</i> (mag)	<i>K_s</i> (mag)	Key
2011/10/06	55840.6	16.1	16.20 ± 0.12	16.10 ± 0.14	15.75 ± 0.14	TCS
2011/10/10	55844.6	20.1	16.08 ± 0.12	–	15.72 ± 0.27	TCS
2011/10/12	55847.4	23.0	16.15 ± 0.08	15.70 ± 0.15	15.63 ± 0.34	TCS
2011/10/29	55864.5	40.0	16.28 ± 0.13	–	–	TCS
2011/10/31	55866.5	42.0	16.54 ± 0.15	16.06 ± 0.15	15.77 ± 0.21	TCS
2011/11/04	55869.5	45.0	16.74 ± 0.15	16.27 ± 0.12	16.17 ± 0.27	TCS

^aPhase in days with respect to the adopted explosion date JD = 2455824.5 ± 0.7

Table B4. Johnson-Cousins optical and 2MASS NIR magnitudes and associated errors for the stellar sequence used in the calibration process of SN 2011fu’s photometry.

Star	<i>U</i> (mag)	<i>B</i> (mag)	<i>V</i> (mag)	<i>R</i> (mag)	<i>I</i> (mag)	<i>J</i> (mag)	<i>H</i> (mag)	<i>K_s</i> (mag)
1	19.26 ± 0.10	18.66 ± 0.02	17.74 ± 0.02	17.20 ± 0.03	16.65 ± 0.03	–	–	–
2	18.25 ± 0.10	18.17 ± 0.02	17.52 ± 0.02	17.10 ± 0.05	16.79 ± 0.03	–	–	–
3	18.26 ± 0.06	18.25 ± 0.02	17.55 ± 0.02	17.24 ± 0.02	16.87 ± 0.03	16.24 ± 0.12	15.96 ± 0.19	15.79 ± 0.26
4	18.98 ± 0.09	18.14 ± 0.02	17.18 ± 0.02	16.63 ± 0.02	16.14 ± 0.03	15.45 ± 0.06	14.81 ± 0.08	14.81 ± 0.11
5	17.75 ± 0.04	17.83 ± 0.02	17.25 ± 0.02	16.90 ± 0.02	16.58 ± 0.02	15.91 ± 0.08	15.77 ± 0.15	15.51 ± 0.20
6	15.68 ± 0.09	15.51 ± 0.02	14.88 ± 0.07	14.47 ± 0.07	14.11 ± 0.06	13.49 ± 0.02	13.16 ± 0.03	13.07 ± 0.03
7	17.88 ± 0.04	17.77 ± 0.02	17.14 ± 0.03	16.77 ± 0.03	16.41 ± 0.02	15.83 ± 0.08	15.30 ± 0.10	15.42 ± 0.19
8	19.94 ± 0.05	18.82 ± 0.02	17.58 ± 0.03	16.82 ± 0.02	16.15 ± 0.03	15.11 ± 0.05	14.54 ± 0.05	14.28 ± 0.07
9	17.60 ± 0.04	17.61 ± 0.02	17.07 ± 0.02	16.75 ± 0.03	16.42 ± 0.02	15.97 ± 0.09	15.73 ± 0.15	15.25*
10	19.48 ± 0.10	18.89 ± 0.02	18.02 ± 0.02	17.57 ± 0.02	17.14 ± 0.02	–	–	–
11	15.21 ± 0.04	15.21 ± 0.02	14.71 ± 0.02	14.37 ± 0.02	14.03 ± 0.02	–	–	–
12	16.33 ± 0.05	15.74 ± 0.03	14.87 ± 0.02	14.42 ± 0.06	14.00 ± 0.05	13.24 ± 0.02	12.82 ± 0.03	12.74 ± 0.03
13	19.37 ± 0.07	18.77 ± 0.01	17.99 ± 0.02	17.52 ± 0.03	17.08 ± 0.03	16.35 ± 0.13	15.68 ± 0.13	15.93 ± 0.29
14	16.83 ± 0.04	16.79 ± 0.01	16.24 ± 0.03	15.92 ± 0.06	15.57 ± 0.02	14.95 ± 0.04	14.73 ± 0.06	14.79 ± 0.11
15	18.52 ± 0.04	17.84 ± 0.01	16.97 ± 0.03	16.46 ± 0.03	16.03 ± 0.03	15.26 ± 0.05	14.79 ± 0.07	14.70 ± 0.10

* No value for the uncertainty is given by the 2-MASS catalogue.

Table B5. Sloan Vega magnitudes and associated errors for the stellar sequence used in the calibration process of SN 2011fu’s photometry.

Star	<i>u</i> (mag)	<i>g*</i> (mag)	<i>r</i> (mag)	<i>i</i> (mag)	<i>z</i> (mag)
1	–	–	17.20 ± 0.01	16.64 ± 0.01	16.28 ± 0.01
2	–	17.38 ± 0.01	17.15 ± 0.01	16.77 ± 0.01	16.52 ± 0.03
3	–	17.44 ± 0.01	17.22 ± 0.01	16.83 ± 0.01	16.59 ± 0.01
4	–	17.19 ± 0.01	16.63 ± 0.01	16.13 ± 0.01	15.82 ± 0.01
5	17.85 ± 0.02	17.04 ± 0.01	16.89 ± 0.01	16.54 ± 0.01	16.34 ± 0.02
6	15.52 ± 0.03	–	14.45 ± 0.01	14.06 ± 0.01	13.83 ± 0.01
7	17.59 ± 0.05	16.96 ± 0.01	16.75 ± 0.01	16.36 ± 0.01	16.12 ± 0.01
8	–	17.73 ± 0.01	16.83 ± 0.01	16.16 ± 0.01	15.71 ± 0.01
9	17.59 ± 0.05	16.88 ± 0.01	16.72 ± 0.01	16.38 ± 0.01	16.16 ± 0.02
10	–	17.96 ± 0.01	17.57 ± 0.01	17.12 ± 0.01	16.84 ± 0.03
11	15.05 ± 0.02	–	14.32 ± 0.02	13.98 ± 0.02	13.84 ± 0.04
12	16.15 ± 0.03	–	14.38 ± 0.04	13.97 ± 0.01	13.68 ± 0.01
13	–	17.87 ± 0.01	17.53 ± 0.01	17.05 ± 0.01	16.78 ± 0.02
14	–	–	15.88 ± 0.01	15.52 ± 0.01	15.29 ± 0.01
15	–	15.95 ± 0.01	16.46 ± 0.01	16.00 ± 0.01	15.71 ± 0.01

* The stars were only measured in one epoch in this band.